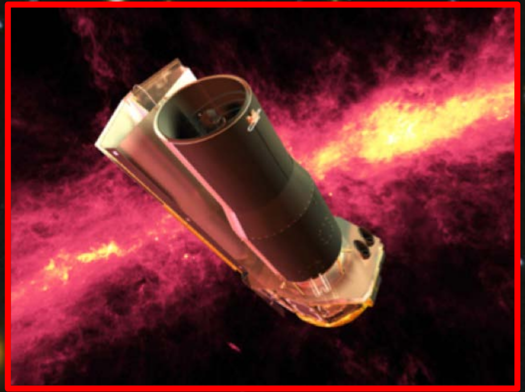
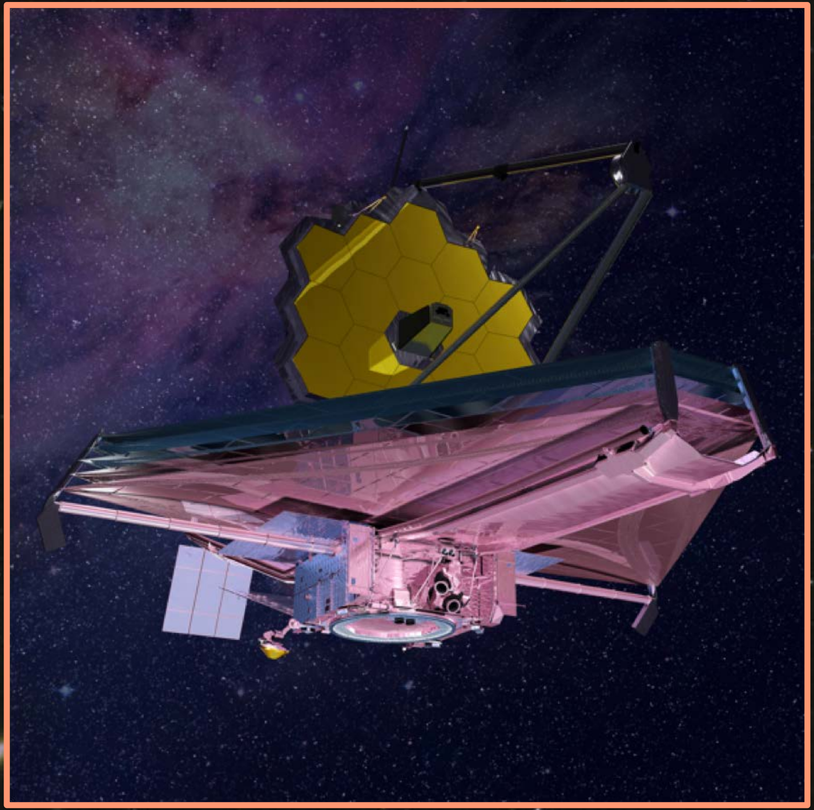
John N. Bahcall Lecture
Goddard Space Flight Center
March 13 2018





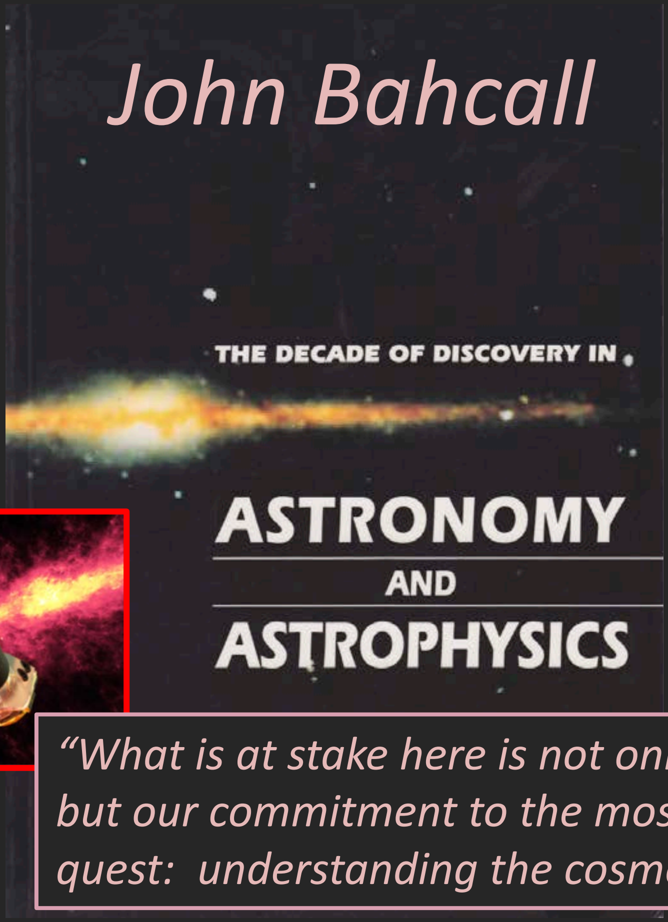
Galaxies at Cosmic Dawn: Exploring the First Billion Years with Hubble and Spitzer – Implications for JWST

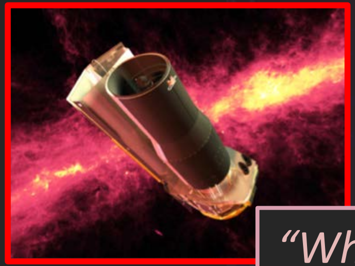
Garth Illingworth
University of California Santa Cruz

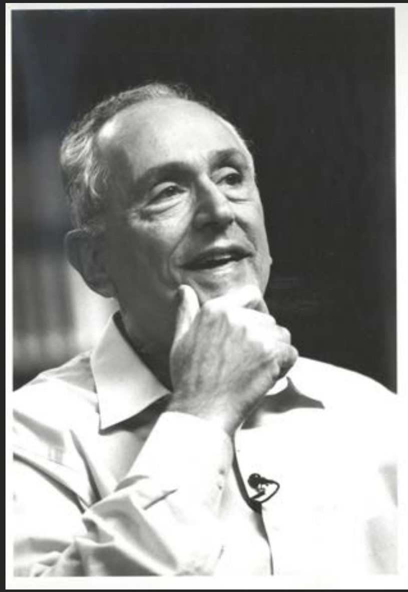
firstgalaxies.org

figure credit: Adolf Schaller

John Bahcall







"What is at stake here is not only a piece of stellar technology but our commitment to the most fundamental human quest: understanding the cosmos." – John N. Bahcall

1970s—1980s John's continuing efforts to support Hubble were crucial and inspiring (and a model for what was needed from scientists for a major mission to be successful).

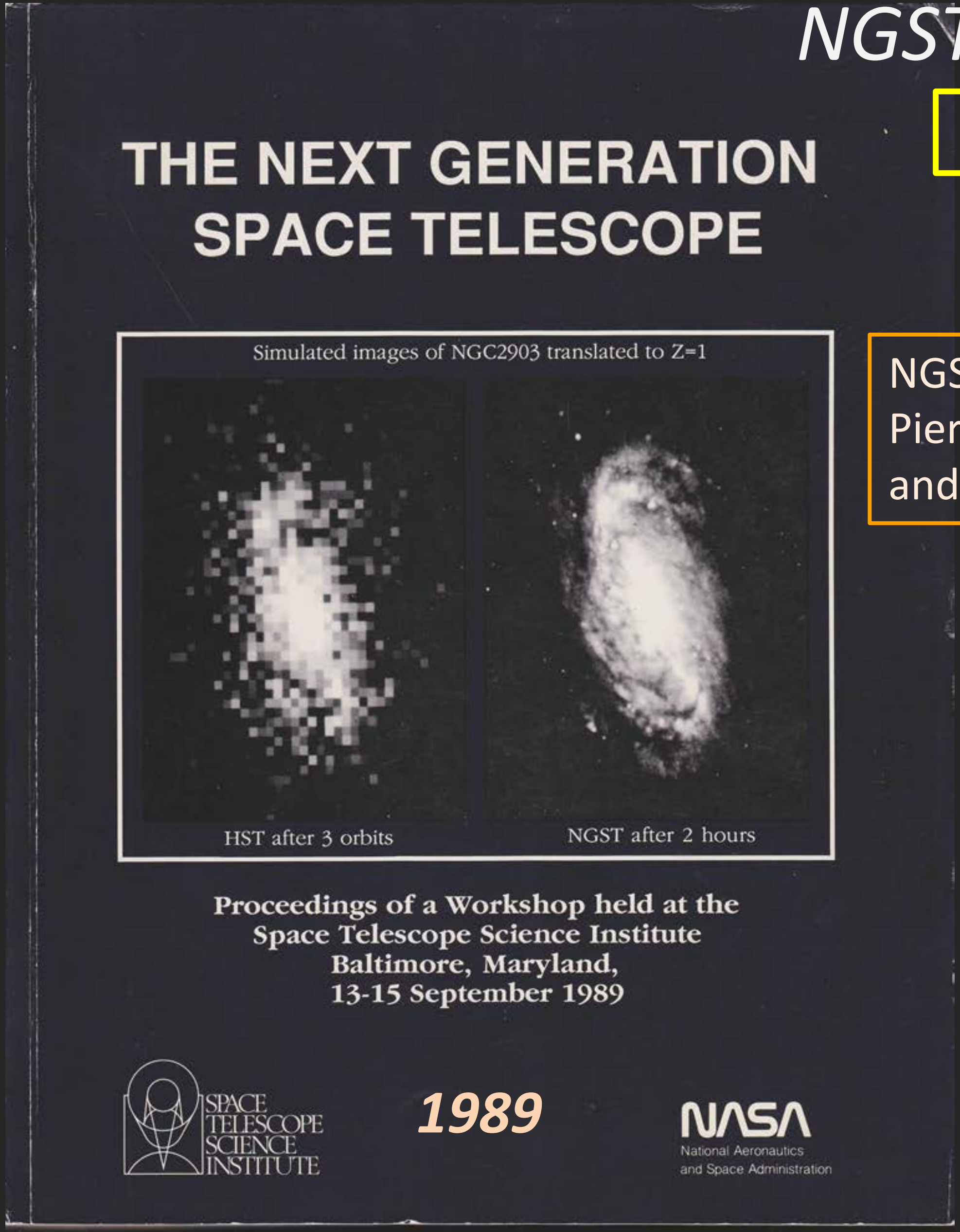
1989 – John's introductory remarks and participation in the NGST 1989 workshop
1991 – John was Chair, Astronomy Decadal Survey. I was Chair of UV-Optical in Space Panel.
2000 – John was Chair, SIRTF Legacy Science TAC. I was Panel Chair.

gdi

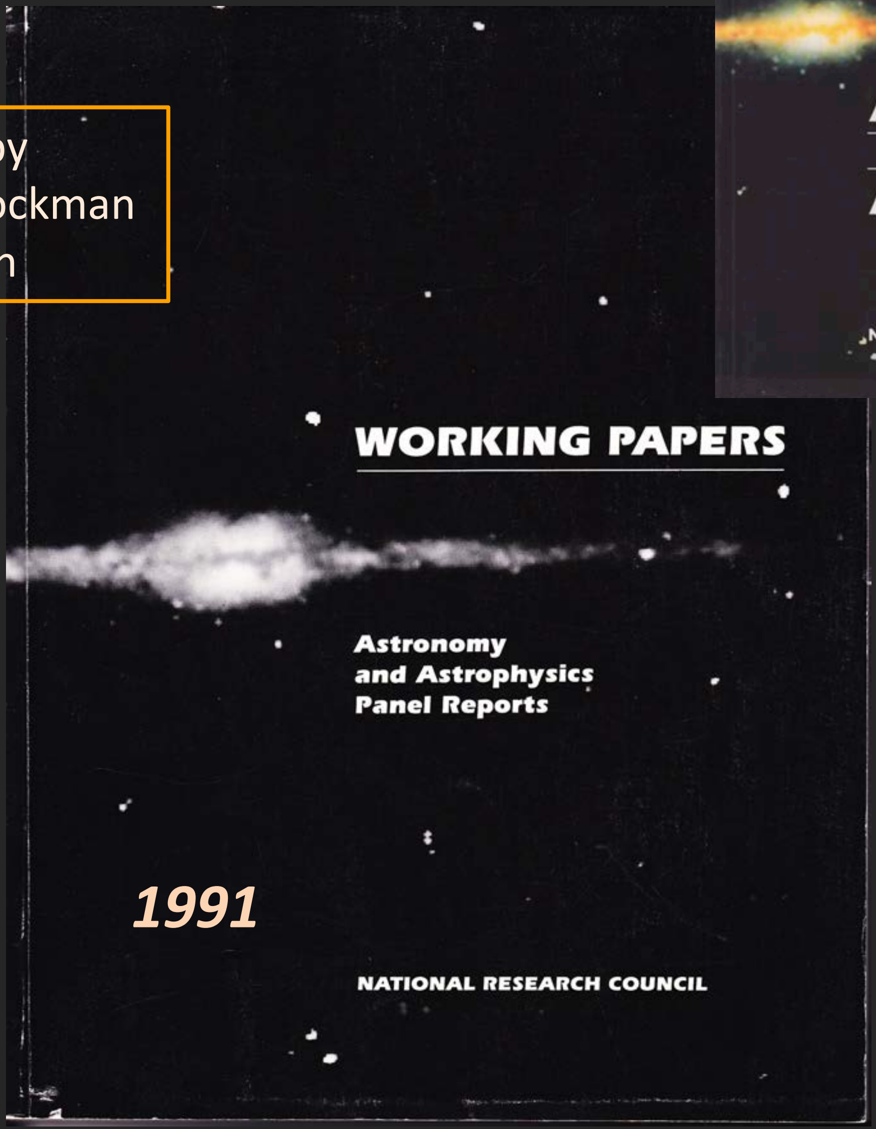
NGST (JWST) – key early events

30 years from NGST mission concept to JWST launch!

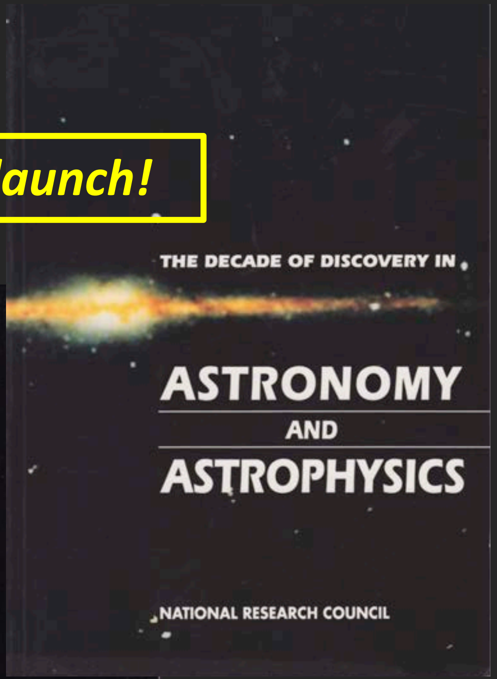
NGST in mid-1980s by Pierre Bely, Peter Stockman and Garth Illingworth



1989



1991

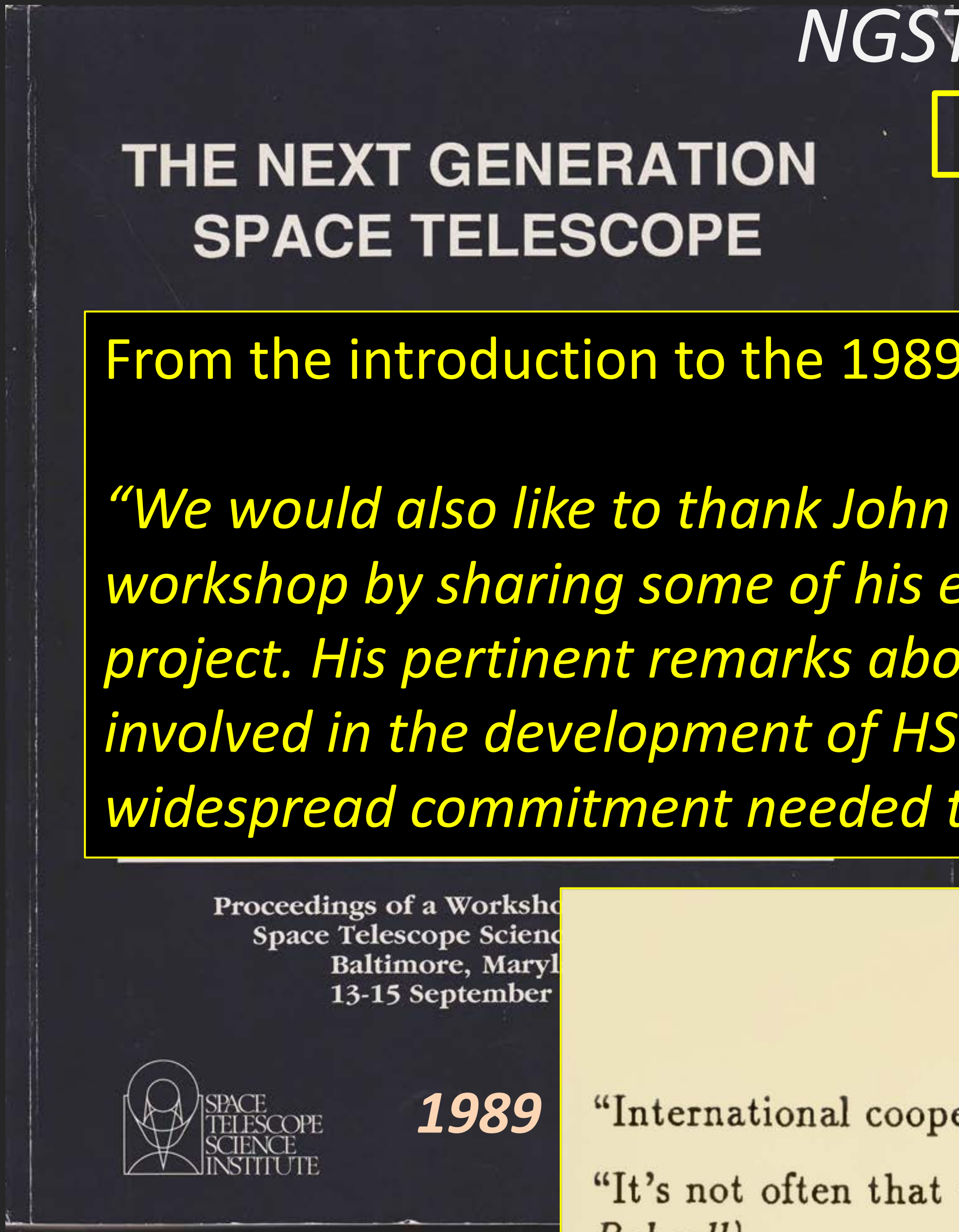


www.ucolick.org/~gdi/early_jwst/

gdi

NGST (JWST) – key early events

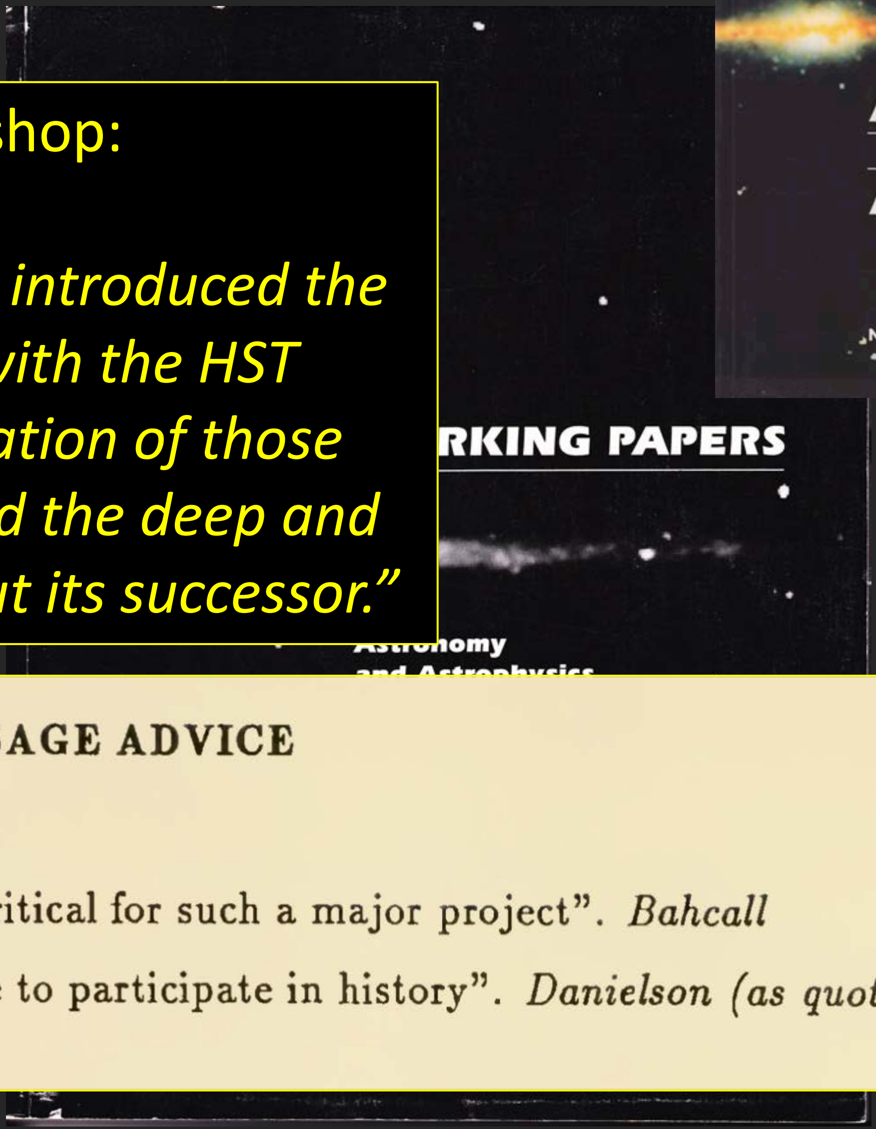
30 years from NGST mission concept to JWST launch!

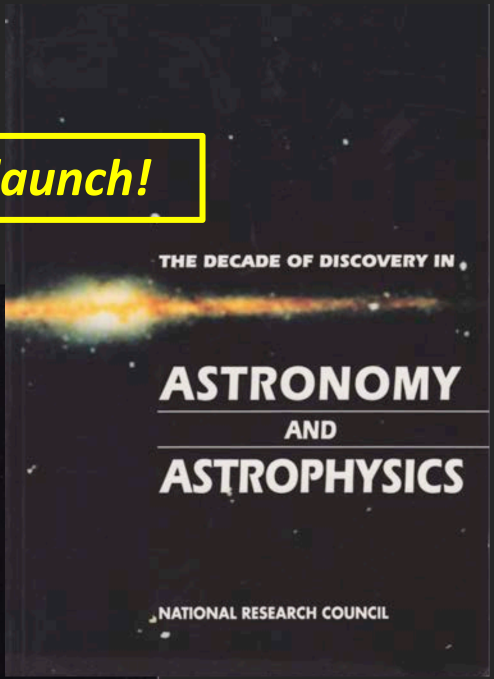


1989

From the introduction to the 1989 NGST workshop:

"We would also like to thank John Bahcall who introduced the workshop by sharing some of his experiences with the HST project. His pertinent remarks about the dedication of those involved in the development of HST emphasized the deep and widespread commitment needed to bring about its successor."





SAGE ADVICE

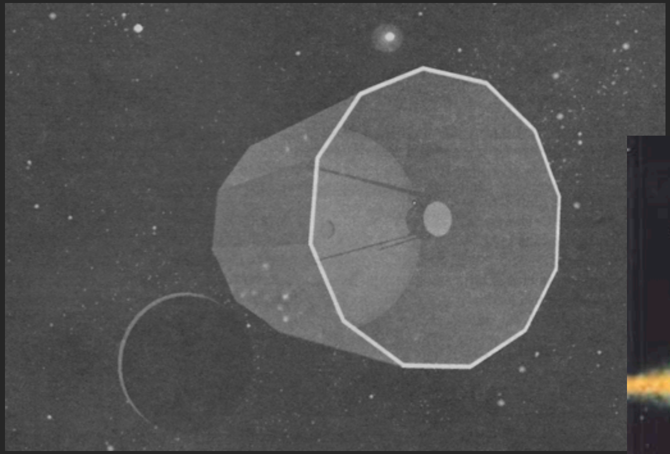
"International cooperation may be critical for such a major project". *Bahcall*

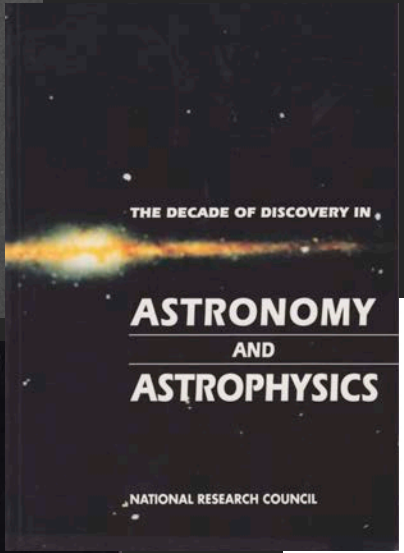
"It's not often that we have a chance to participate in history". *Danielson (as quoted by Bahcall)*

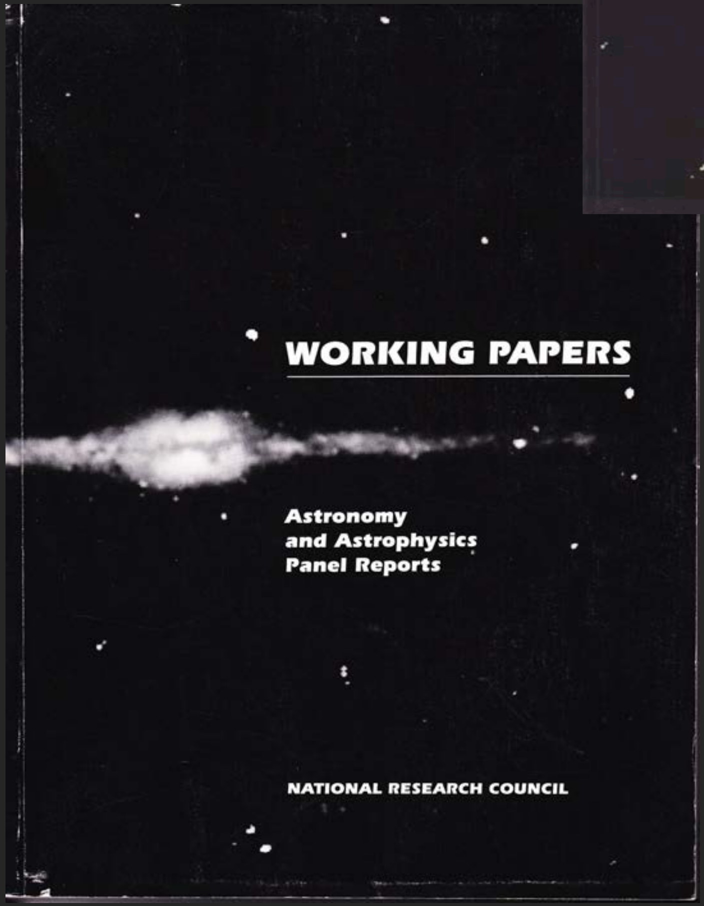
www.ucolick.org/~gdi/early_jwst/

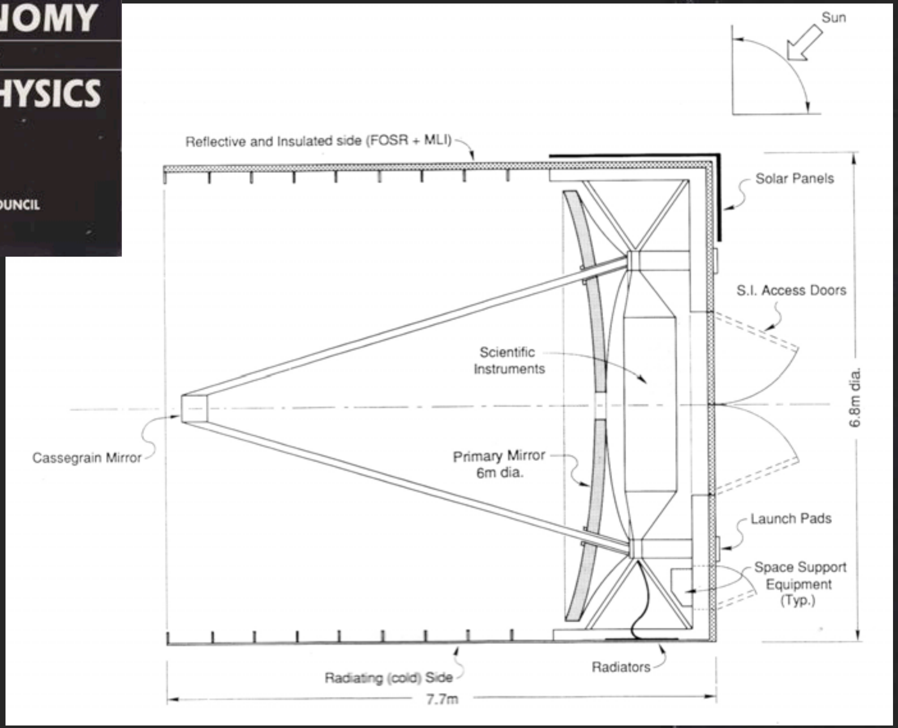
gdi

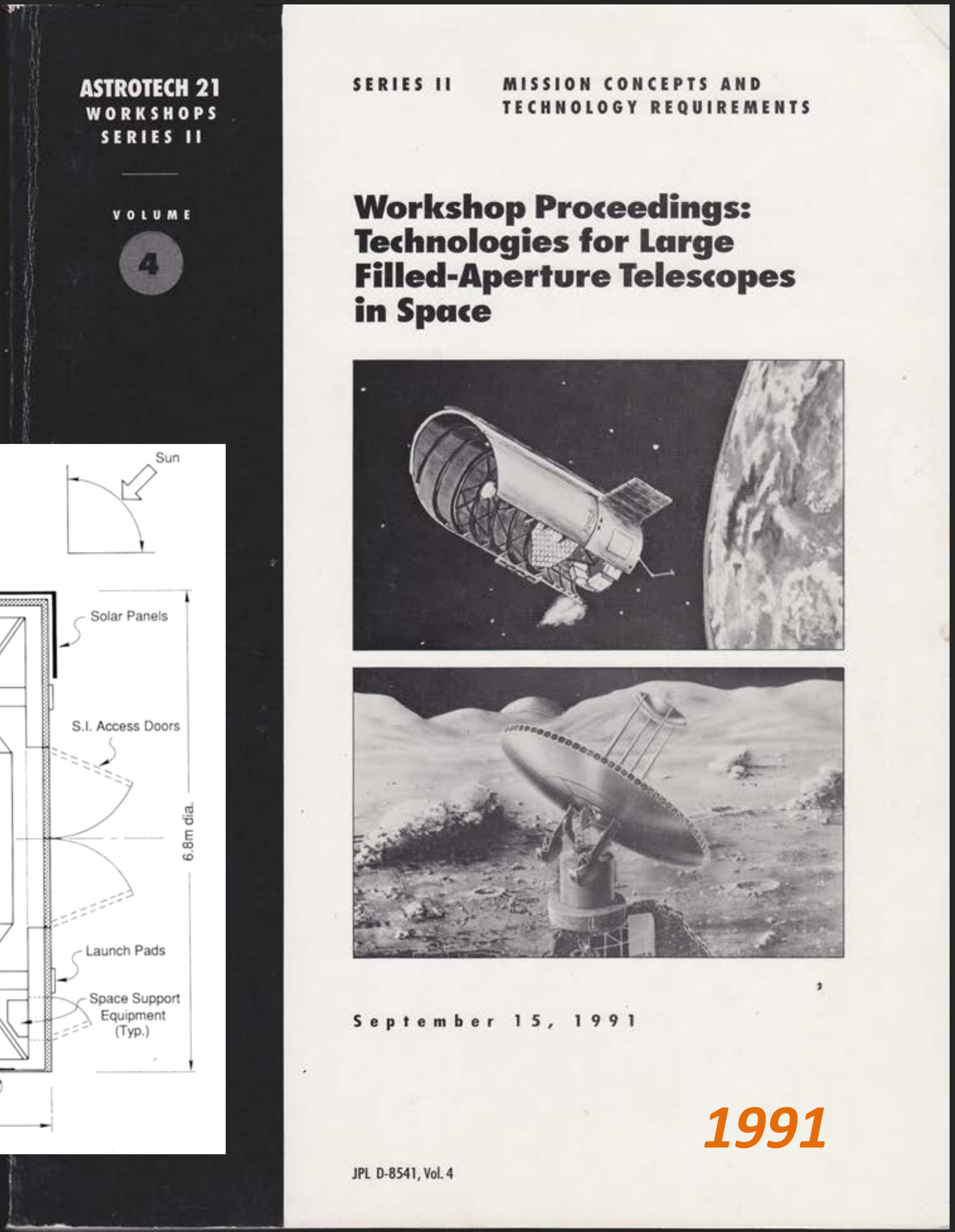
NGST & the 1990 Decadal











1991

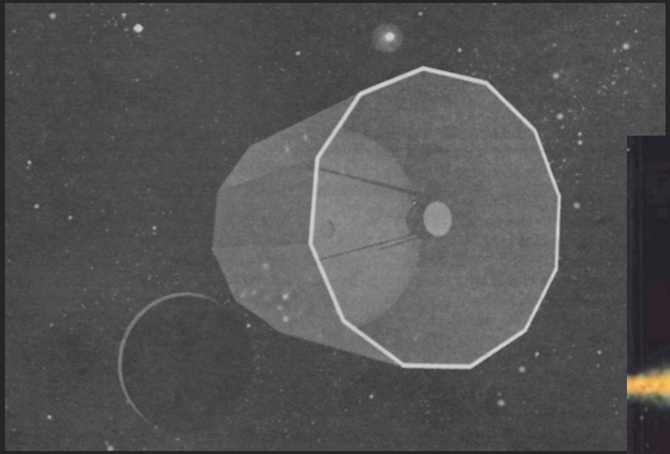
see [2016 STScI Newsletter article](newsletter.stsci.edu/early-webb-history)

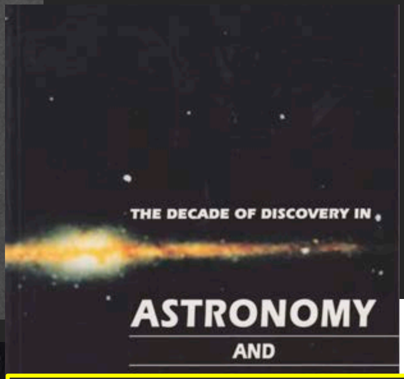
NGST: The Early Days of JWST

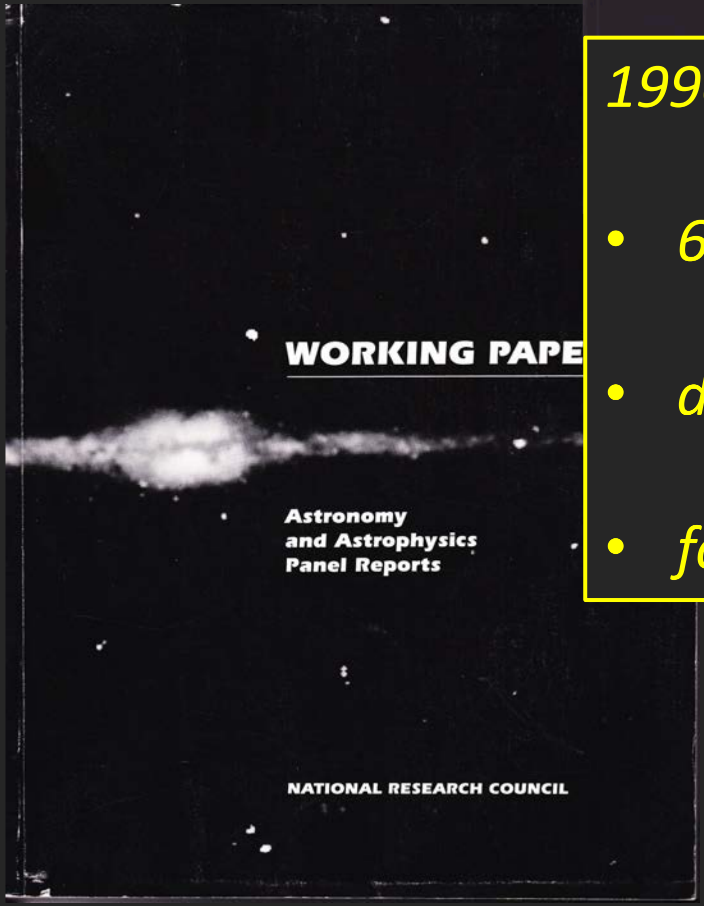
newsletter.stsci.edu/early-webb-history

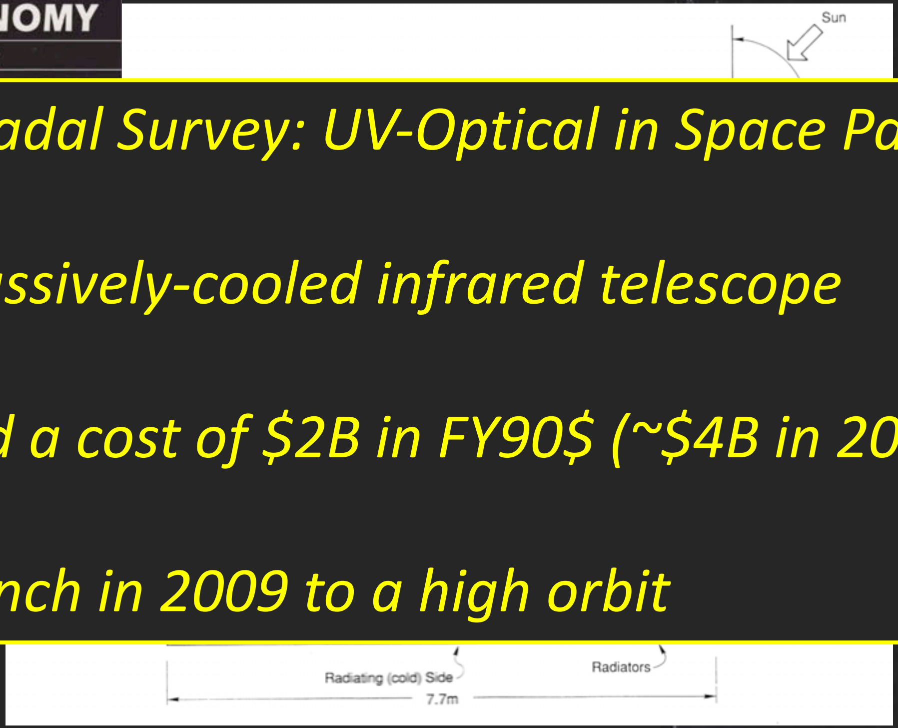
gdi

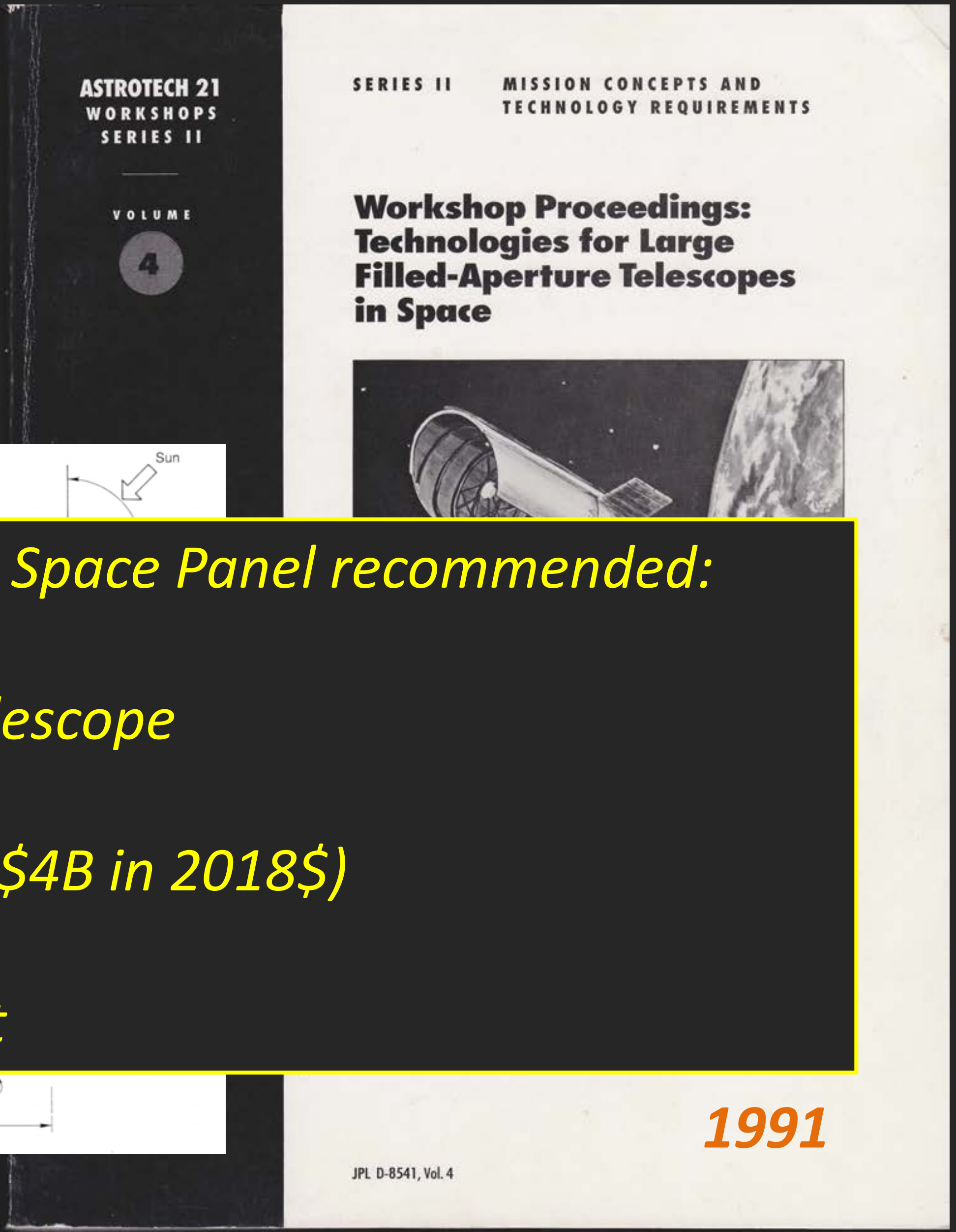
NGST & the 1990 Decadal











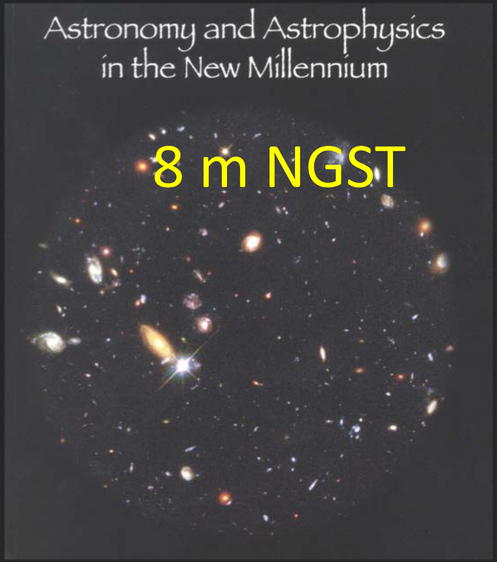
1990 Decadal Survey: UV-Optical in Space Panel recommended:

- *6-m passively-cooled infrared telescope*
- *derived a cost of $2B in FY90$ (~$4B in 2018$)*
- *for launch in 2009 to a high orbit*

see 2016 STScI Newsletter article
NGST: The Early Days of JWST
newsletter.stsci.edu/early-webb-history

gdi

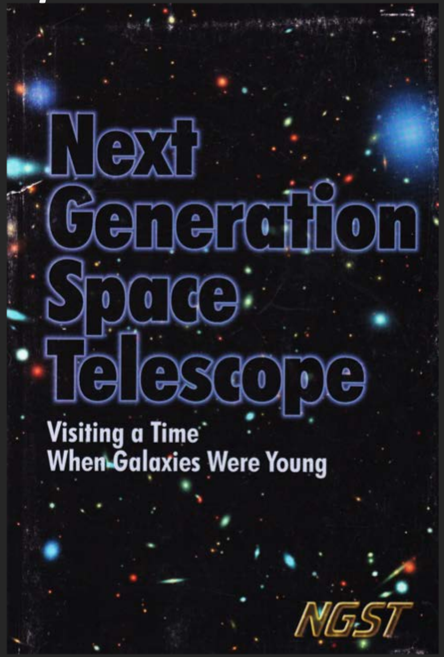
NGST ⇨ JWST – key steps in the 1990s leading to development



1996: *HST and Beyond* study (chair Dressler) with 3 recommendations including an IR telescope "….of aperture 4 m or larger, optimized for imaging and spectroscopy over …. 1-5 μm."

Crucial change: Dan Goldin: "I see Alan Dressler here. All he wants is a four meter optic that goes from a half micron to 20 microns. And I said to him, "Why do you ask for such a modest thing? Why not go after six or seven meters?""

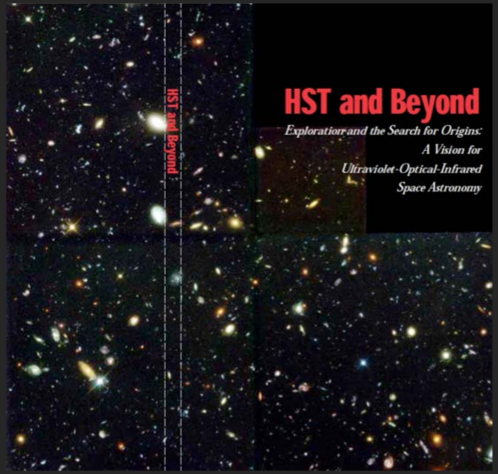
1996 American Astronomical Society meeting



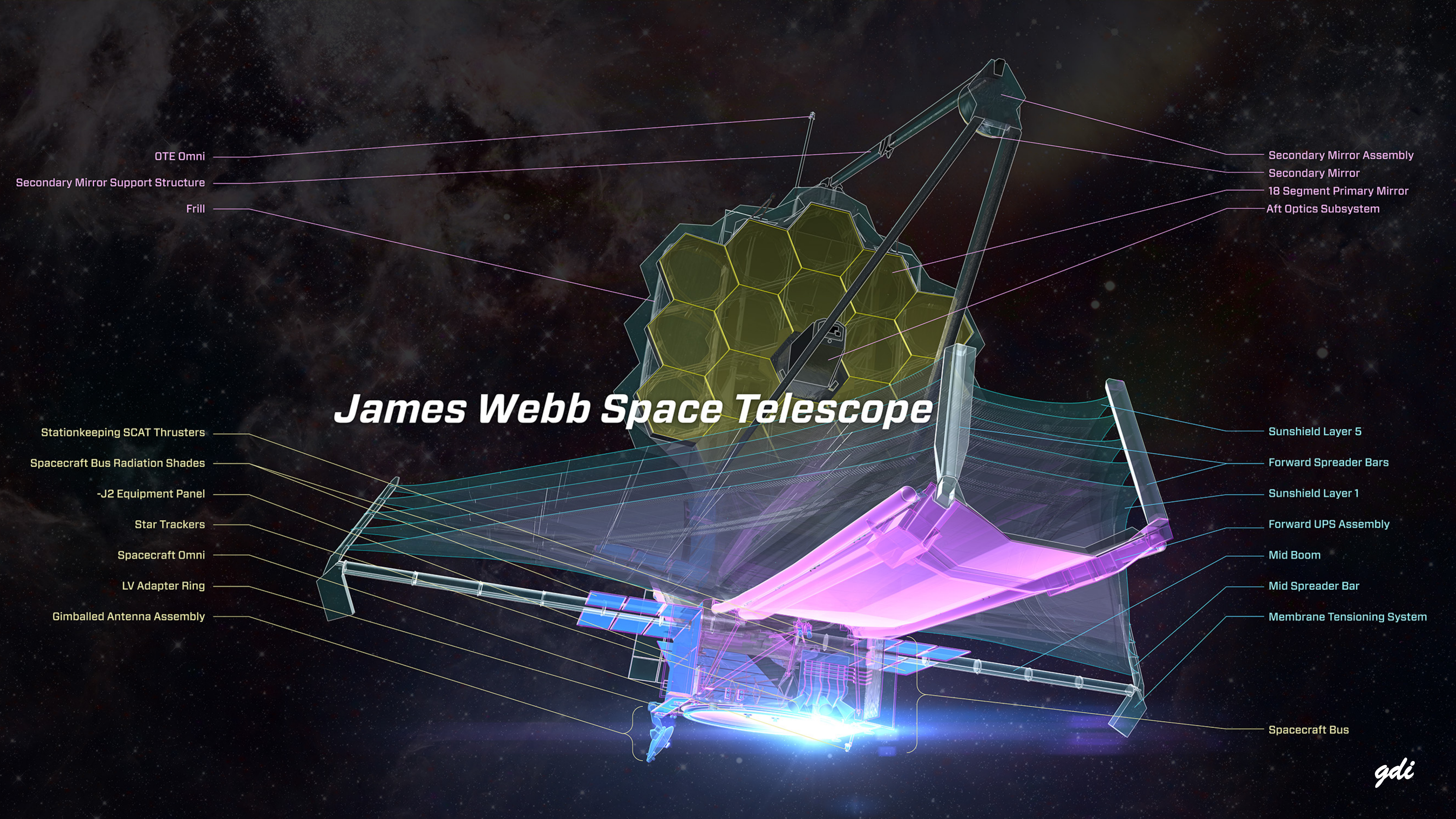
Goddard's central role started with a crucial step involving Ed Weiler and John Mather

1999: SMD AA Weiler signs Formulation Authorization – NASA starts NGST

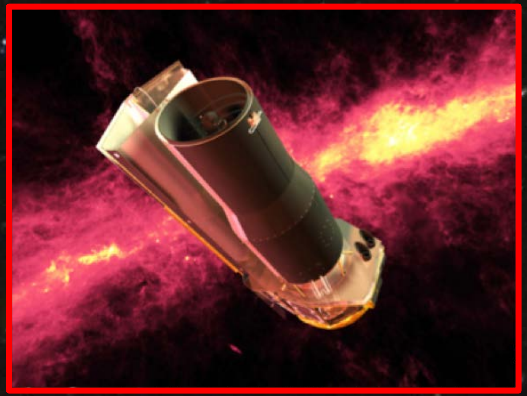
Goddard's technical excellence and project management experience has been central to the accomplishments of the JWST program – and will be to its ultimate success



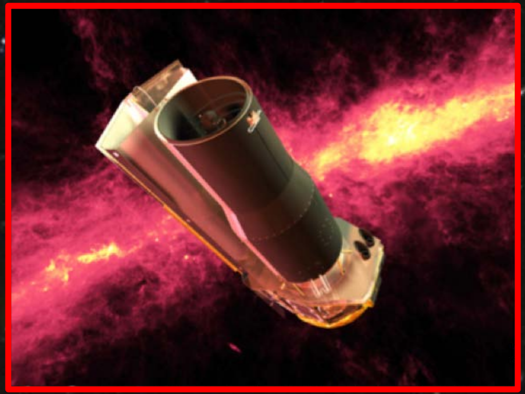
gdi



James Webb Space Telescope
OTE Omni
Secondary Mirror Support Structure
Frill
Secondary Mirror Assembly
Secondary Mirror
18 Segment Primary Mirror
Aft Optics Subsystem
Stationkeeping SCAT Thrusters
Spacecraft Bus Radiation Shades
-J2 Equipment Panel
Star Trackers
Spacecraft Omni
LV Adapter Ring
Gimballed Antenna Assembly
Sunshield Layer 5
Forward Spreader Bars
Sunshield Layer 1
Forward UPS Assembly
Mid Boom
Mid Spreader Bar
Membrane Tensioning System
Spacecraft Bus
gdi



galaxies at cosmic dawn



galaxies at cosmic dawn

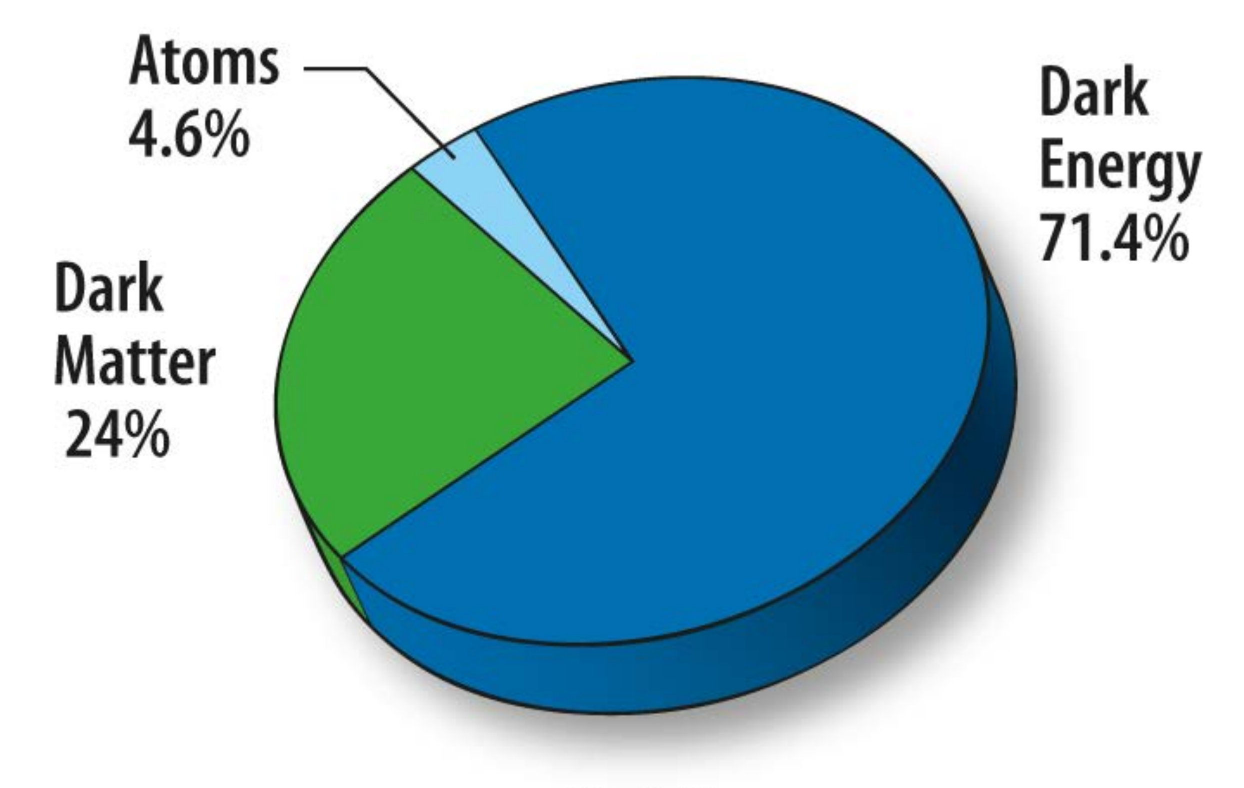
science collaborators & science team members

Rychard Bouwens, Pascal Oesch, Pieter van Dokkum, Ivo Labbé,
Marijn Franx, Mauro Stefanon, Renske Smit, Dan Magee, Holland Ford
& the HUDF09/XDF/HLF, 3D-HST and ACS GTO science teams

gdi

our strange universe

it is all dark matter & dark energy – and a little bit of ordinary matter “icing on the cake”

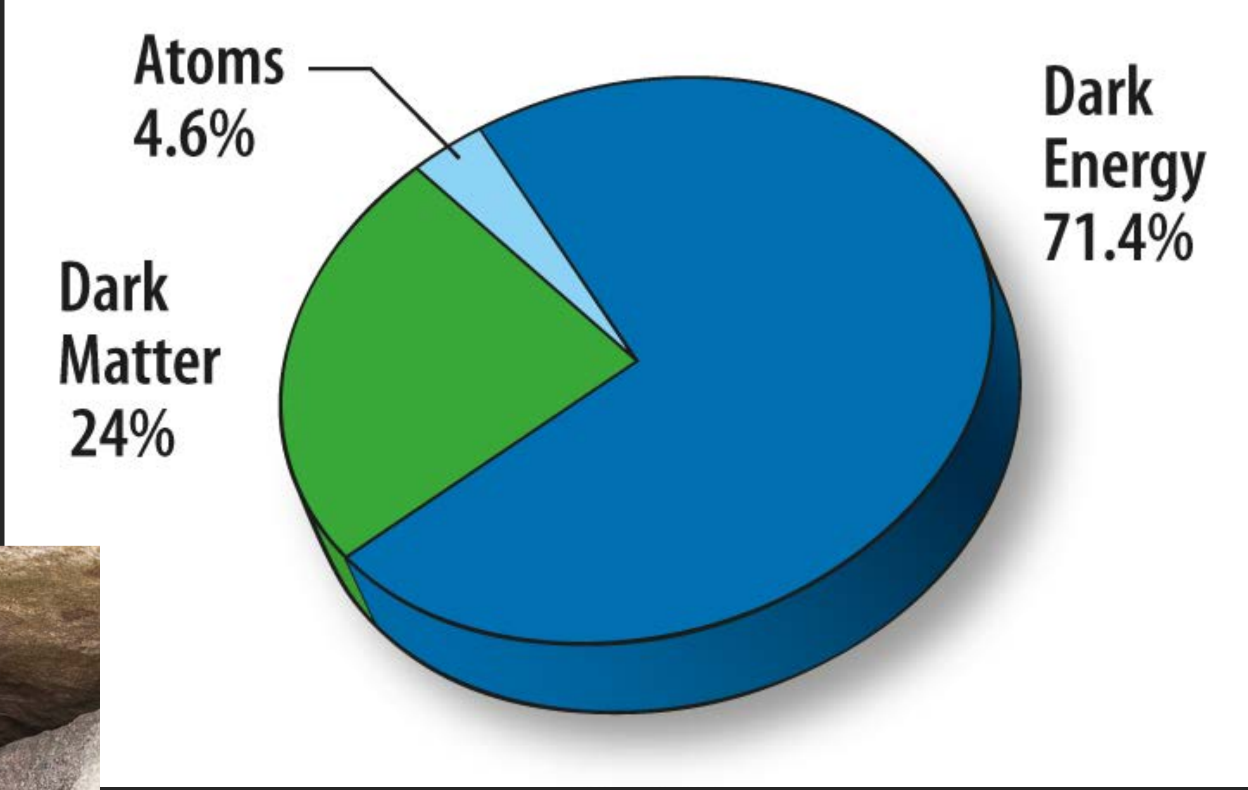


from WMAP and Planck telescopes

gdi

our strange universe

it is all dark matter & dark energy – and a little bit of ordinary matter "icing on the cake"



ordinary matter is, by comparison, a bit mousey…

from WMAP and Planck telescopes

gdi

history of everything

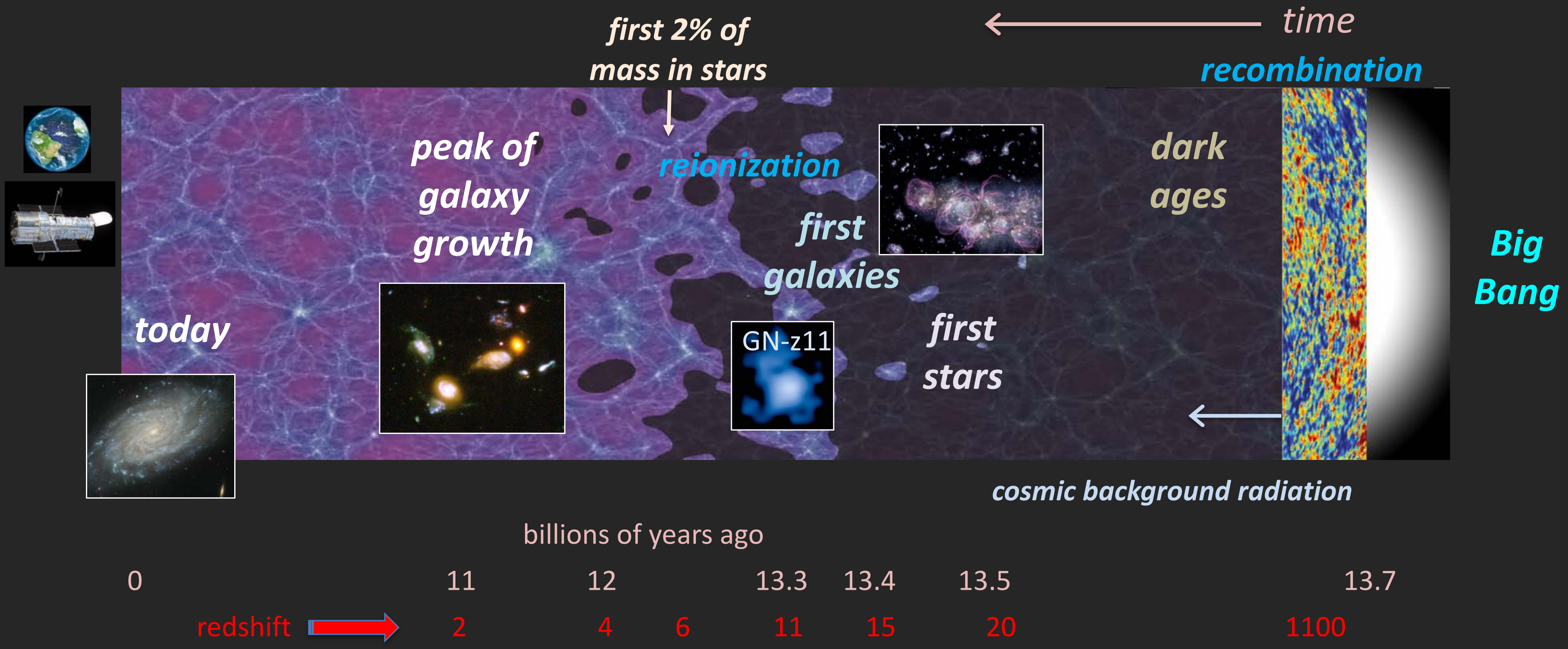


figure credit: insert adapted from Brant Robertson UCSC

history of everything

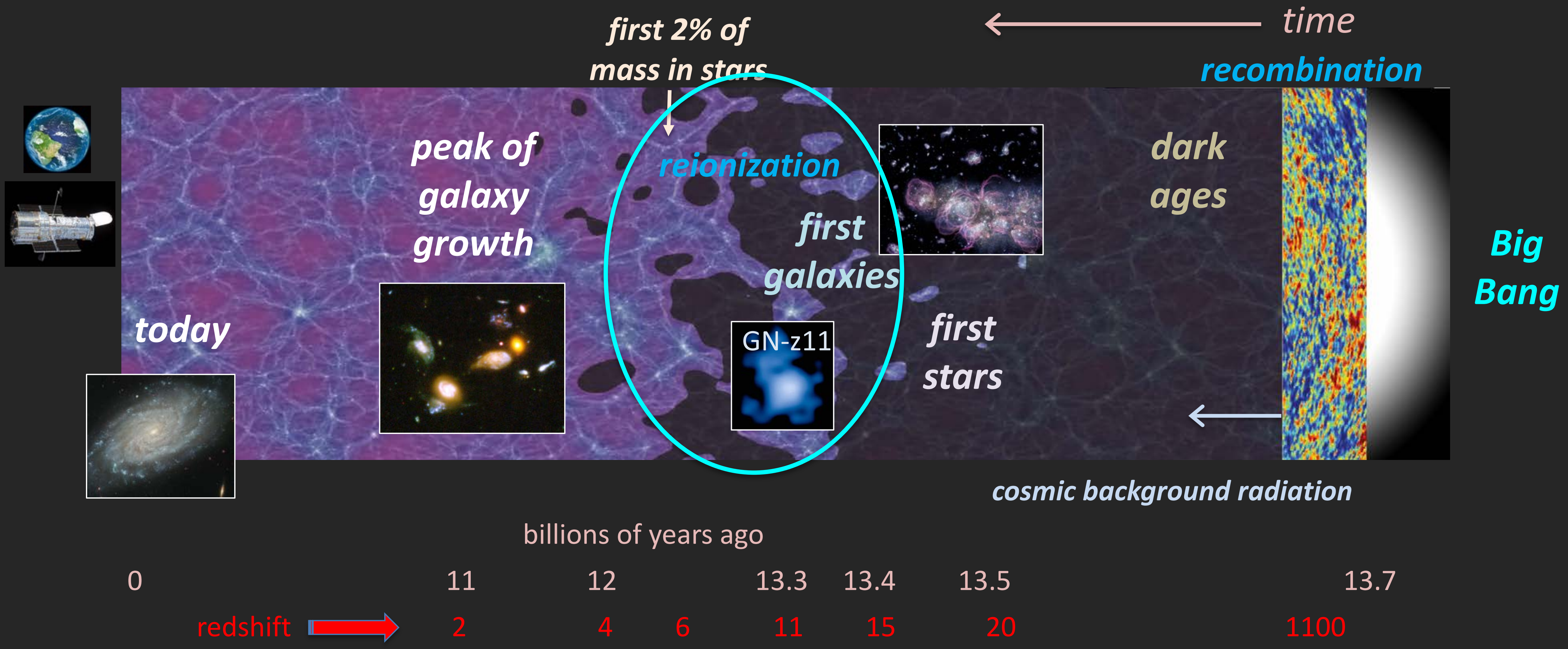
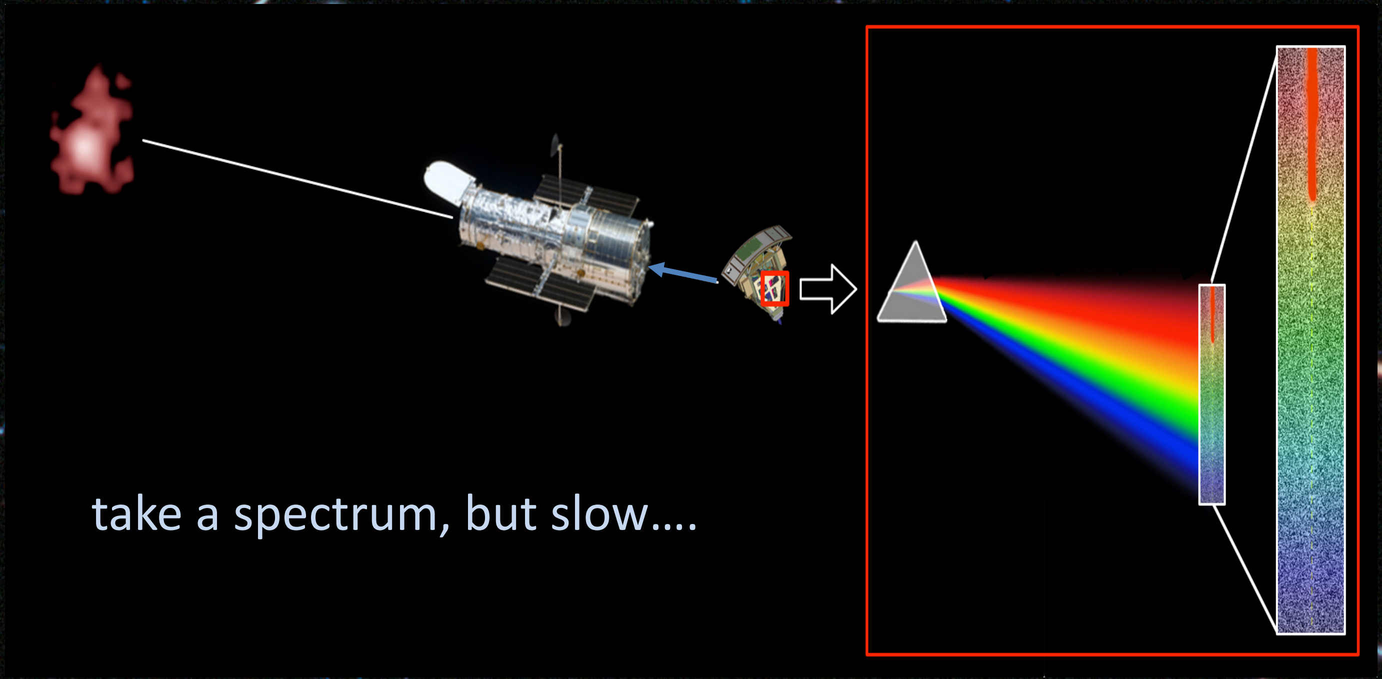


figure credit: insert adapted from Brant Robertson UCSC

gdi

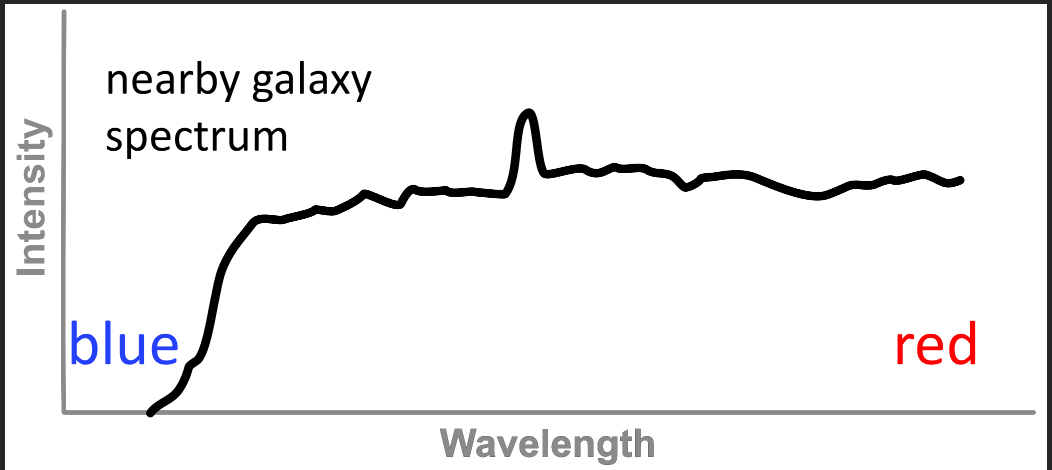
how we determine redshifts



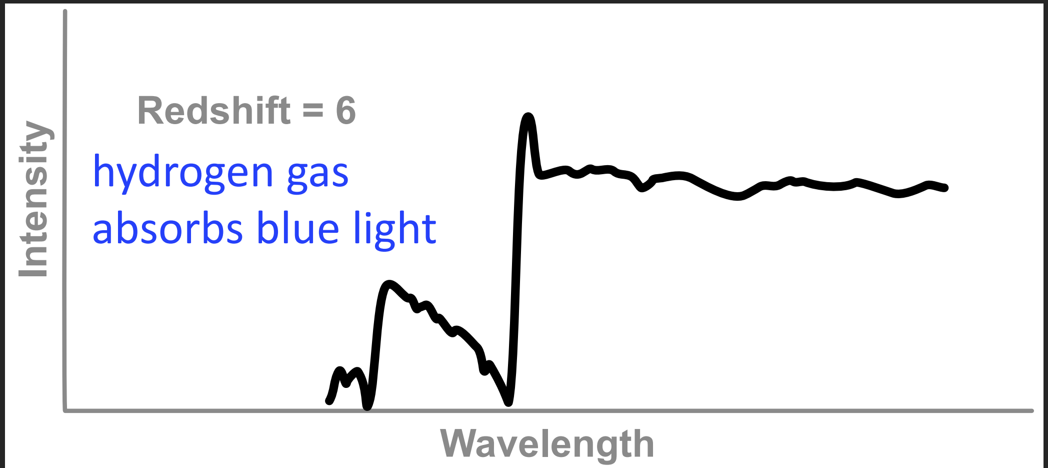
gdi

redshifts ("z")

☞ change in wavelength gives redshift

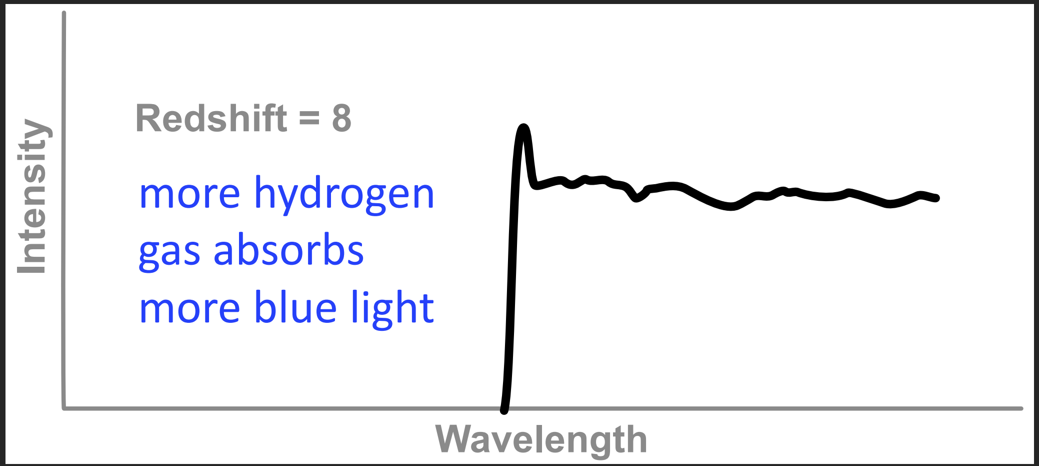


z=6 spectrum shifted to red

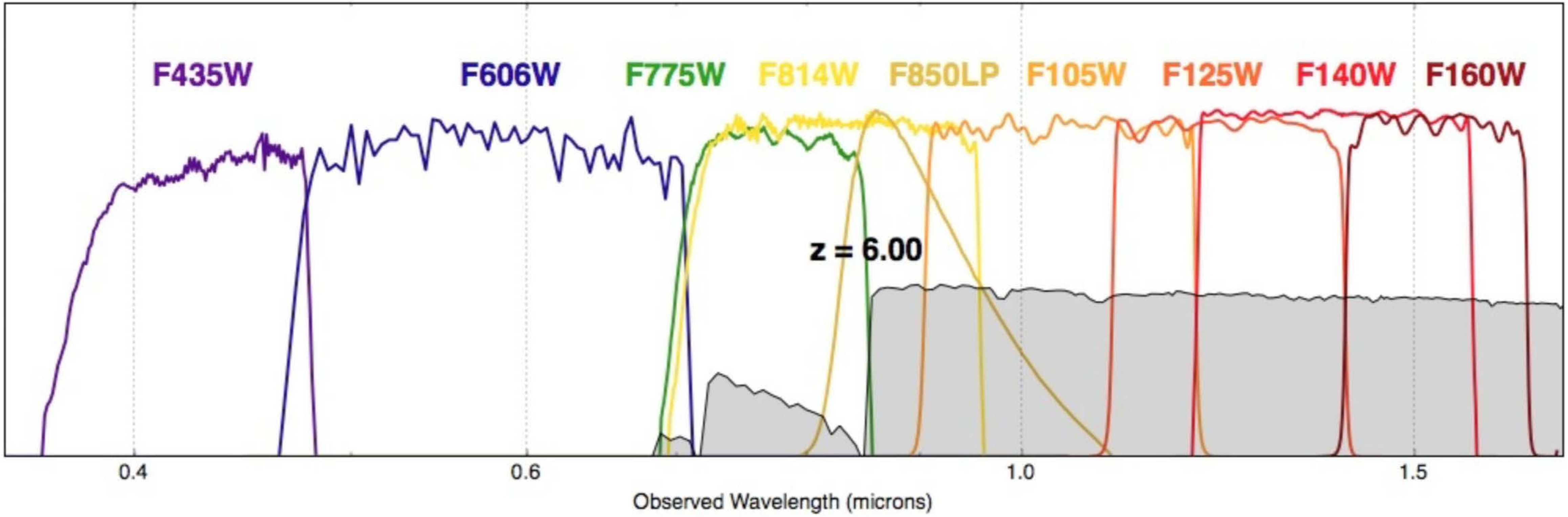


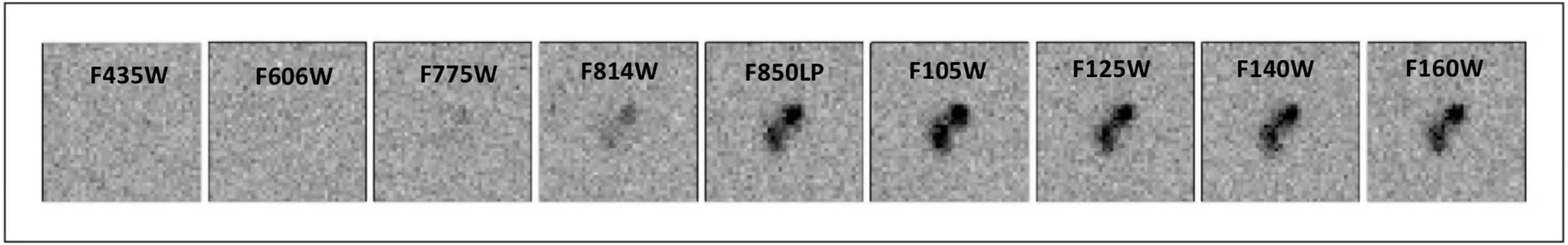
at z=6 galaxy is moving away at 96% of the speed of light!

z=8 spectrum shifted even more to red



gdi





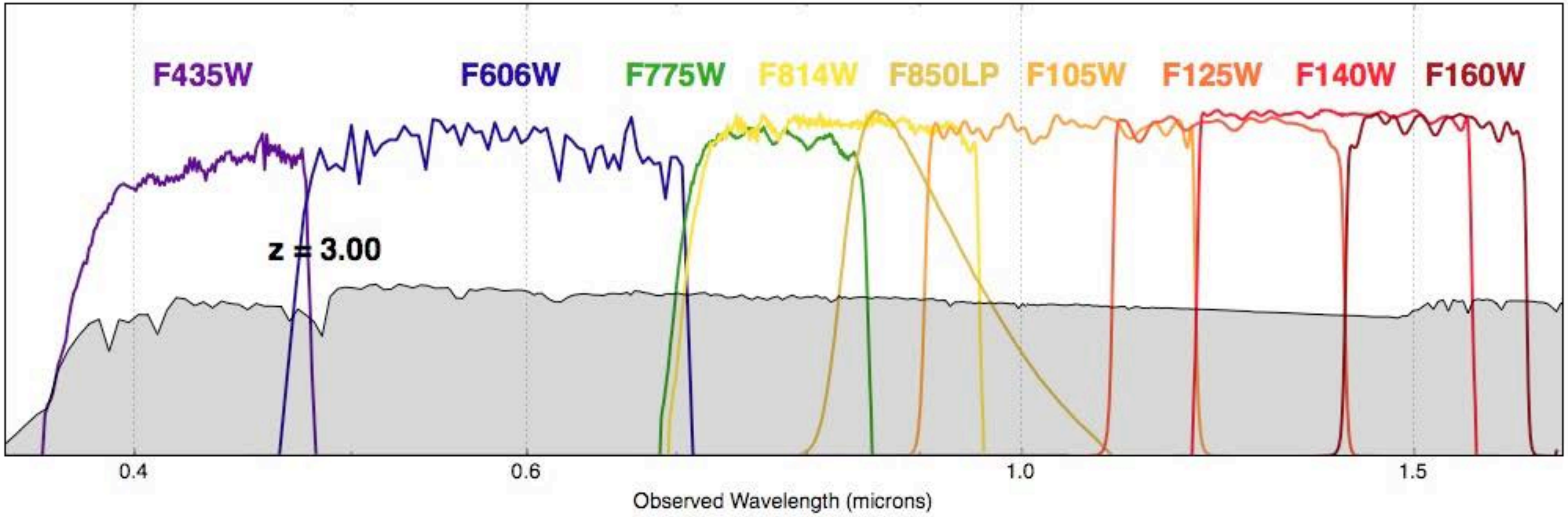
optical ACS

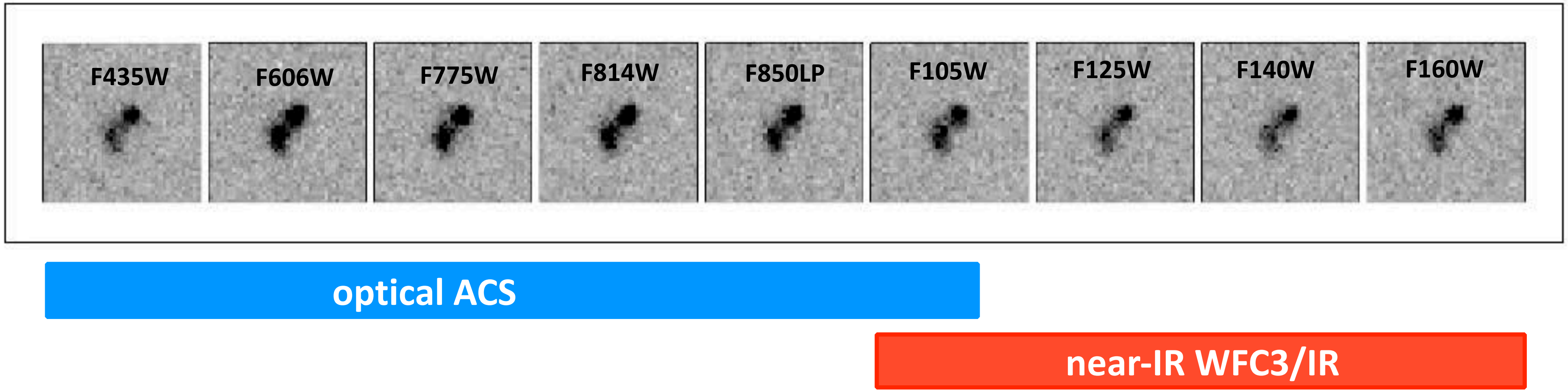
near-IR WFC3/IR

xdf.ucolick.org/

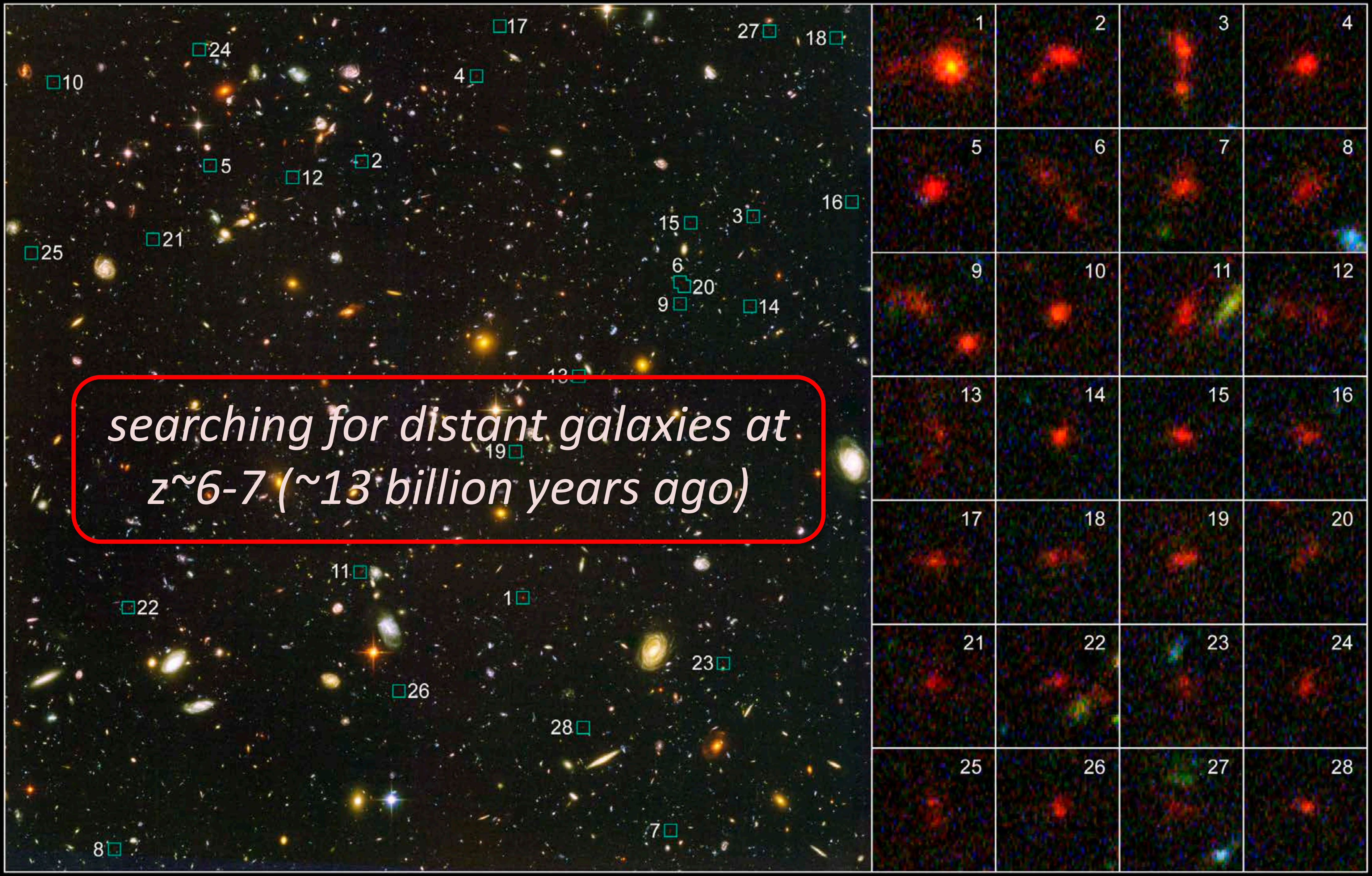
gdi

ACS+WFC3/IR: efficient detection of galaxies to z~10+



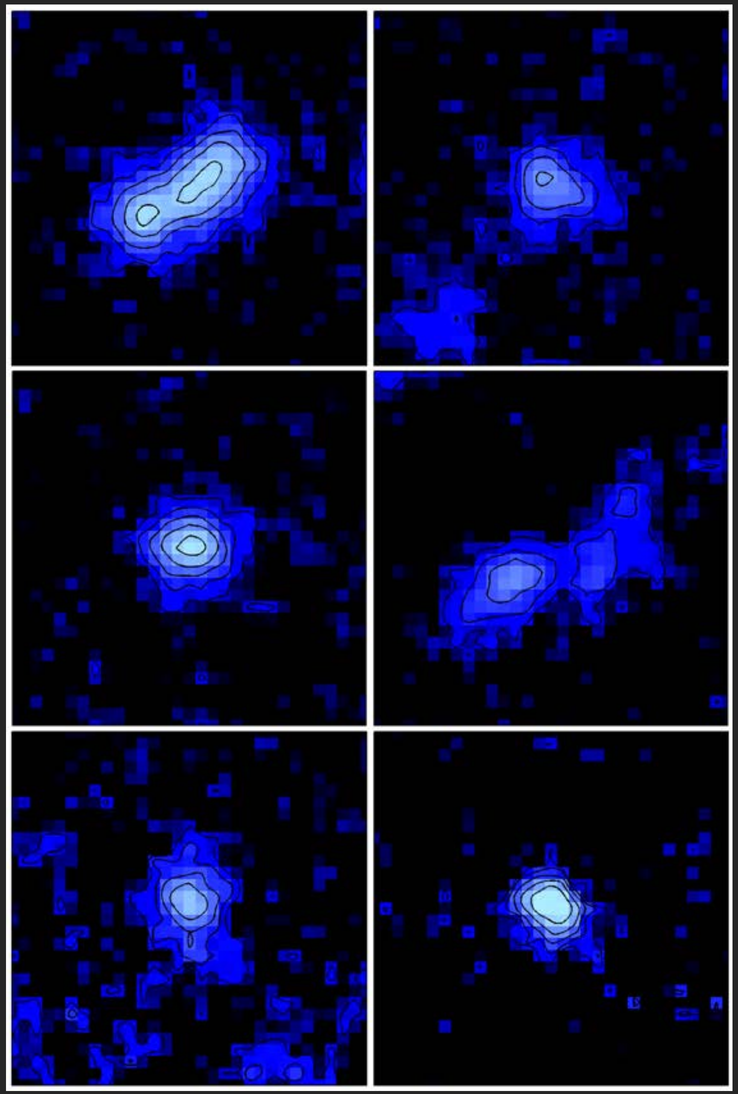


xdf.ucolick.org/



very distant galaxies look red in our images

1.8” (~10 kpc)



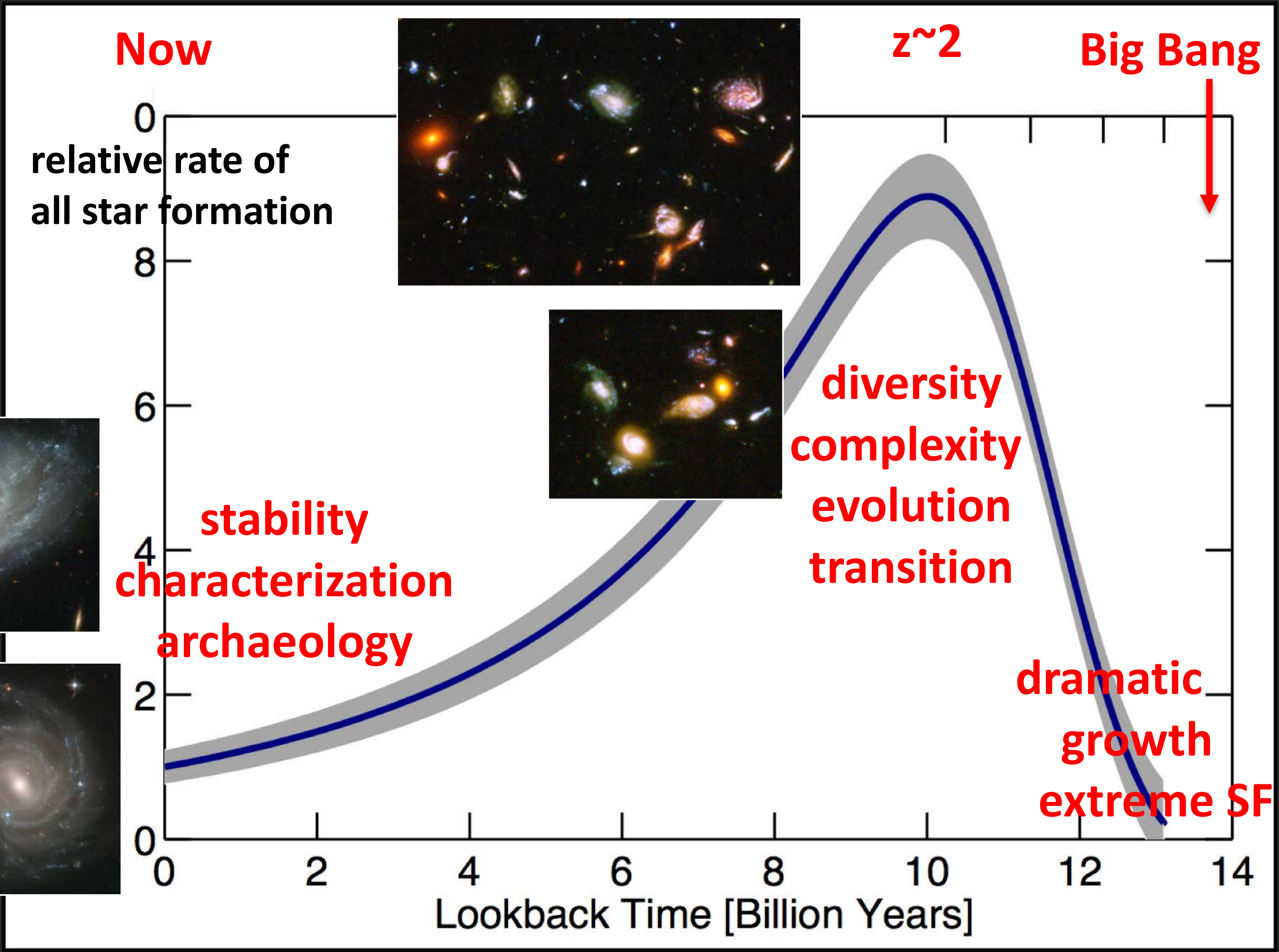
what have we learned about galaxies in the first several billion years

a sample of bright galaxies about 1 billion years after the Big Bang

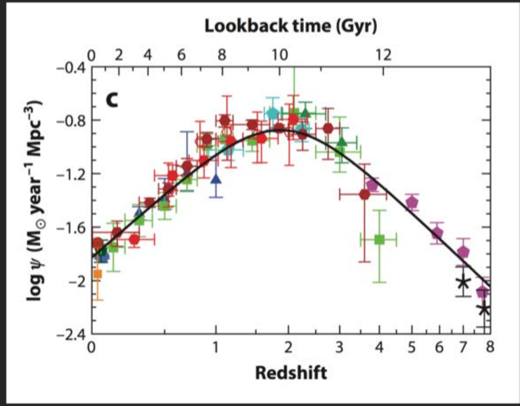
have been represented in blue to better convey what they really look like

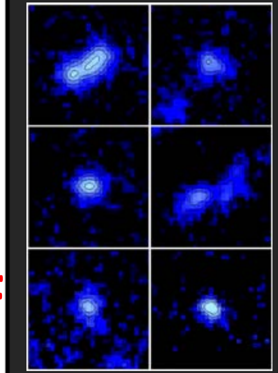
gdi

cosmic star formation over all time



revealing the star formation rate density over 96% of time

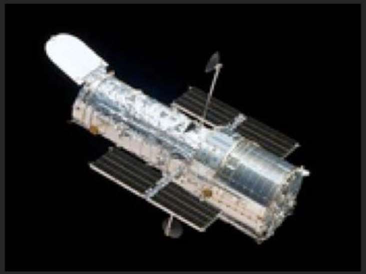


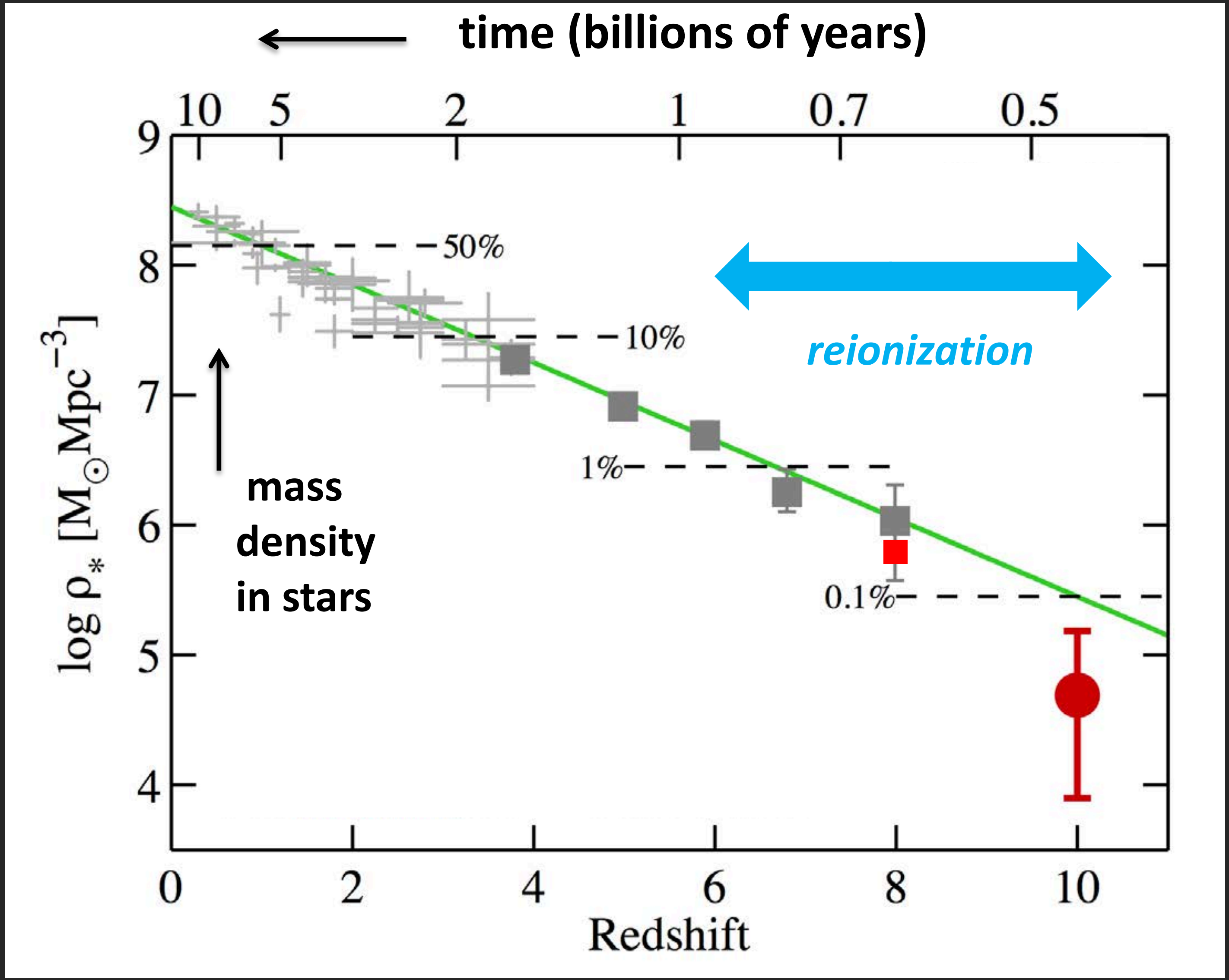


linear figure credit:
Pascal Oesch

gdi

buildup of mass in the universe in stars





only ~2% of stellar mass density built up by the end of reionization

only ~0.1% at the start of reionization at z~10 (see later)

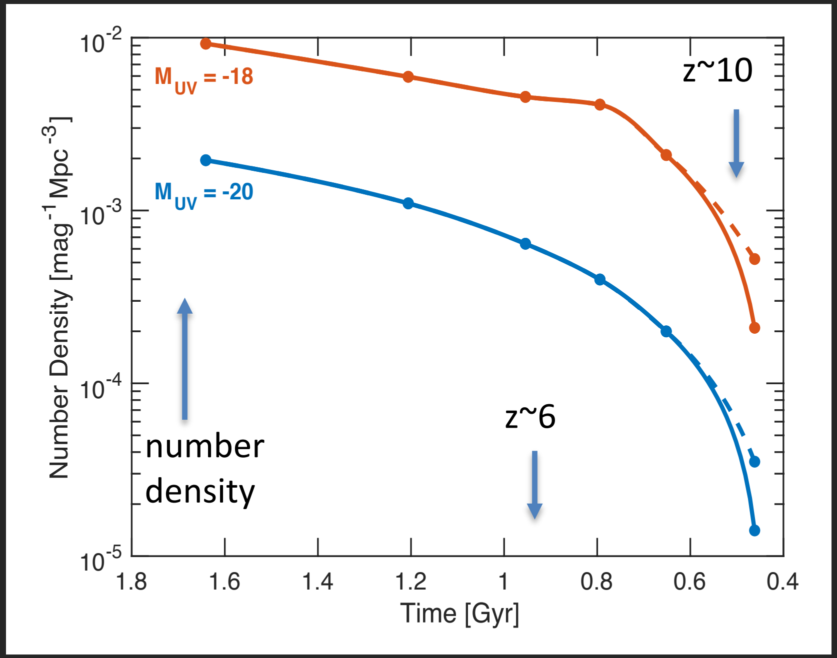
Oesch+2014

gdi

Cosmic Dawn – the time of the first stars and galaxies

the first billion years → we see galaxies building up extremely rapidly from redshift z~10 to z~6 (480 million years to 1 billion years)

a time of rapid growth of the dark matter halos within which galaxies form

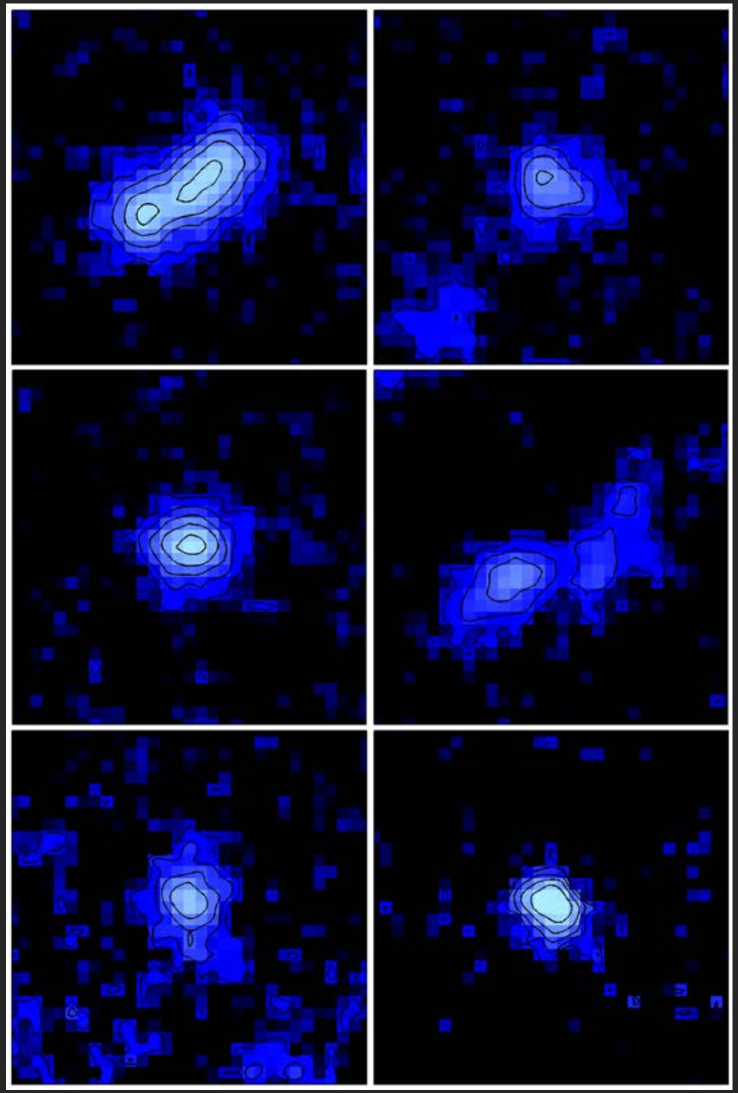


a time when significant quantities of heavier elements were produced in stars and ejected into the gas in galaxies

a time when the universe was reionized...

HST and Spitzer have let us explore in this fascinating period

gdi

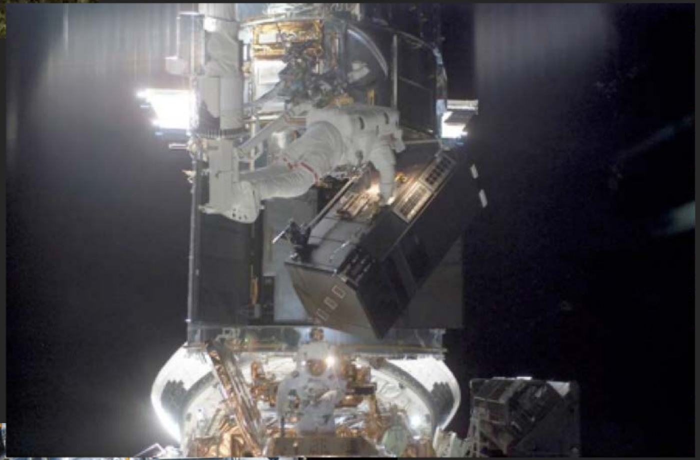


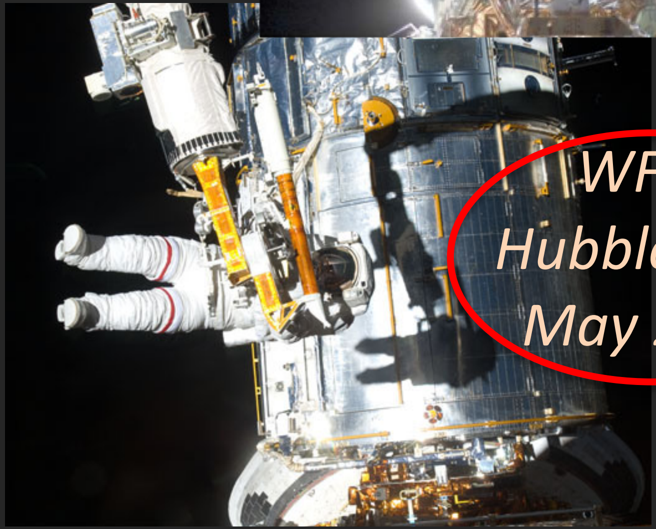
how have we learned about galaxies in the first billion years

gdi

the telescopes and cameras that enabled the exploration of the early universe

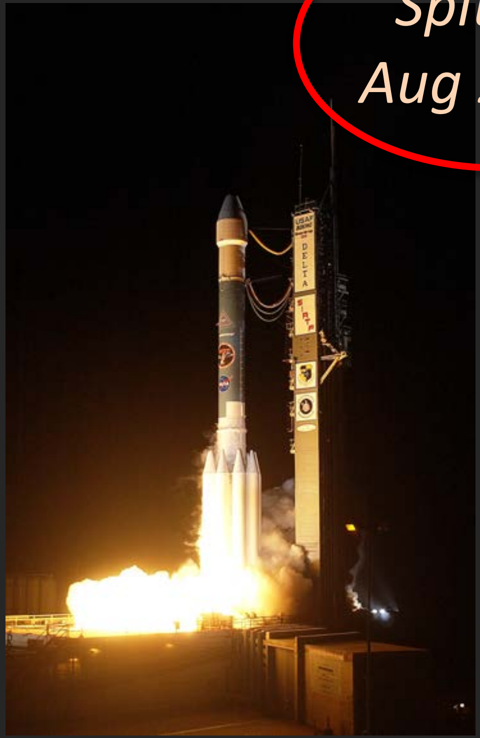
ACS
Hubble SM3B
Mar 2002





WFC3
Hubble SM4
May 2009

Spitzer
Aug 2003



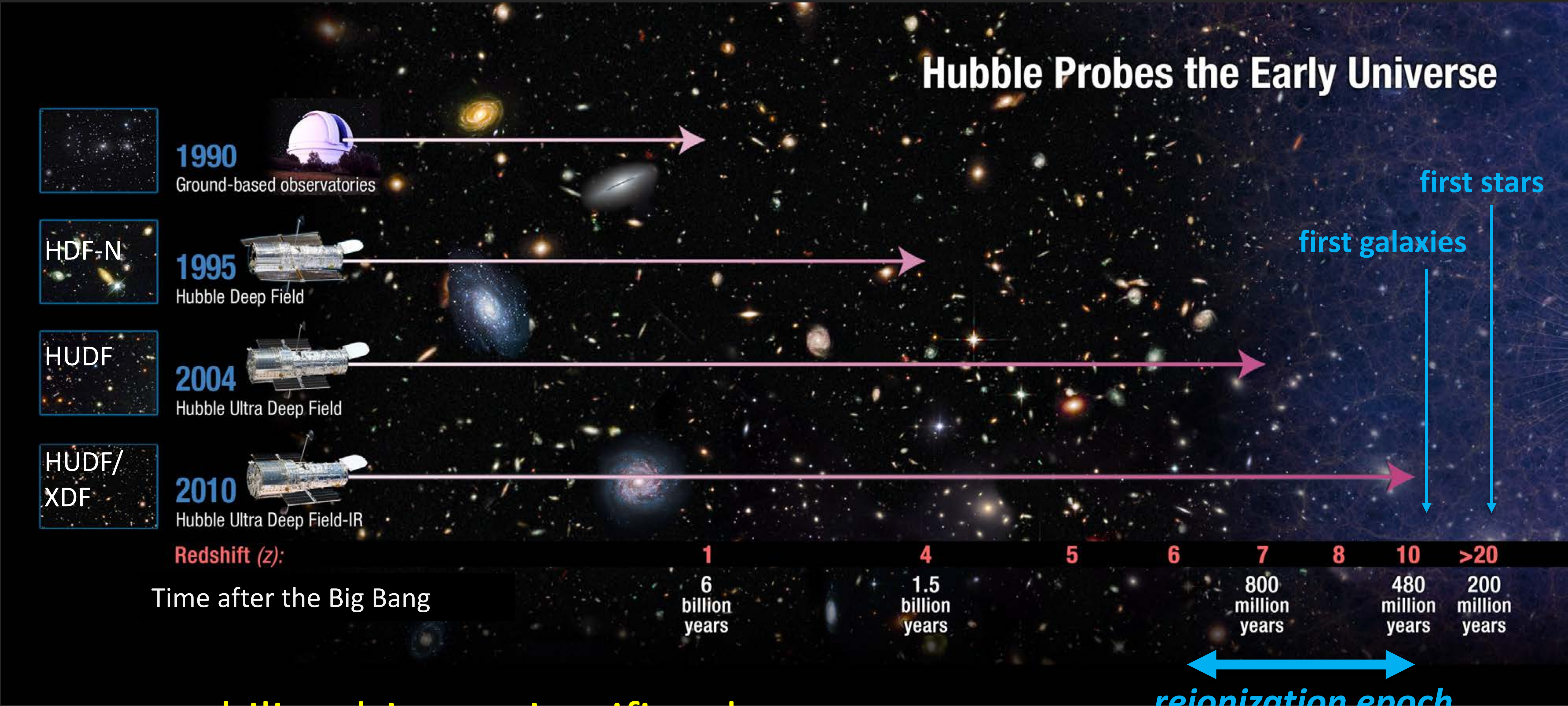
upgraded Hubble
ACS in 2002
WFC3 in 2009

launched Spitzer in 2003

each new servicing mission resulted in a dramatic change in our ability to explore the early universe

gdi

redshift limits increase with new capability



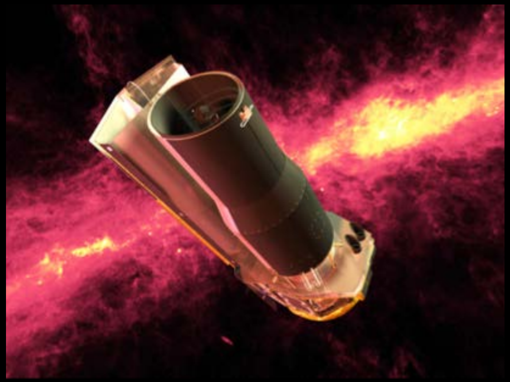
capability-driven scientific advances

gdi

Hubble's partners for distant galaxies

Chandra
Great Observatory

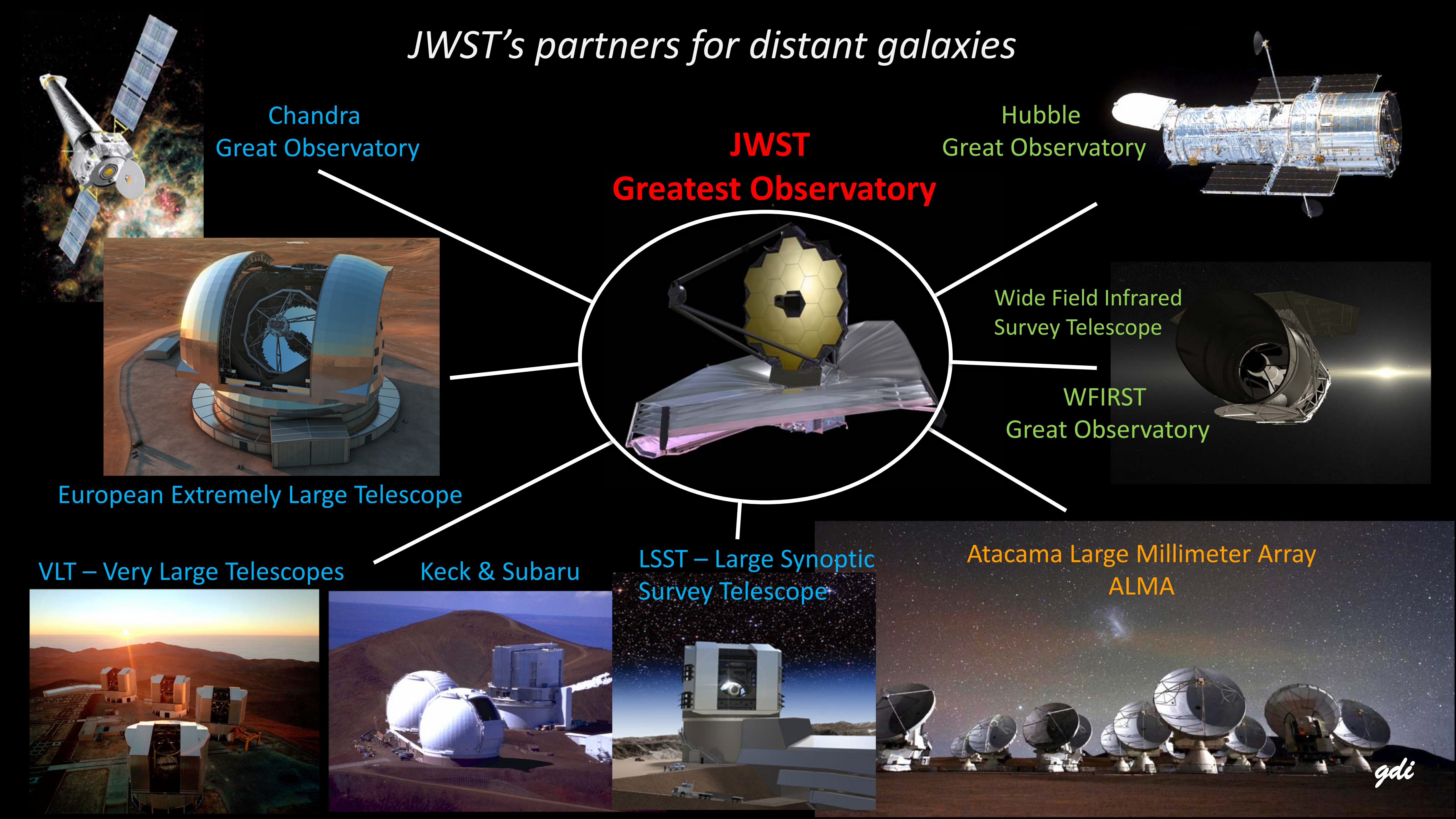
Hubble
Great Observatory

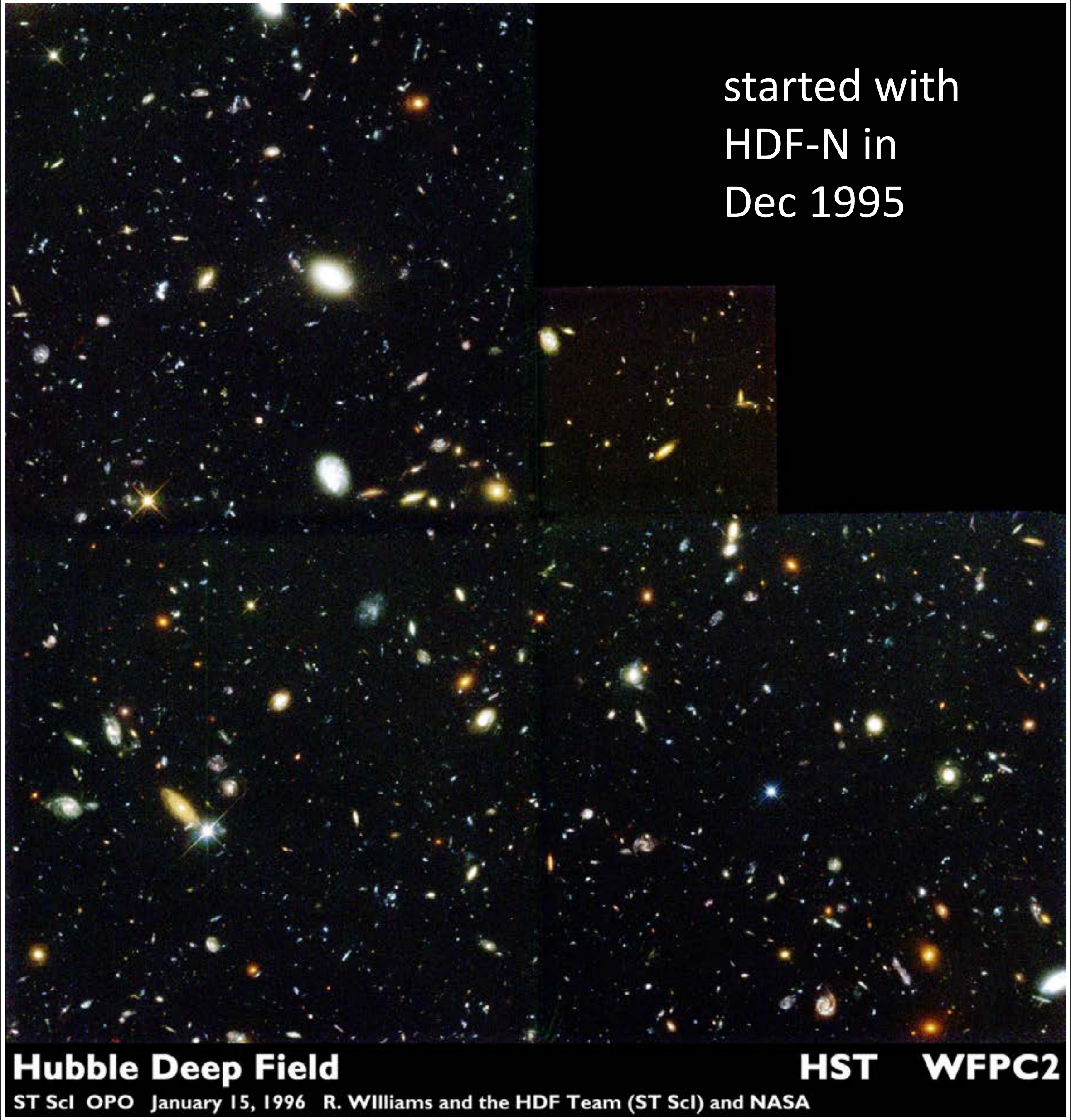


Spitzer
Great Observatory

VLT – Very Large telescope

Keck & Subaru telescopes


JWST's partners for distant galaxies
Chandra
Great Observatory
JWST
Greatest Observatory
Hubble
Great Observatory
Wide Field Infrared
Survey Telescope
WFIRST
Great Observatory
European Extremely Large Telescope
VLT – Very Large Telescopes
Keck & Subaru
LSST – Large Synoptic
Survey Telescope
Atacama Large Millimeter Array
ALMA
gdi



the survey datasets used for high-redshift galaxy studies

gdi

Hubble and Spitzer survey fields for high-redshift galaxies

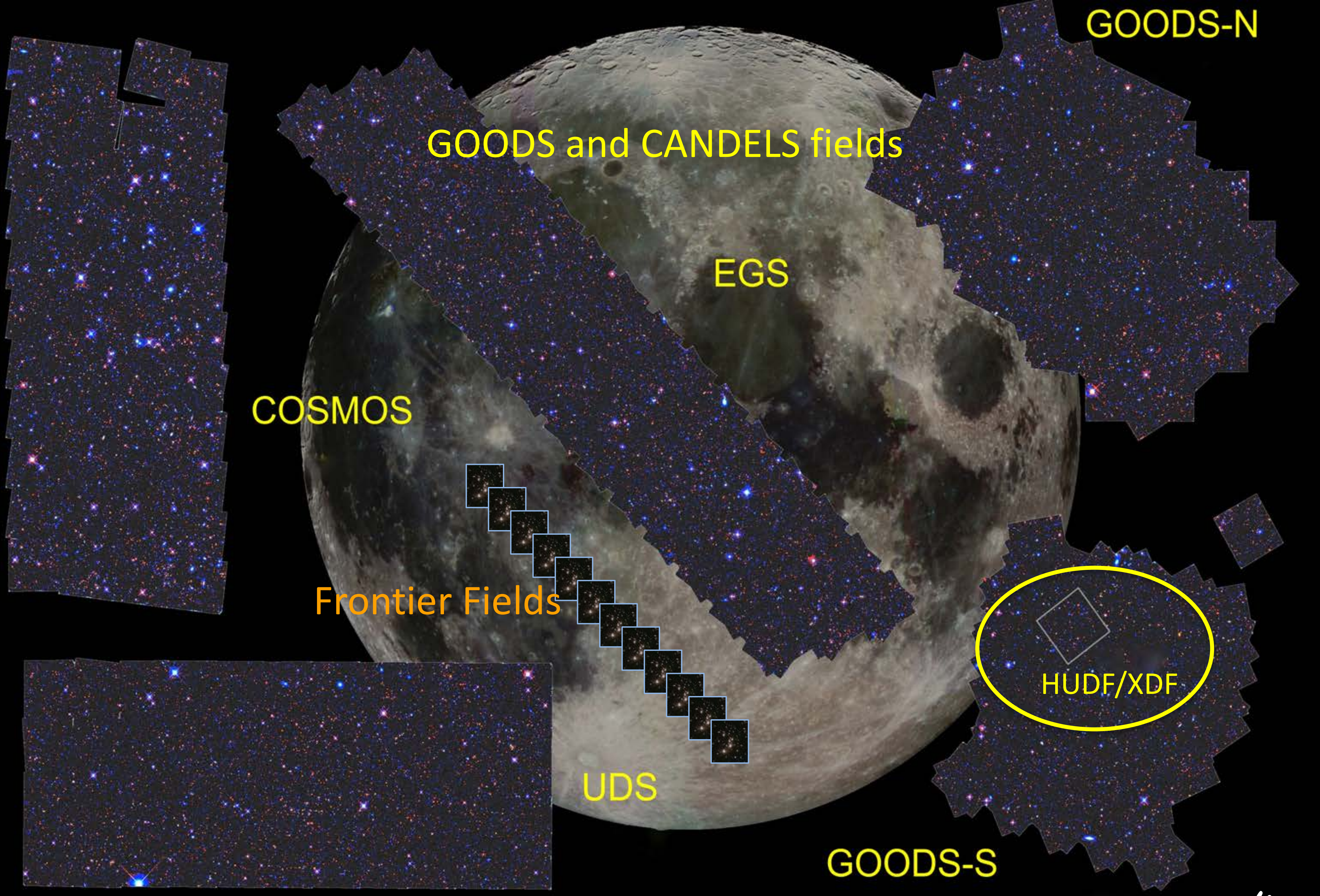


figure credit: Harry Ferguson and the CANDELS team

gdi

XDF/HUDF (*eXtreme Deep Field)*

deepest ever Hubble image

2963 HST images

from 800 orbits of Hubble

for a 23 day total exposure on the HUDF!

all optical ACS data and all infrared WFC3/IR data on the HUDF from 2003-2013 from 19 programs

reaches ~31 AB mag 5σ
or >32.5 AB mag 1σ

xdf.ucolick.org

HUBBLE SPACE TELESCOPE
XDF ▪ EXTREME DEEP FIELD

A decade of imaging on the Hubble Ultra Deep Field
The deepest image of the Universe

GDI+2013

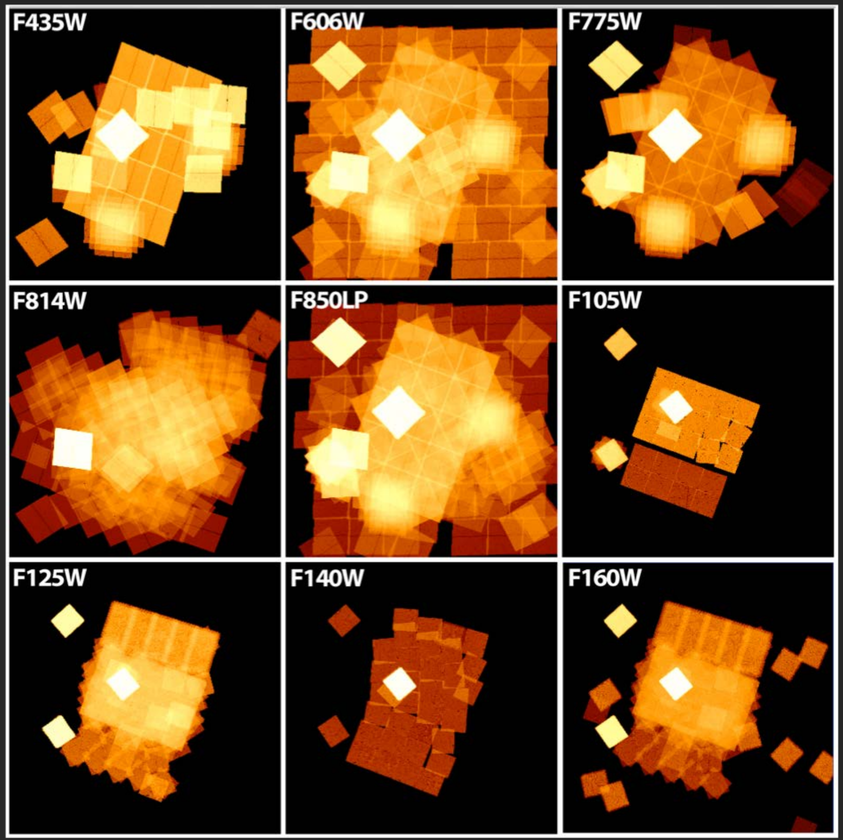
2012
NASA, ESA,
G. ILLINGWORTH, D. MAGEE, AND P. OESCH (UNIVERSITY OF CALIFORNIA, SANTA CRUZ),
R. BOUWENS (LEIDEN UNIVERSITY), AND THE XDF TEAM

gdi

CDF-S/GOODS-S/CANDELS-S – a remarkable and unique region

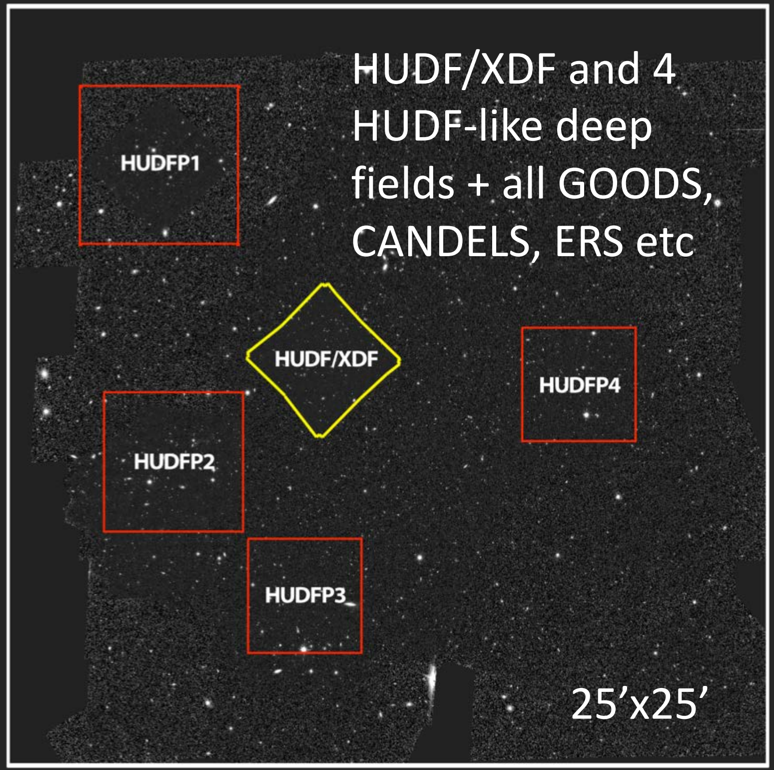
Hubble Legacy Field South (HLF-GOODS-S)

7211 exposures from 2442 orbits



ACS + WFC3/IR – 10 filters in V1.5 release (+098M)

5.8 Msec or ~70% of a Hubble Cycle



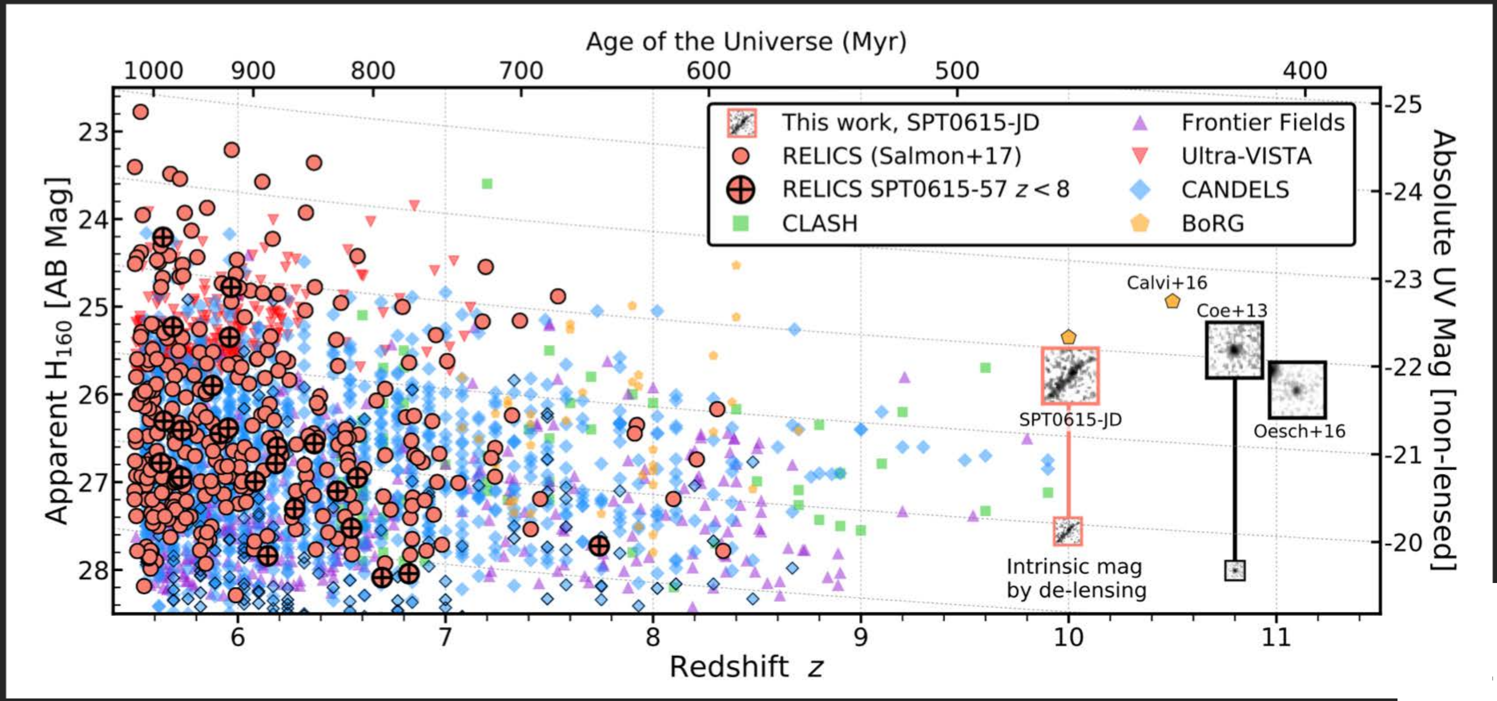
152 GB of aligned astrometric HST images

firstgalaxies.org/hlf & archive.stsci.edu/prepds/hlf/

GDI Magee+2016

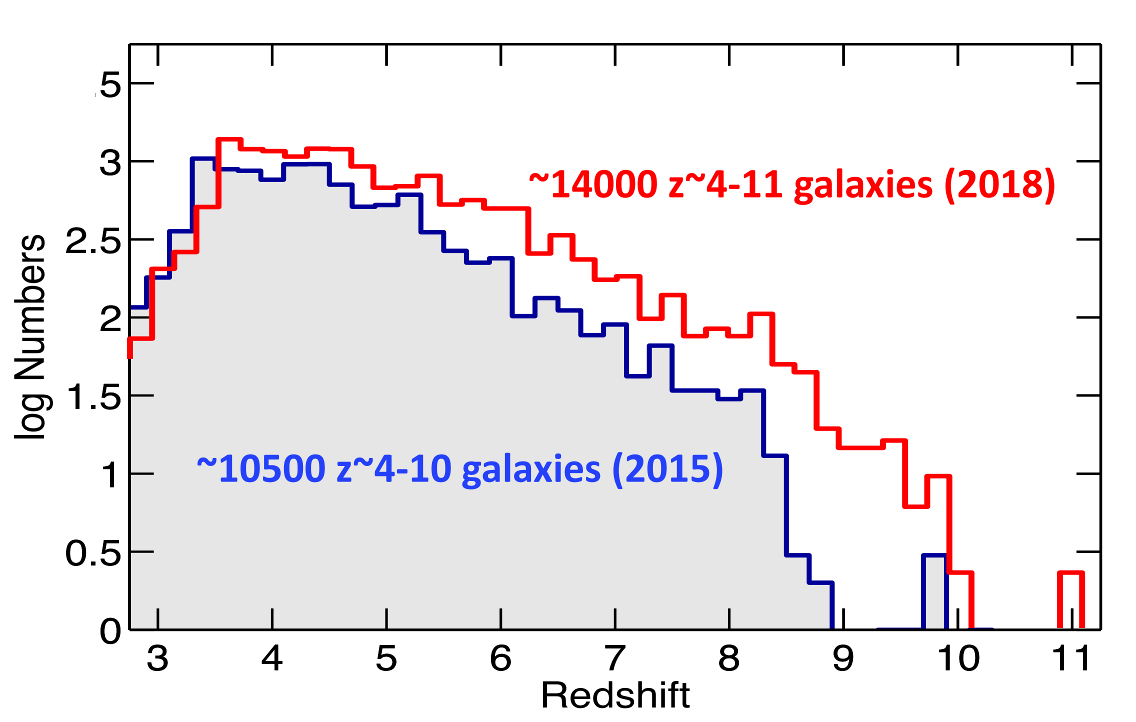
high-redshift galaxies

WFC3/IR opened up the reionization epoch at z>6



Salmon+2017

Hubble Frontier Fields have now added many more galaxies



Bouwens GDI Oesch+15

very large samples of z>4 galaxies from Hubble

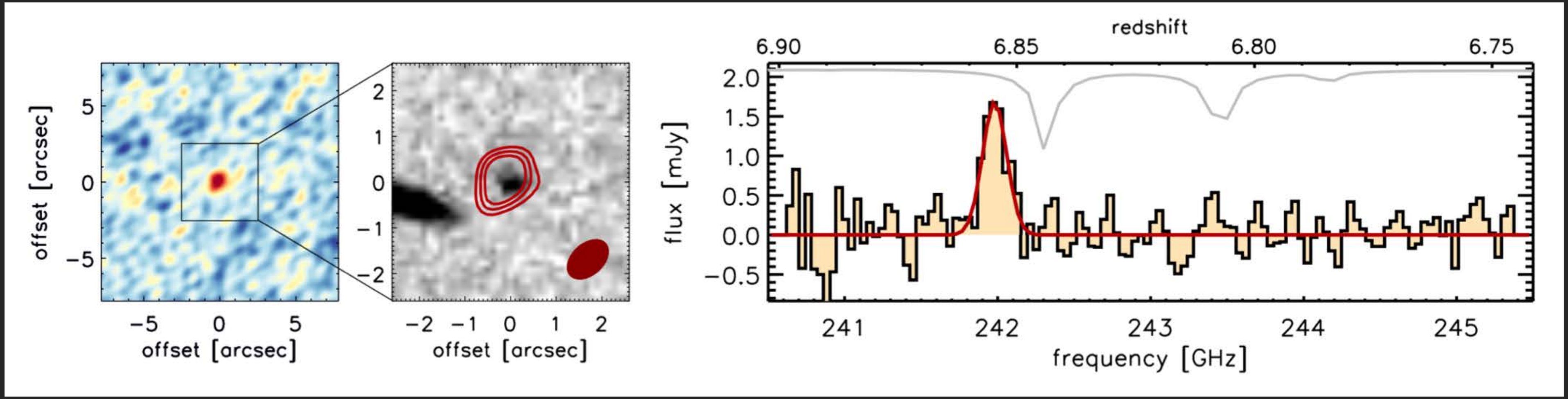
ground surveys also now have resulted in very large samples but largely at z<5-6 – (and are not shown)

gdi

another remarkable new measurement from ALMA

gdi

ALMA [C II] 157.74μm redshifts and velocity structure in two z~6.8 galaxies



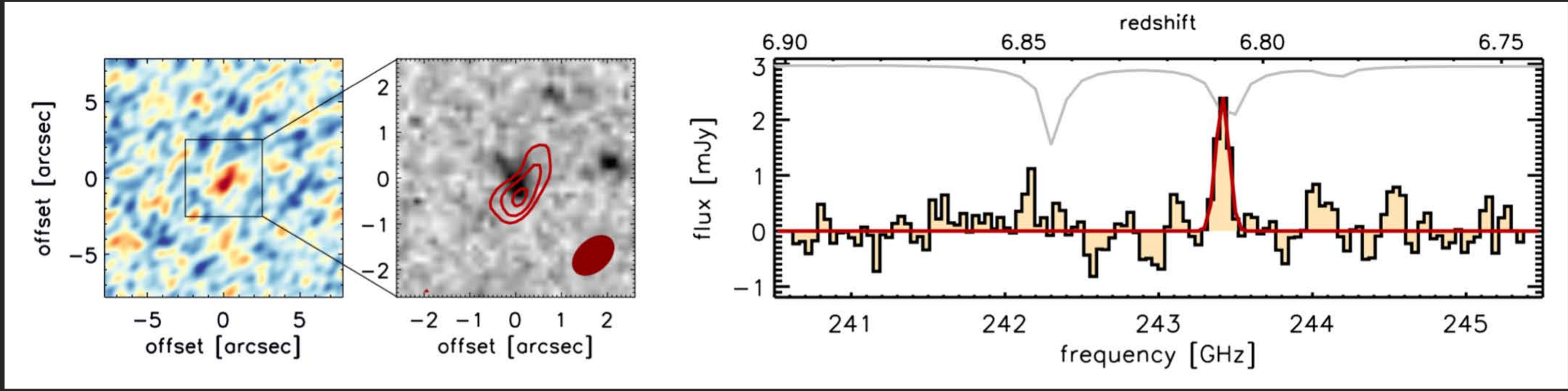
COS-3018555981

ALMA narrowband

HST image + ALMA (+ ALMA beam)

$z_{[CII]}$ = 6.8540±0.0003

$z_{[CII]}$ = 6.8076±0.0002



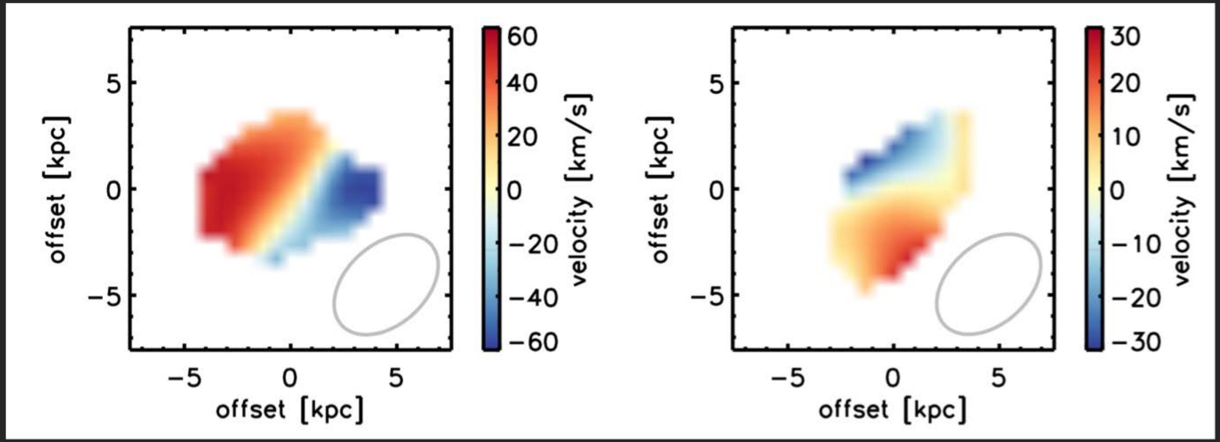
COS-2987030247

Smit + 2017

gdi

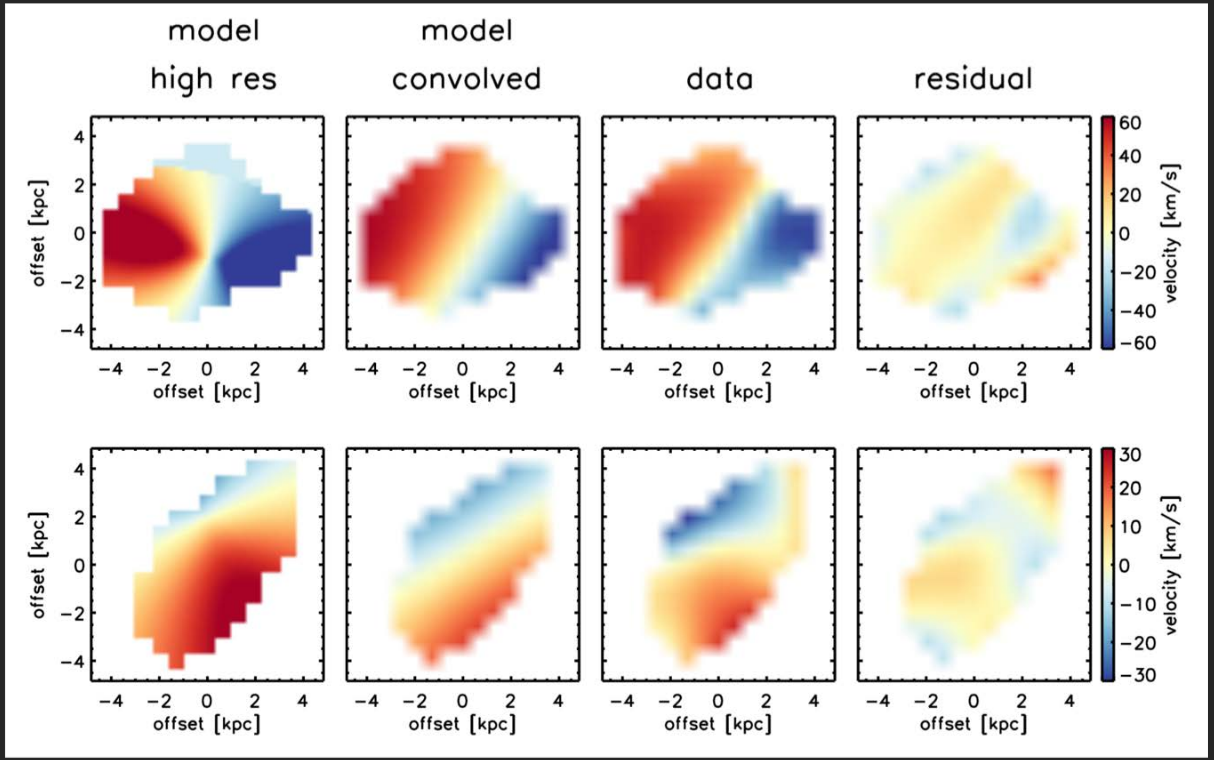
ALMA [C II] 157.74μm redshifts and velocity structure in two z~6.8 galaxies

consistent with rotation but could be more complex (merging?; gas flows?)



velocity structure in the two galaxies

rotation models compared to data

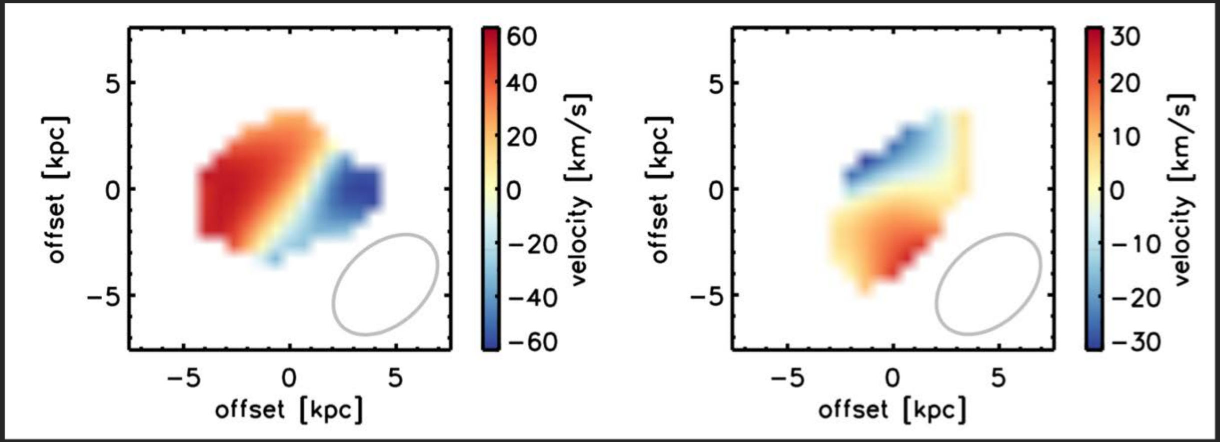


Smit + 2017

gdi

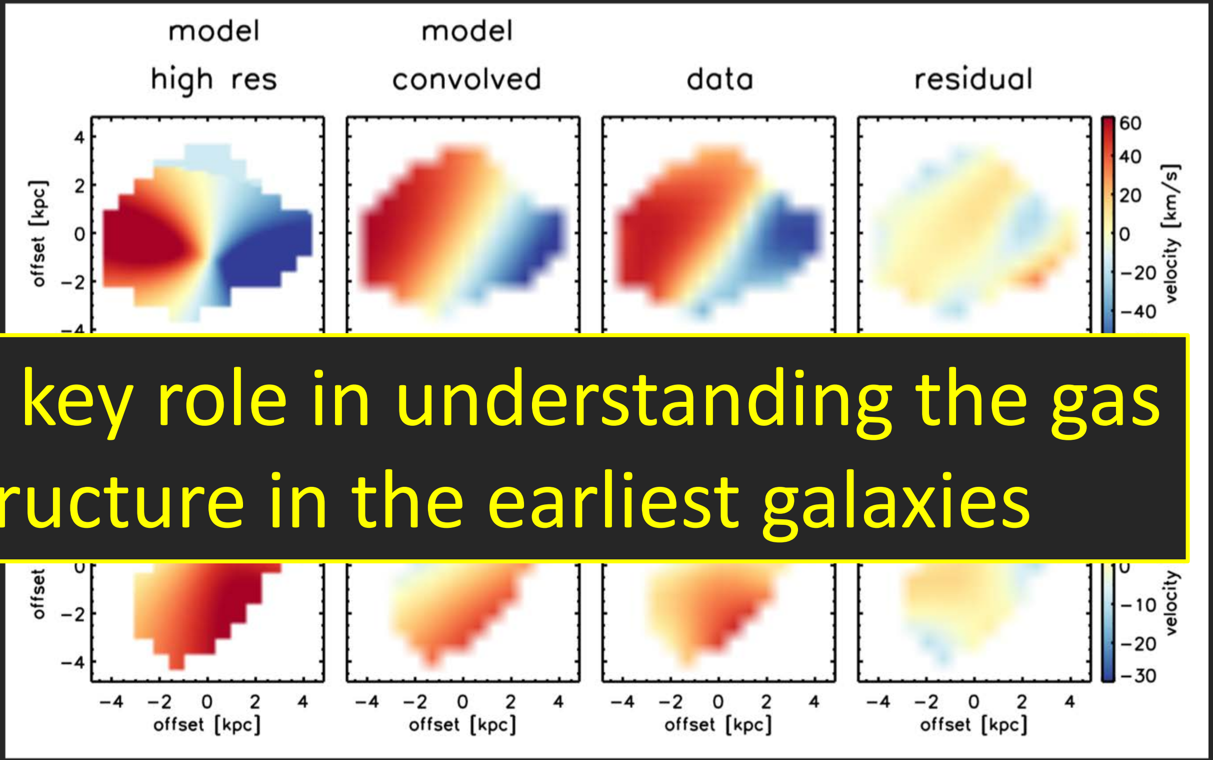
ALMA [C II] 157.74μm redshifts and velocity structure in two z~6.8 galaxies

consistent with rotation but could be more complex (merging?; gas flows?)

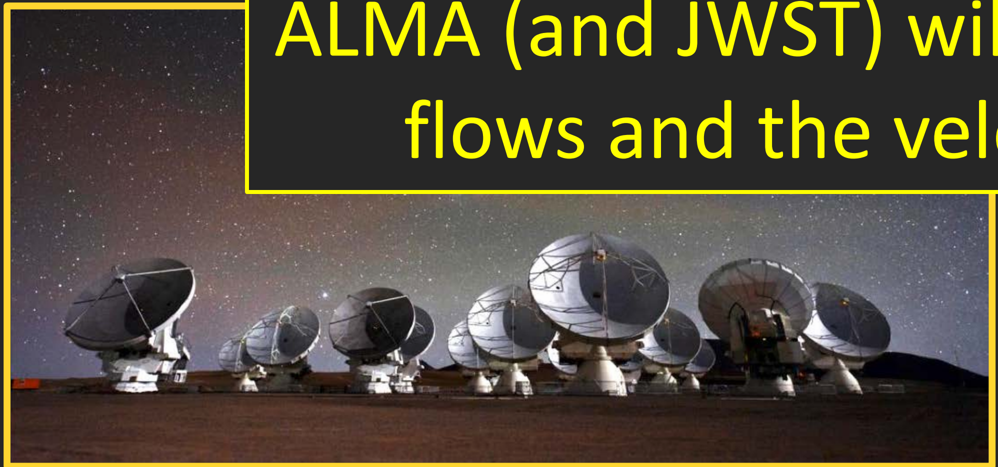


velocity structure in the two galaxies

rotation models compared to data



ALMA (and JWST) will play a key role in understanding the gas flows and the velocity structure in the earliest galaxies



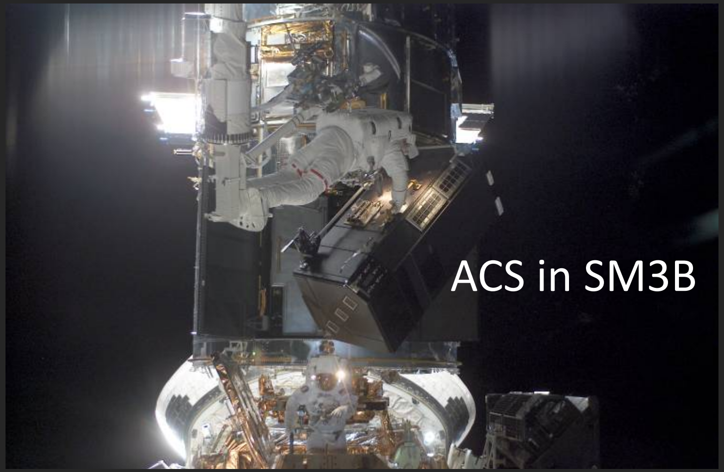
Smit + 2017

gdi

luminosity functions – the census of galaxies:
a key input for understanding galaxy build-up and reionization

over ***10,000*** high redshift Hubble-selected
galaxies from z~4 to z~10!

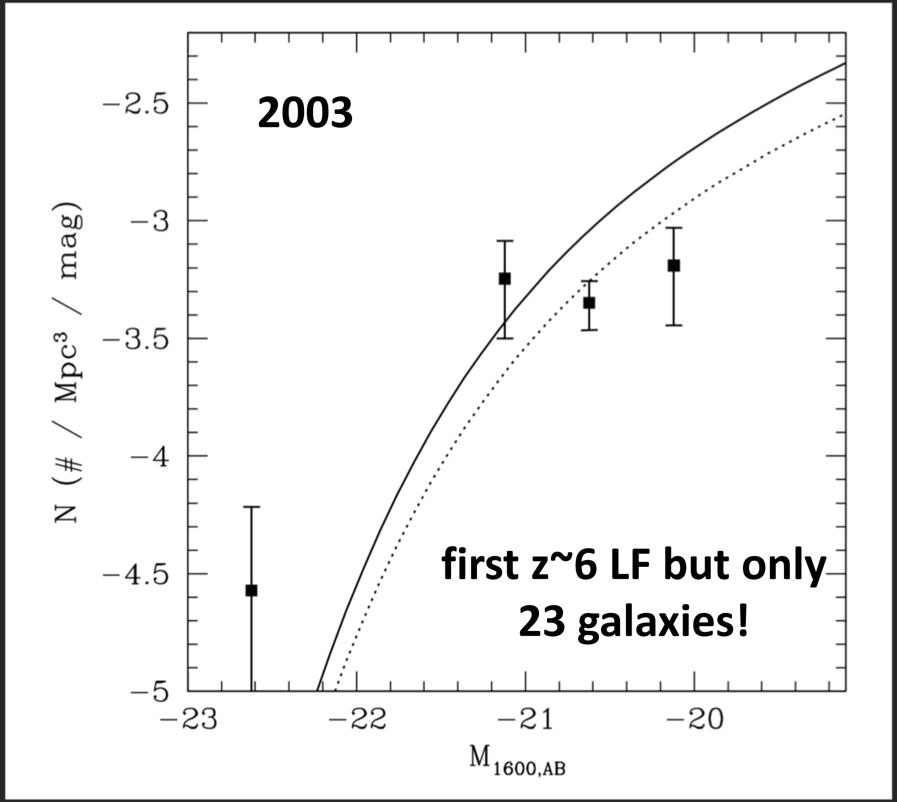
gdi



ACS enabled the first redshift z~6 sample

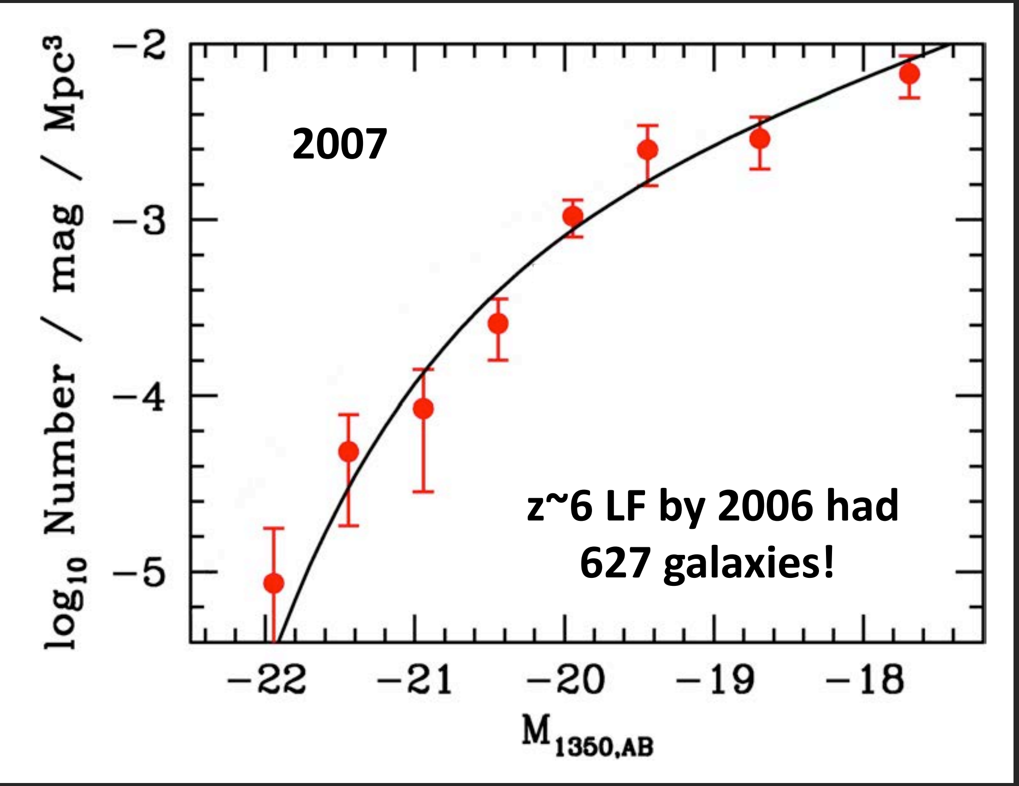
- ACS: 10-20X "discovery efficiency" of WFPC2 (more galaxies)
- enhanced wavelength coverage (higher redshift galaxies)

first luminosity function at 1 Gyr (z~6)



Bouwens GDI+2003

3 years later: 27X (627) as many z~6 galaxies



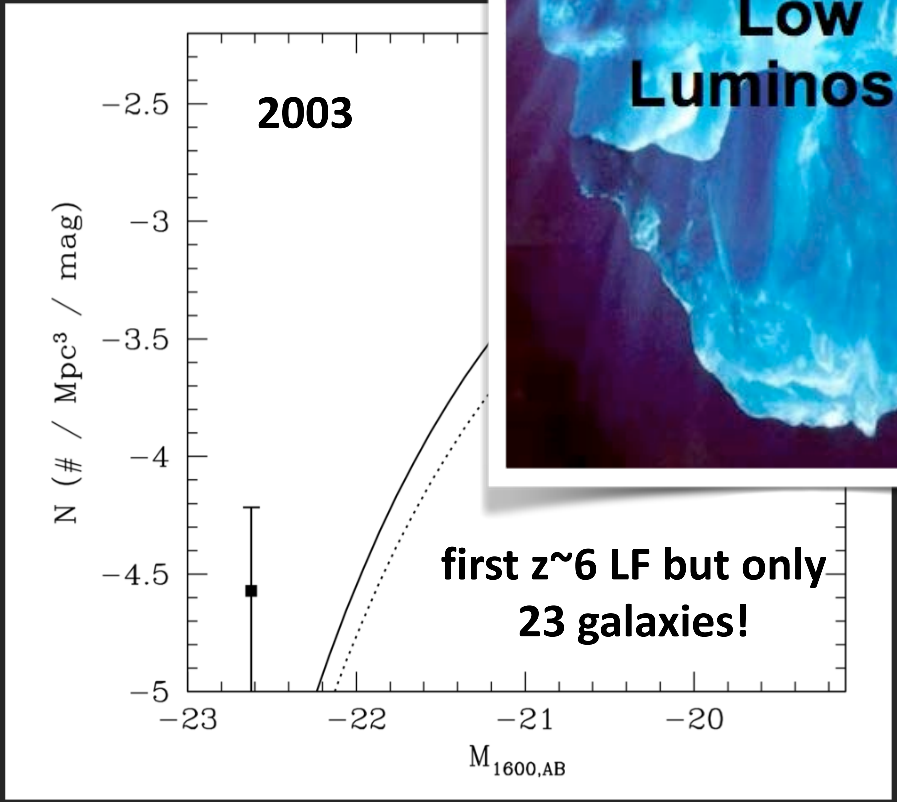
Bouwens GDI+2007

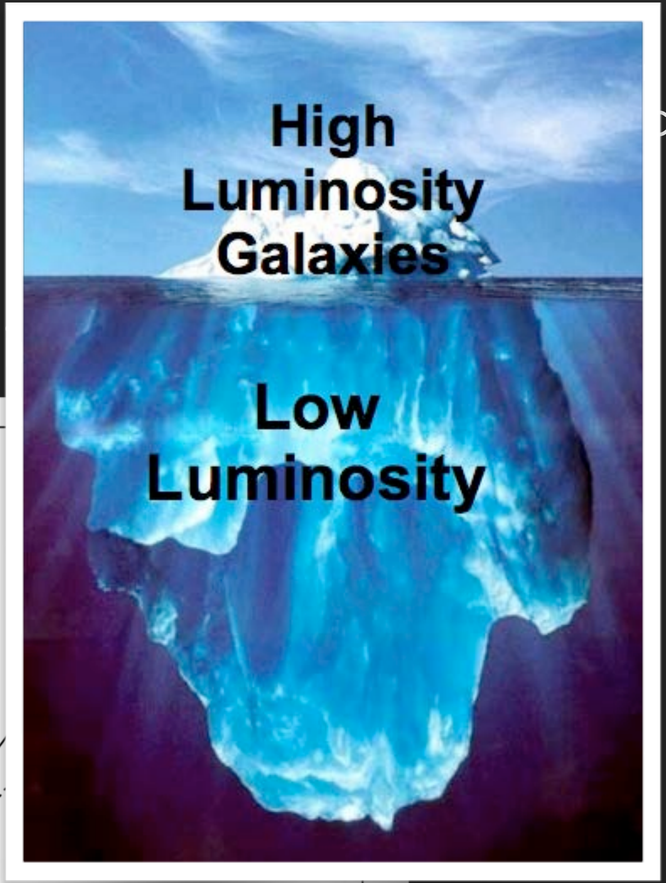
gdi

ACS enabled the first redshift z~6 sample

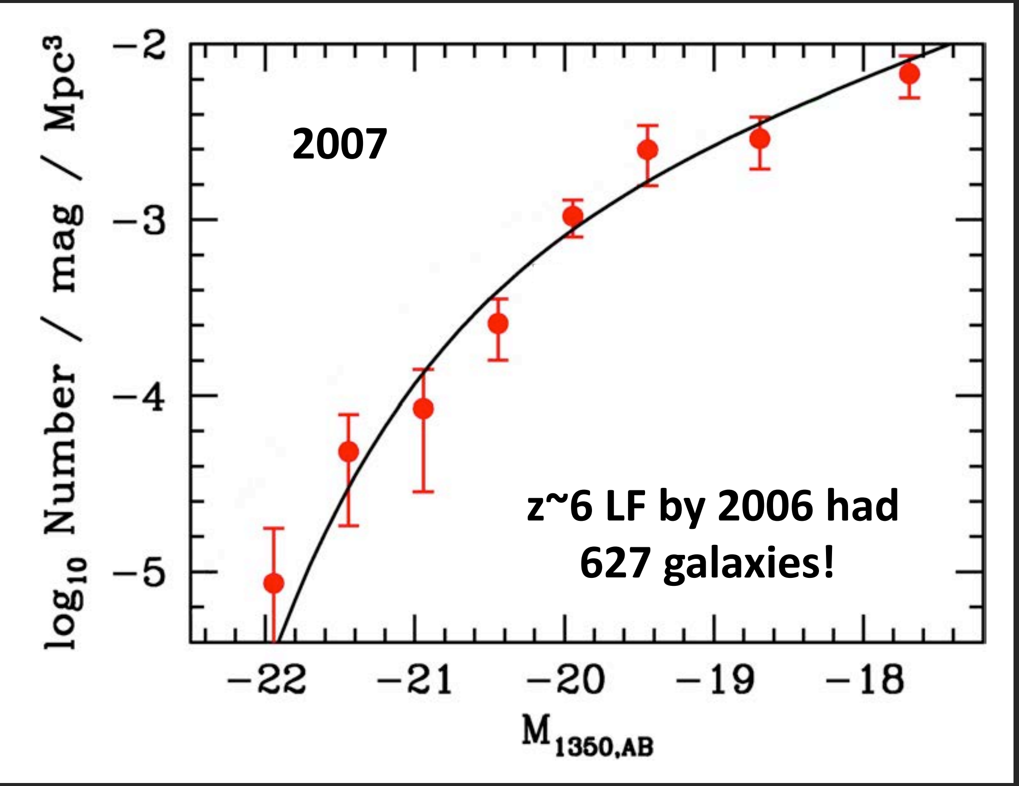
efficiency” of WFPC2 (more galaxies)
coverage (higher redshift galaxies)

first luminosity funct





3 years later: 27X (627) as many z~6 galaxies

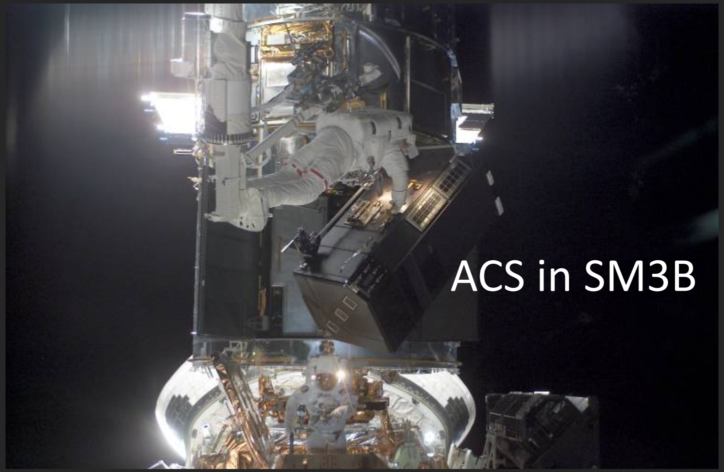


Bouwens GDI+2003

Bouwens GDI+2007

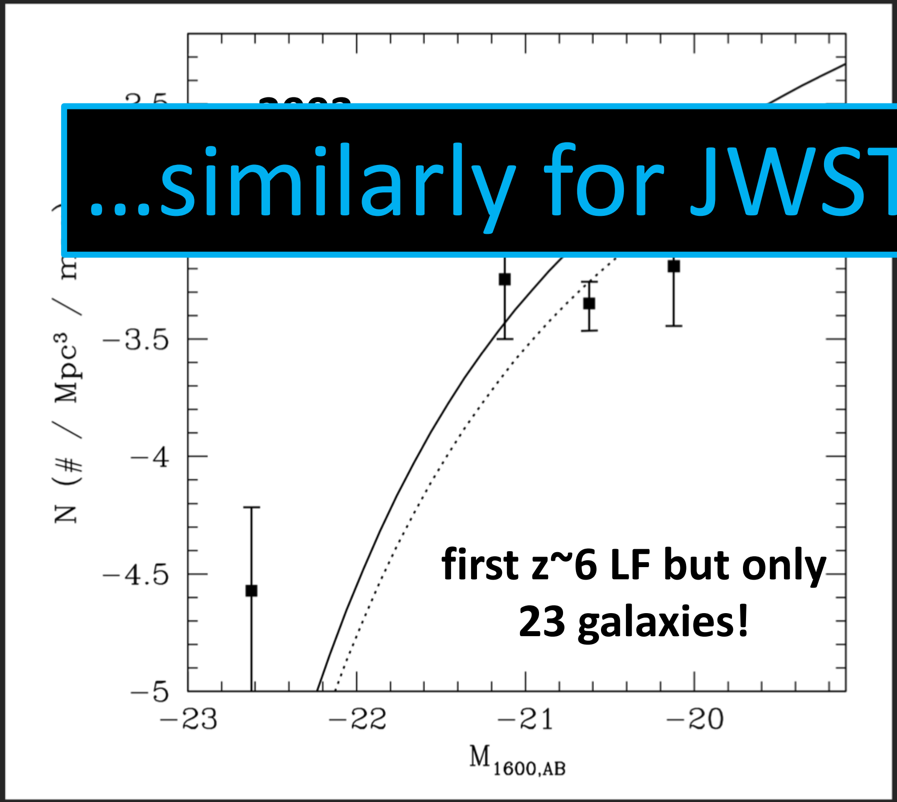
gdi

ACS enabled the first redshift z~6 sample



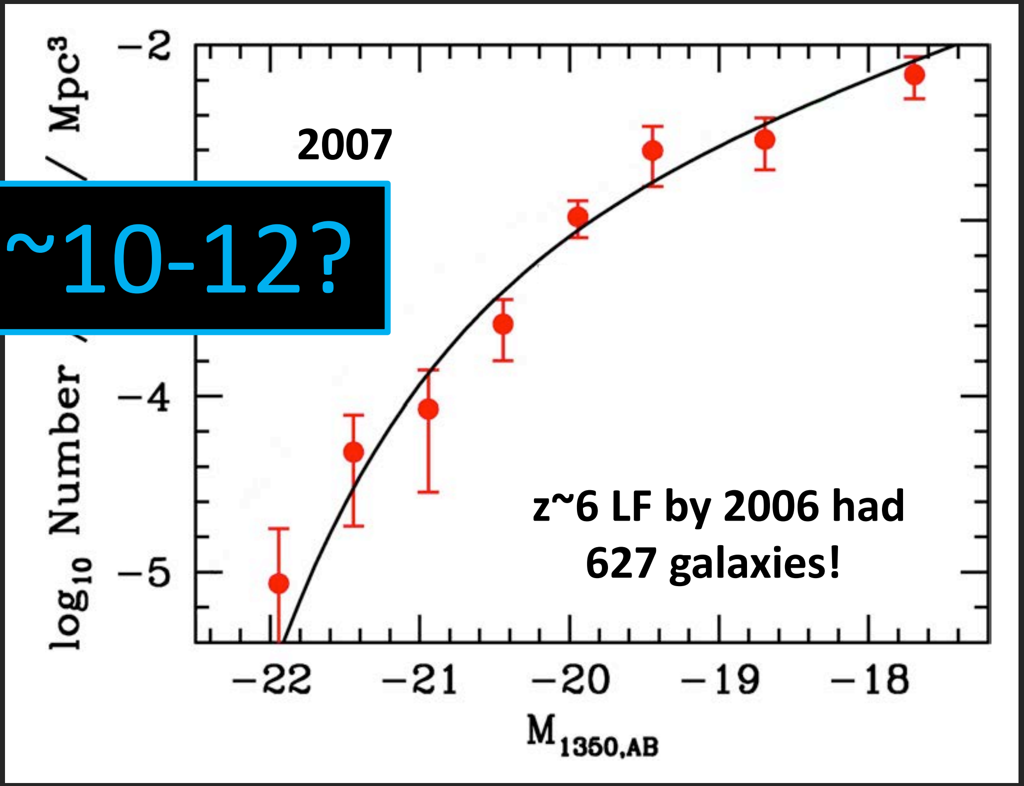
- ACS: 10-20X “discovery efficiency” of WFPC2 (more galaxies)
- enhanced wavelength coverage (higher redshift galaxies)

first luminosity function at 1 Gyr (z~6)



...similarly for JWST and z~10-12?

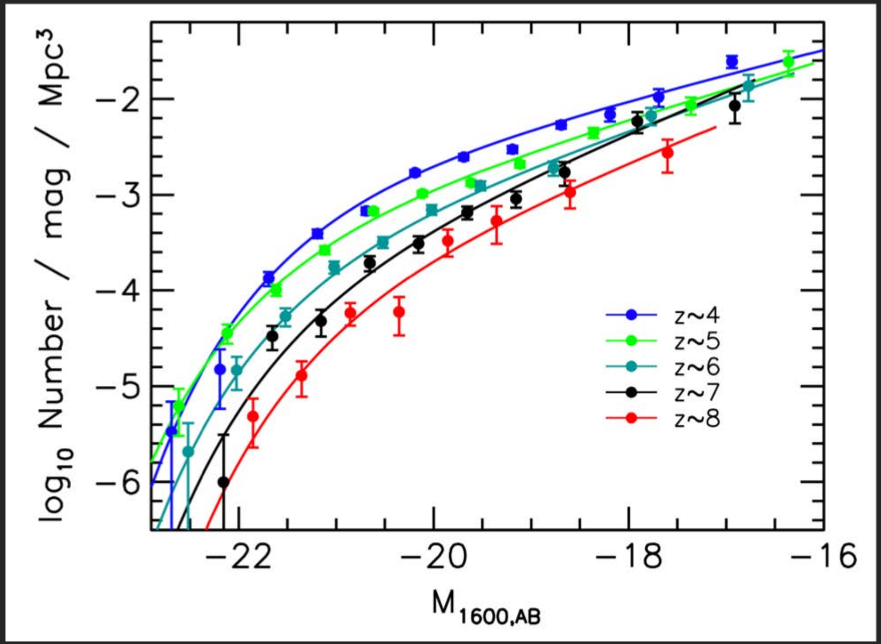
3 years later: 27X (627) as many z~6 galaxies

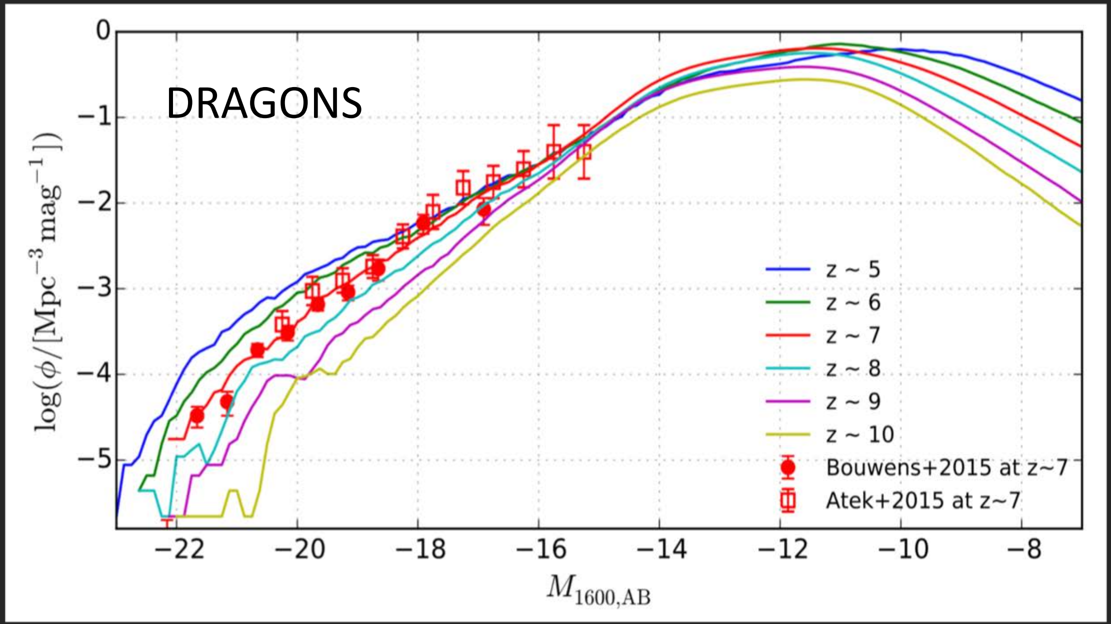


Bouwens GDI+2003

Bouwens GDI+2007

pushing LFs to fainter limits to derive UV luminosity densities





theoretical LFs Liu+2016

need to go faint to very low luminosities since majority of UV luminosity density with α~-2 comes from very faint galaxies

expect flattening or turn-over in the UV LF at low luminosities

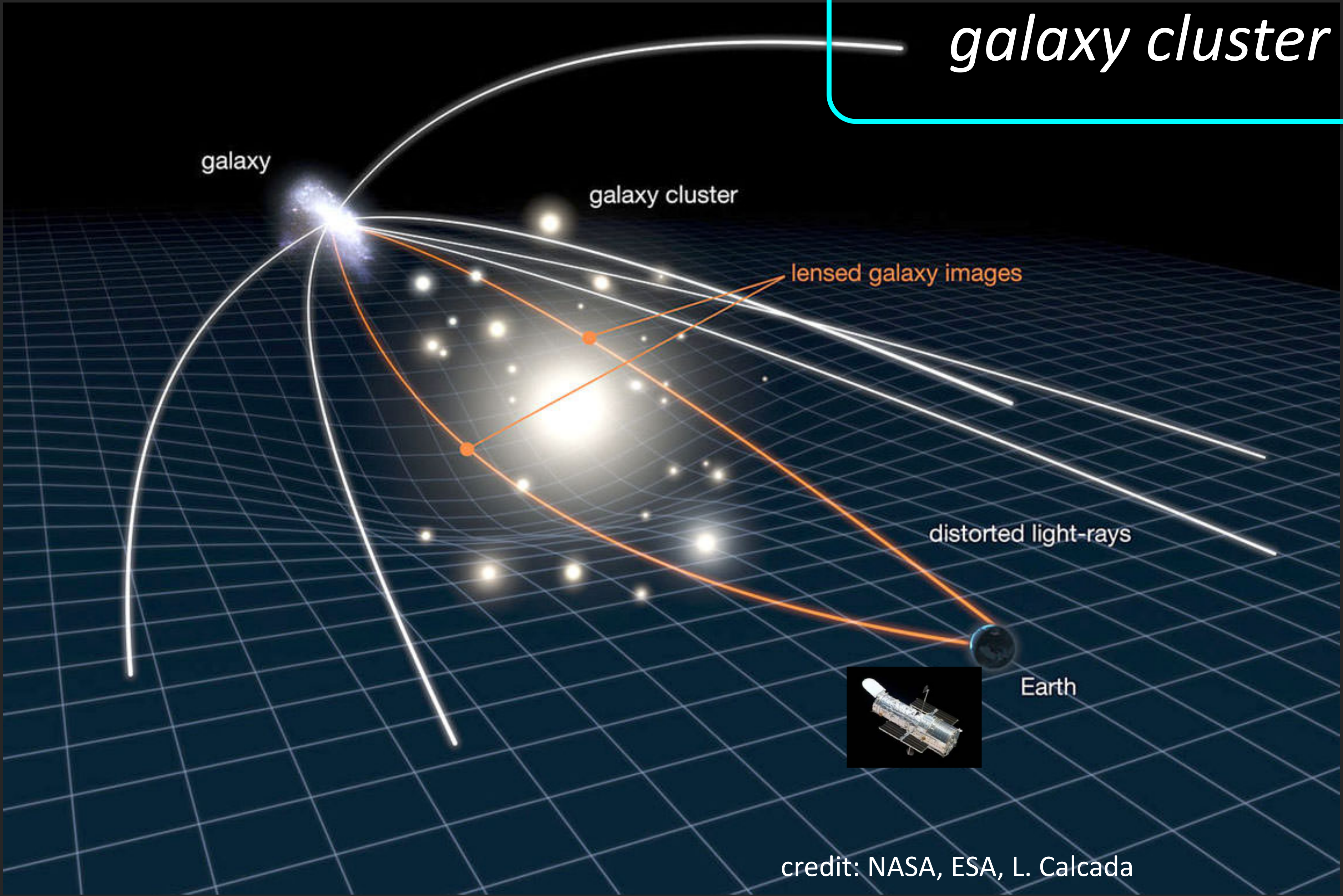
but how do we go fainter than XDF?

Bouwens GDI Oesch+2015

see also McLure+2013, Finkelstein+2015, Bowler+2015, Parsa+2016, Alavi+2016

gdi

galaxy cluster "lenses"



credit: NASA, ESA, L. Calcada

the way to both go faint and to see what faint galaxies really look like…

by combining Hubble with a "cosmic telescope"

gdi

Frontier Fields
HFF

long history of galaxy cluster imaging programs with HST from WFPC2 to ACS to WFC3:

- ACS GTO Team
- CLASH
- HFF
- RELICS
- + others

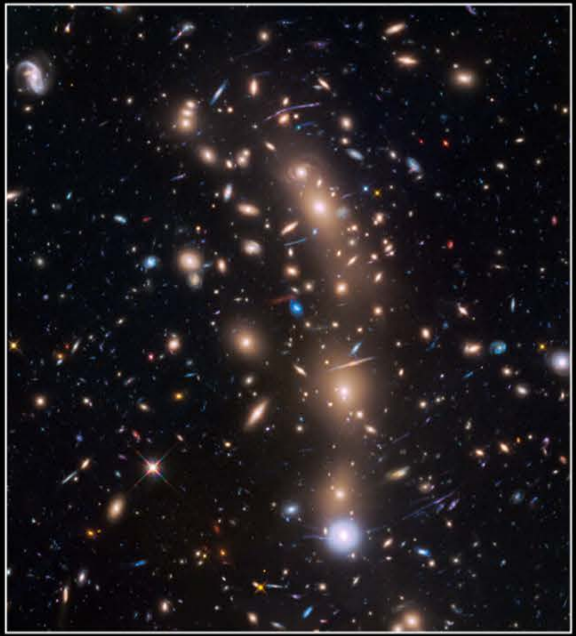
6 clusters + 6 parallel fields

840 orbits of ACS and WFC3/IR images

1000 hours of Spitzer IRAC

figure credit: Jennifer Lotz

Abell 2744

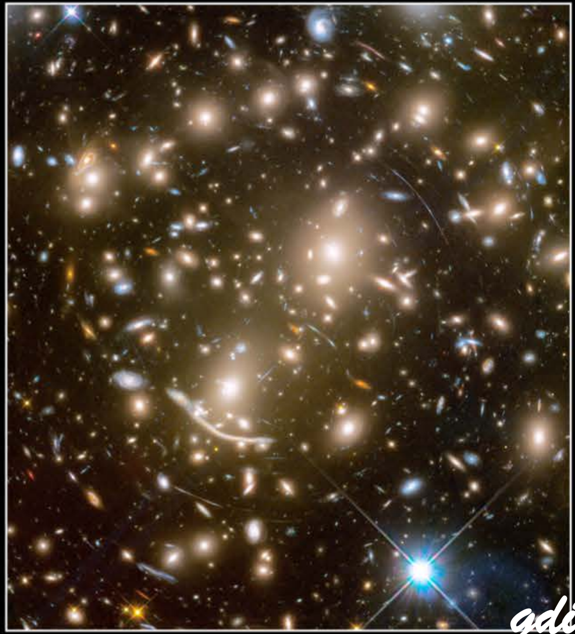


MACSJ0416.1-2403

MACSJ0717.5+3745

MACSJ1149.5+2223

Abell S1063



Abell 370

Frontier Fields HFF

long history of galaxy cluster imaging programs with HST from WFPC2 to ACS to WFC3:

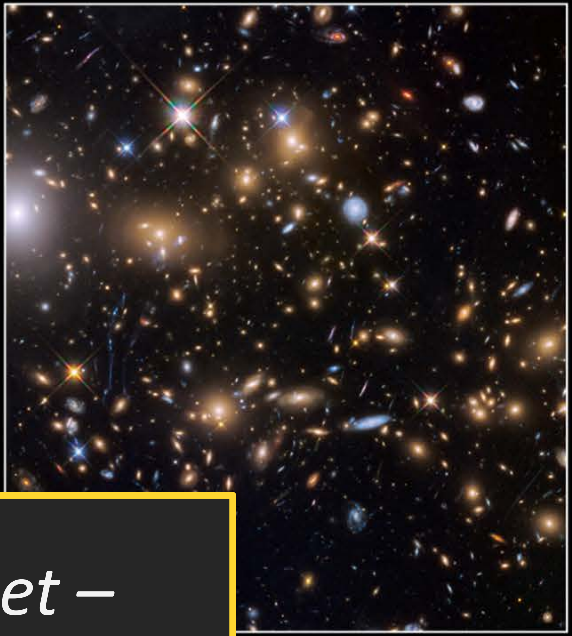
- ACS GTO Team
- CLASH
- HFF
- RELICS
- + others

6 clusters + 6 parallel fields

840 orbits of ACS and WFC3/IR images

1000 hours of Spitzer IRAC

figure credit: Jennifer Lotz

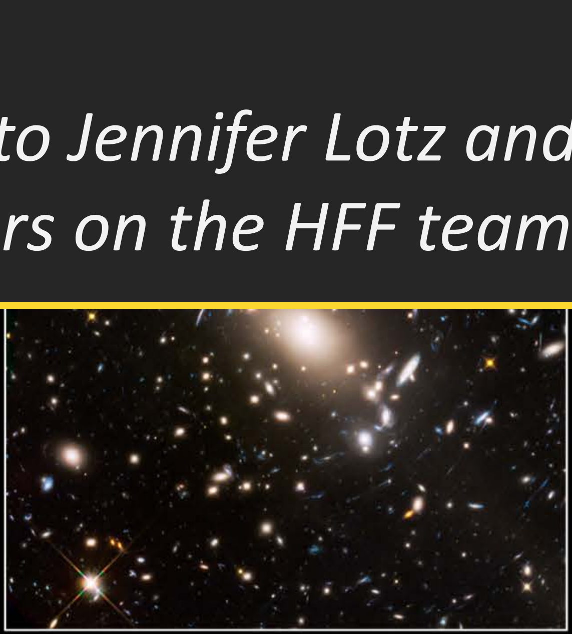


MACSJ0717.5+3745

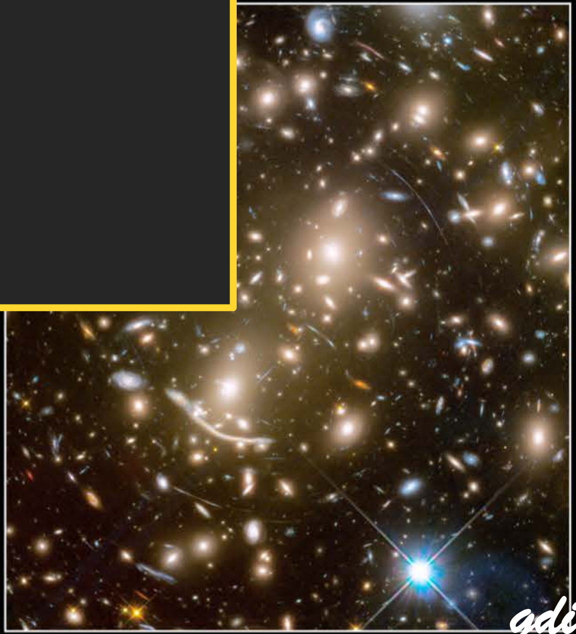
– the HFF is a remarkable dataset – thanks to Matt Mountain

and to Jennifer Lotz and others on the HFF team

MACSJ1149.5+2223



Abell S1063



Abell 370

the challenges of luminosity functions using lensing clusters

☛ model uncertainties at high magnifications ☚

different models yield substantially
different results at high magnification

strongly lensing clusters provide the opportunity to go much fainter than deep fields, but how faint can we reliably push?

gdi

modelling challenging at high magnifications

critical curves for four different models for the HFF cluster Abell 2744 and a source at z~9

from Atek 2017 – adapted from www.stsci.edu/hst/campaigns/frontier-fields/Lensing-Models

gdi

modelling challenging at high magnifications

strongly lensing clusters provide extraordinary opportunities at high magnification to see faint objects, but also extraordinary challenges to measure them accurately

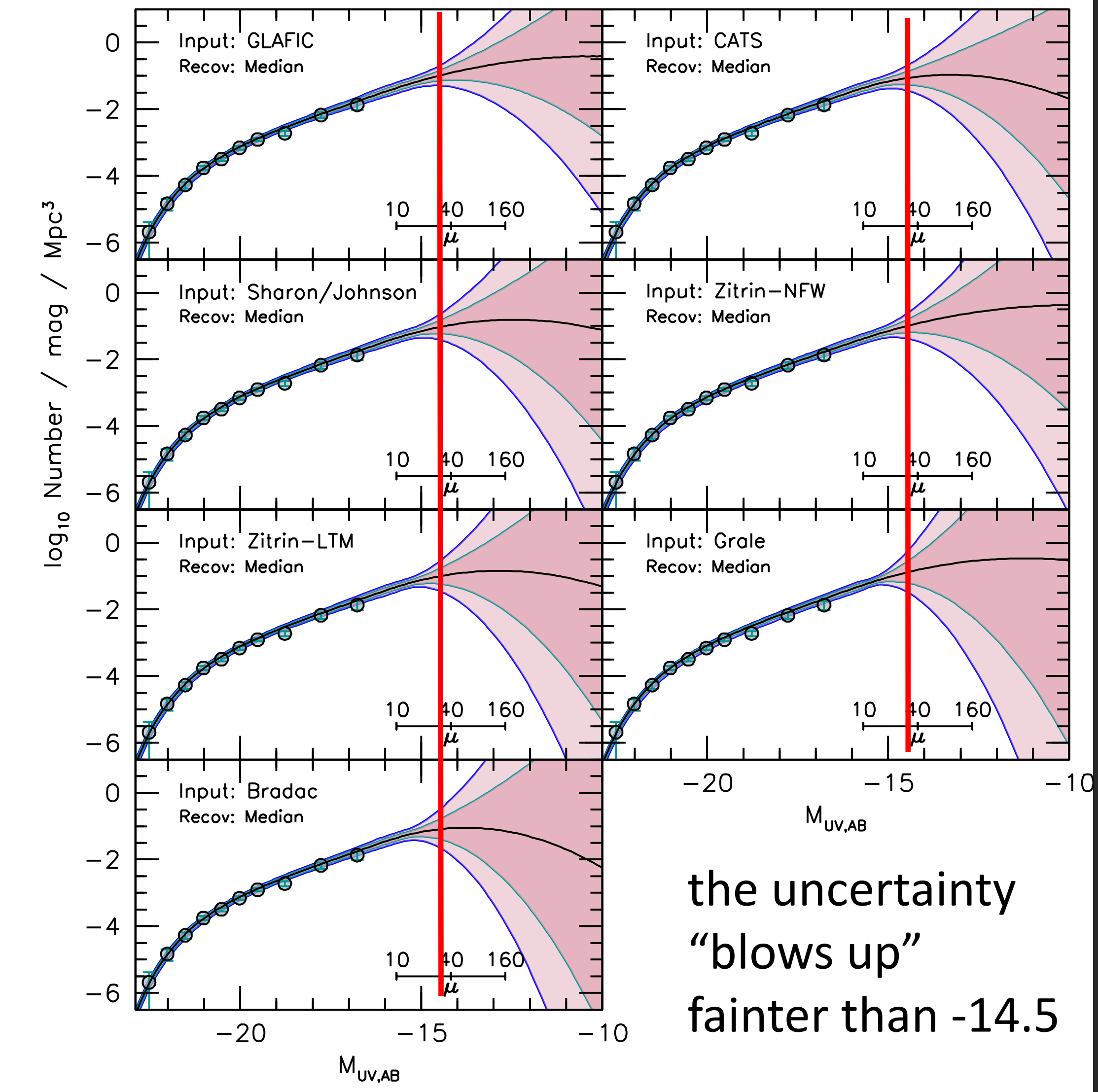
critical curves for four different models for the HFF cluster Abell 2744 and a source at z~9

from Atek 2017 – adapted from www.stsci.edu/hst/campaigns/frontier-fields/Lensing-Models

gdi

limit for reliable LFs from the HFFs

(Hubble Frontier Fields)



the errors in the LF become so large as to make estimates of the LF from the HFF *of limited value* below $M_{UV,AB}$ ~-14.5

forward modelling simulations

model systematics are the limiting factor

Bouwens+2017b

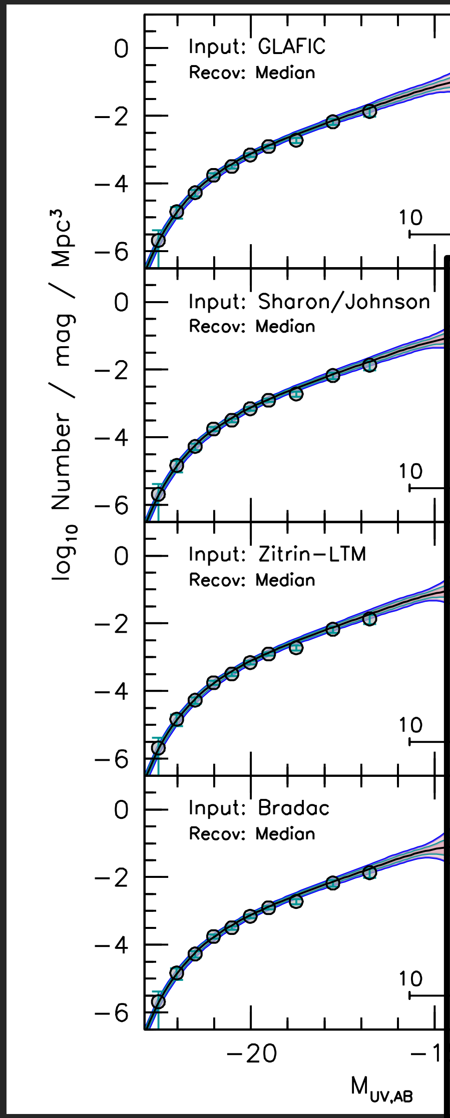
see also Castellano+2015, Atek+2016, 2018, Laporte+2016; Livermore+2017; Ishigaki+2017; Yue+2018

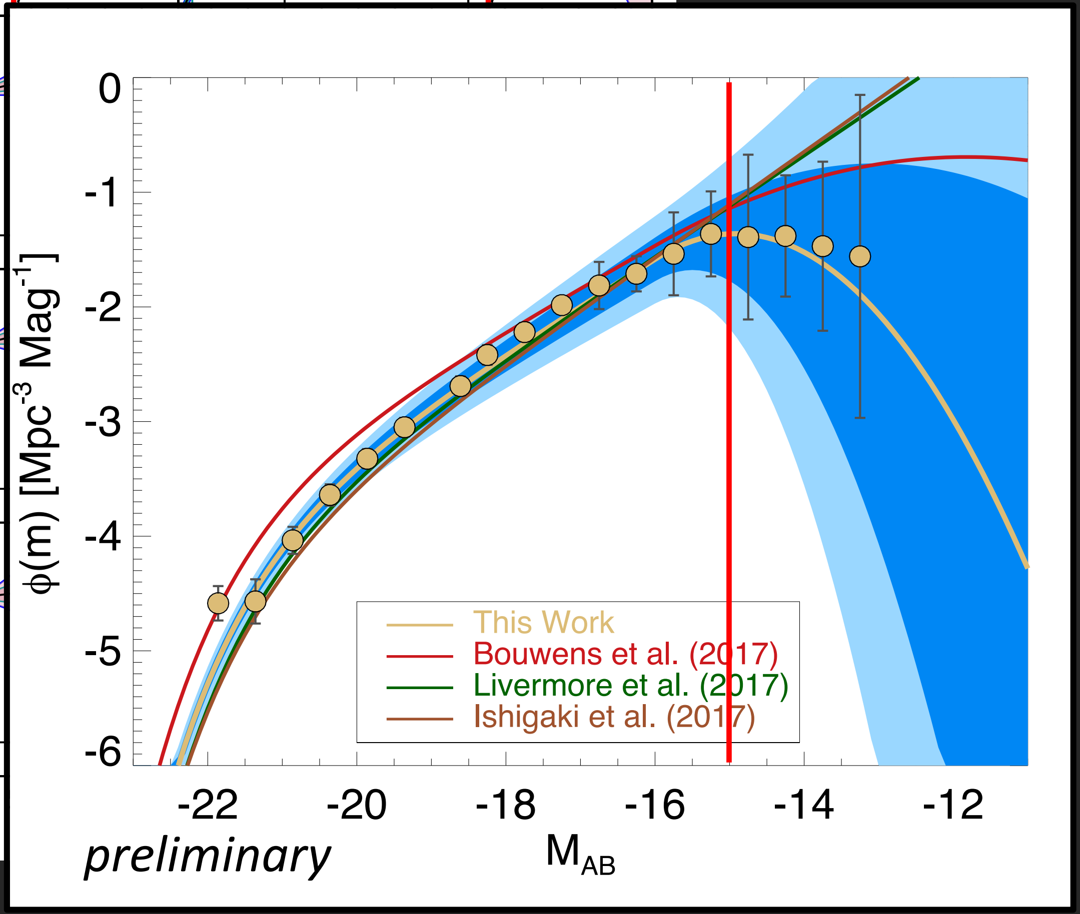
gdi

limit for reliable LFs from the HFFs

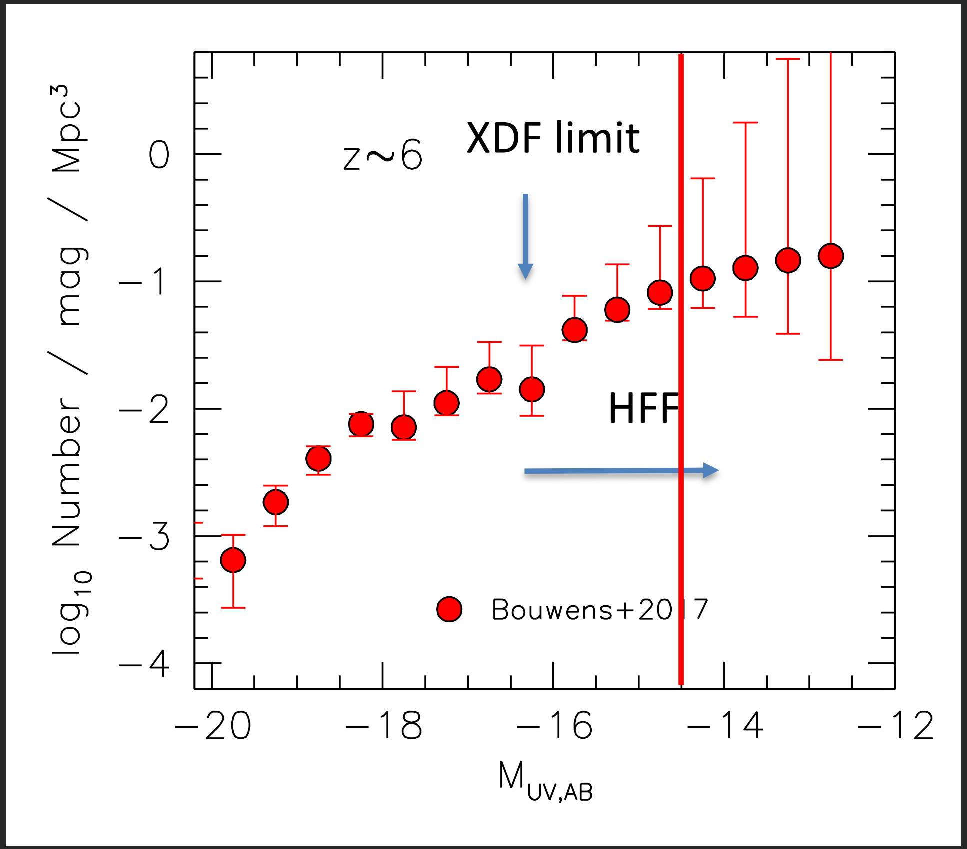
(Hubble Frontier Fields)

HFF gained us 2 mags





Atek+2018



Bouwens+2017b

large systematic errors and large uncertainties below ~-14.5 to -15 mag – confirmed by Atek+2018

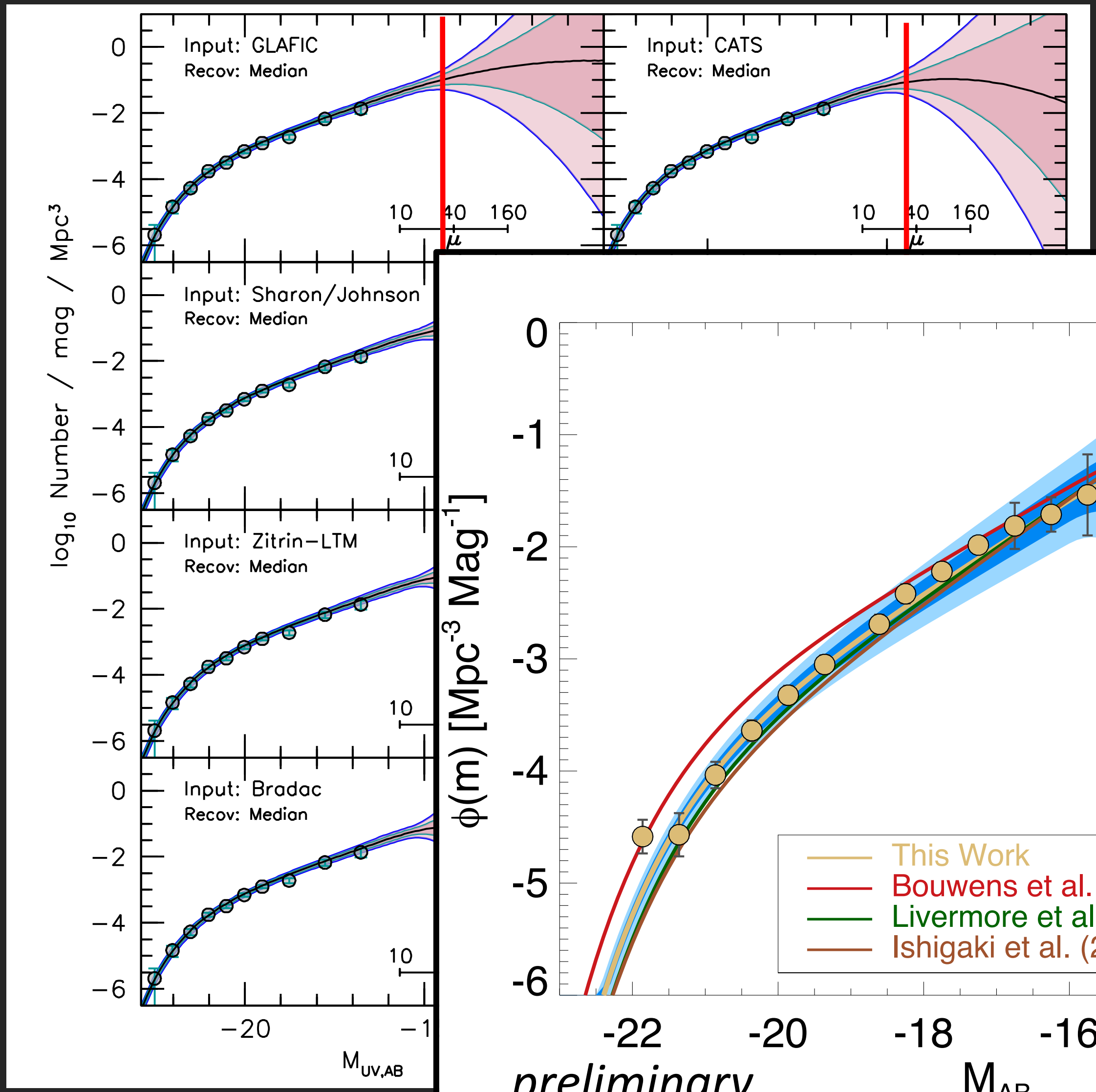
see also Castellano+2015, Atek+2016, 2018, Laporte+2016; Livermore+2017; Ishigaki+2017; Yue+2018

gdi

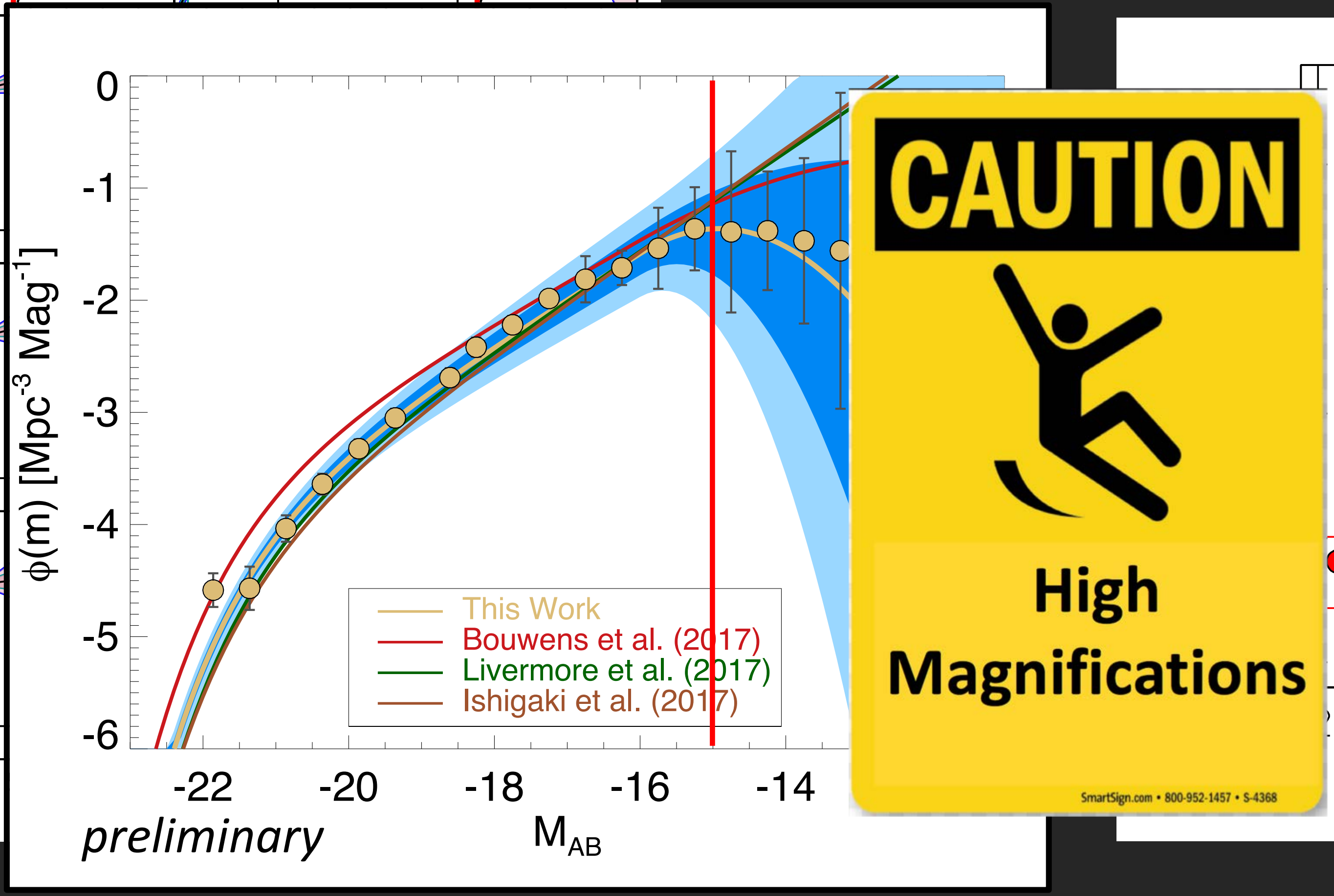
limit for reliable LFs from the HFFs

(Hubble Frontier Fields)

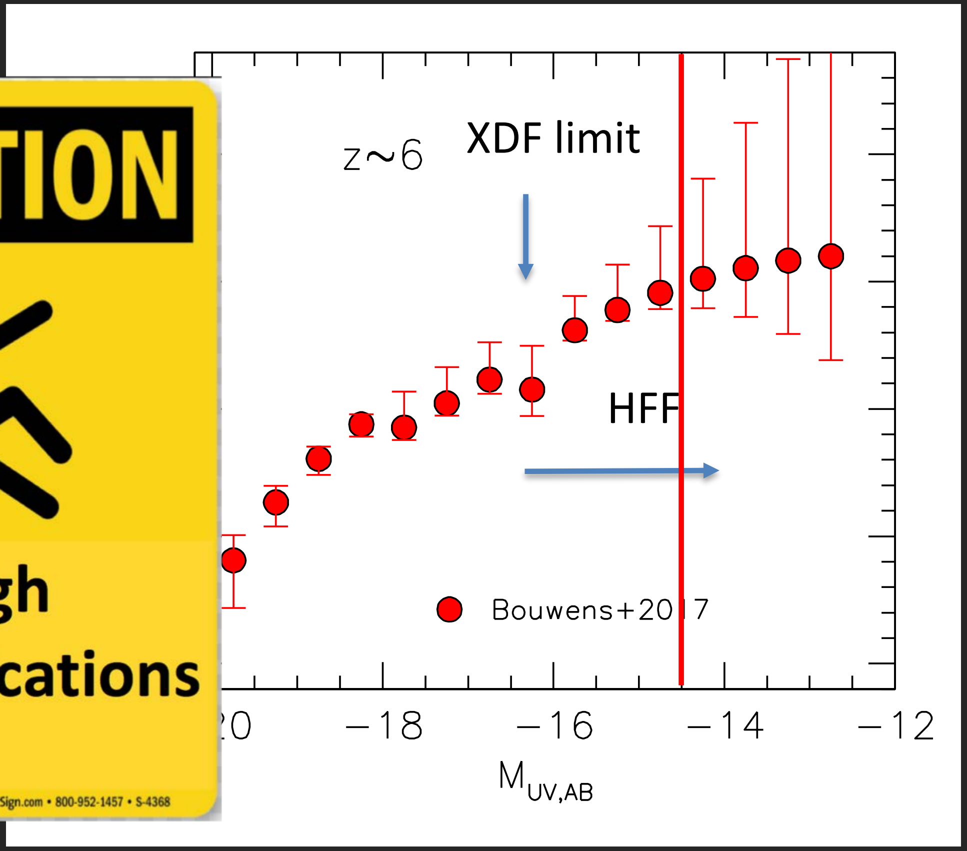
HFF gained us 2 mags



Bouwens+2017b



Atek+2018



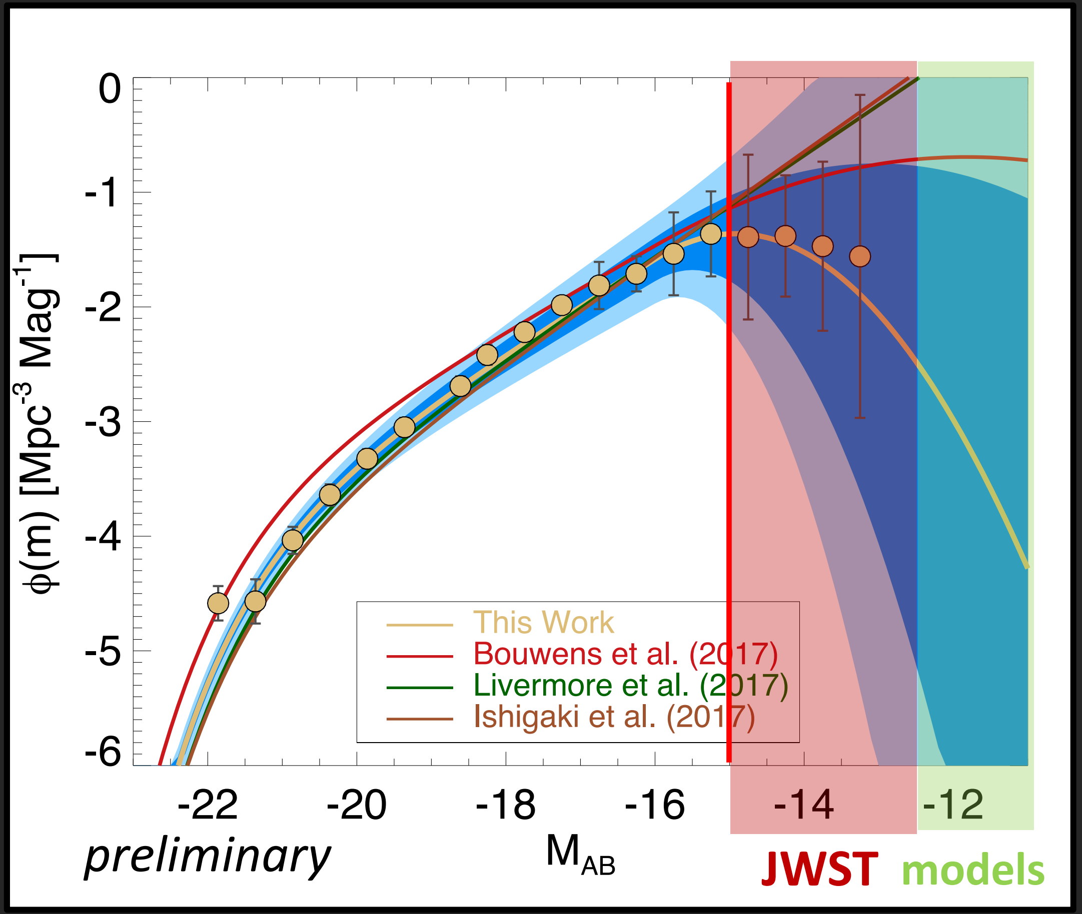
large systematic errors and large uncertainties below ~-14.5 to -15 mag – confirmed by Atek+2018

see also Castellano+2015, Atek+2016, 2018, Laporte+2016; Livermore+2017; Ishigaki+2017; Yue+2018

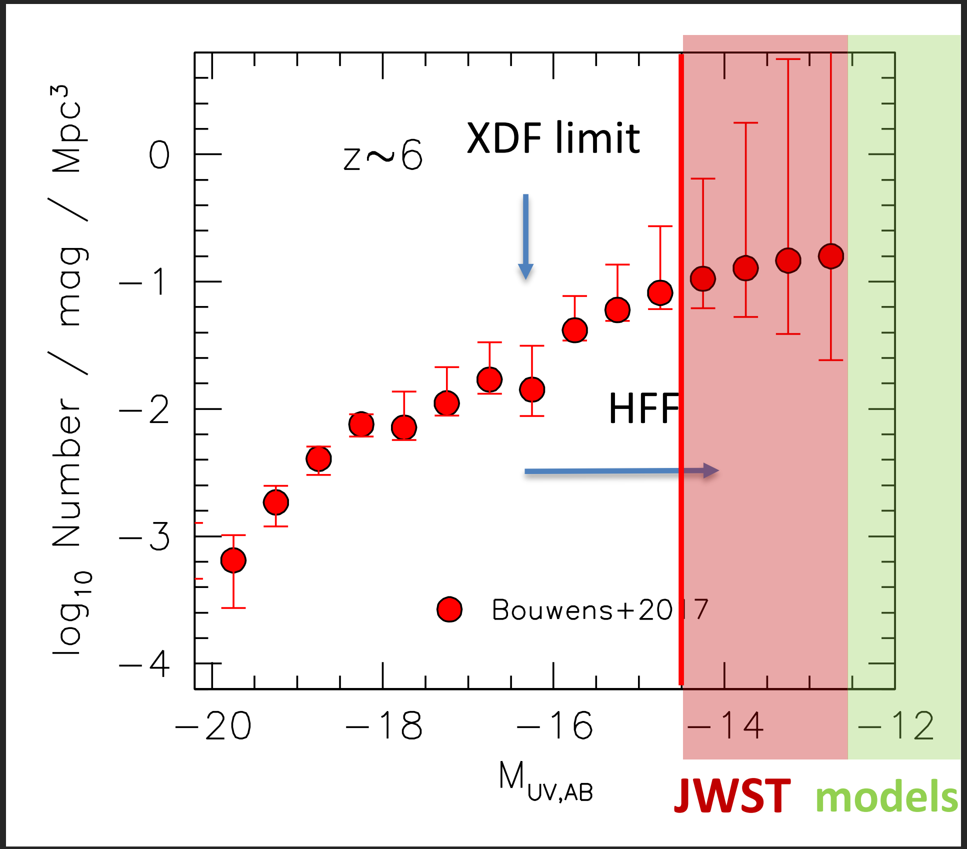
gdi

limit for reliable LFs from the HFFs

(Hubble Frontier Fields)



Atek+2018



Bouwens+2017b

JWST imaging (plus better constrained models) should allow us to see the turnover

see also Castellano+2015, Atek+2016, 2018, Laporte+2016; Livermore+2017; Ishigaki+2017; Yue+2018

gdi

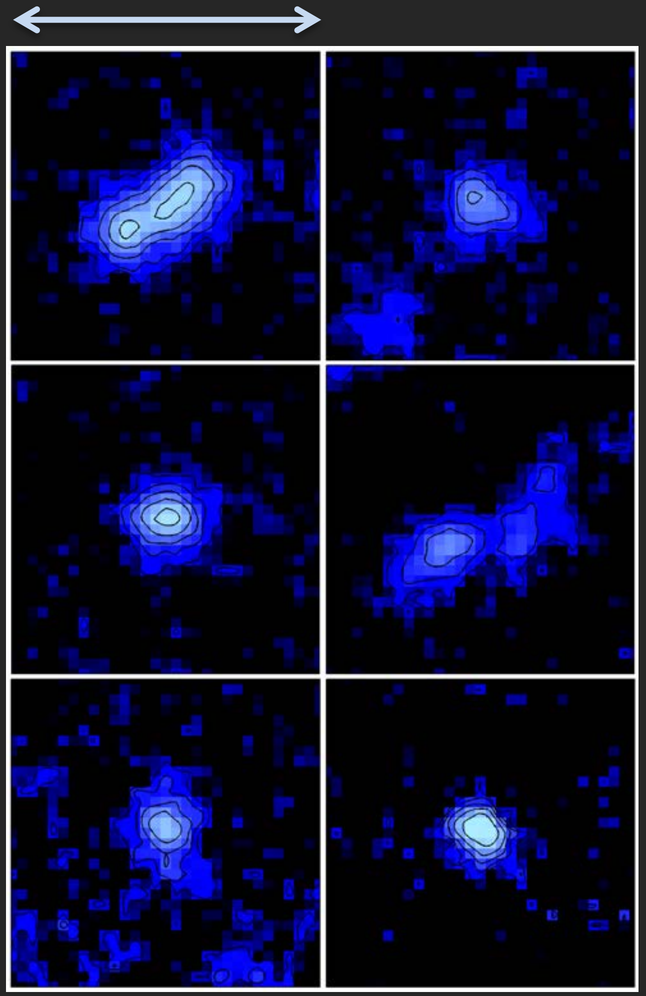

the sizes of galaxies -- how large are early galaxies
gdi

galaxies in the first billion years

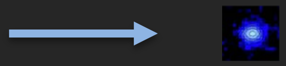
large bright z~6-7 galaxies

a large galaxy now

1.8” (~10 kpc)



small faint distant galaxies



really tiny!

most galaxies in the first billion years have recently been measured for the first time to be surprisingly small!

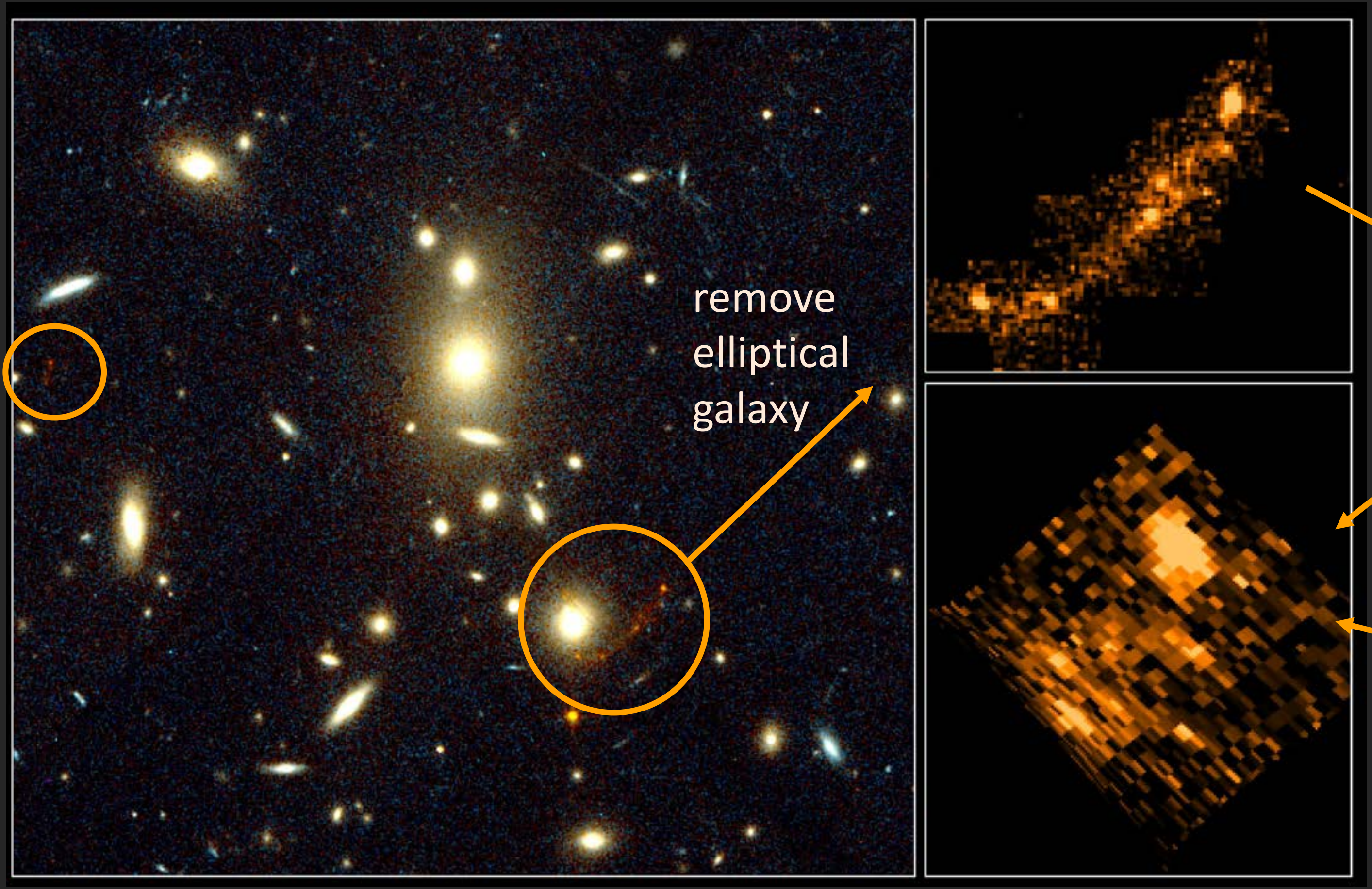
~size of the Hubble point spread function

gdi

a remarkable fold arc in CL1358

cluster of galaxies CL1358 at z=0.33
magnifies faint z=4.92 background galaxy

get ~20X magnified
image of distant galaxy
12.5 million years ago



WFPC2 image

CL1358-G1 at z=4.92

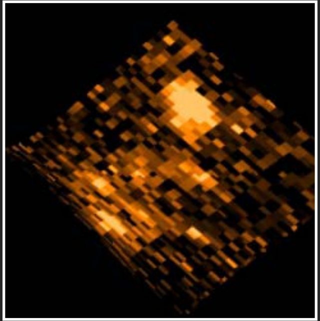
model gravity in cluster
remove distortions
recover source plane image

50% of star formation (>50%)
in the bright "blob" r_e ~130pc

SFR around 40 $M_\odot$/yr

Franx GDI+1997

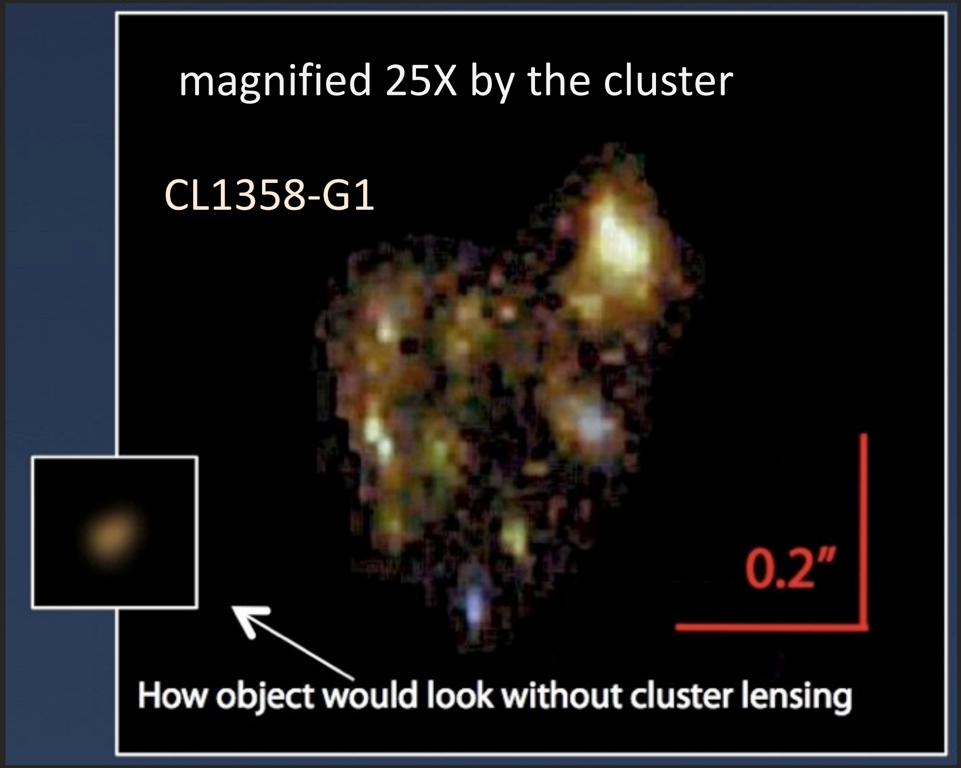
unique insight into the structure of a high redshift galaxy

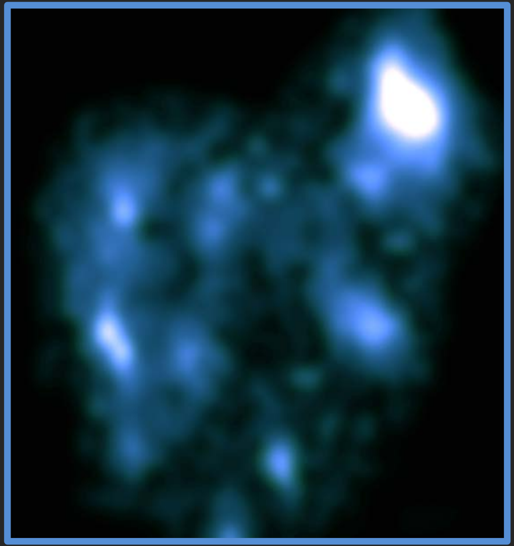


1996 WFPC2 image

found in 1996 – still the best magnified image we have for a galaxy in the first 2 billion years

2004 image from Hubble's Advanced Camera ACS





CL1358-G1 probably looks more like this!

- very rare to see such details
- star-forming regions at high redshift are very small

Zitrin+2011

gdi

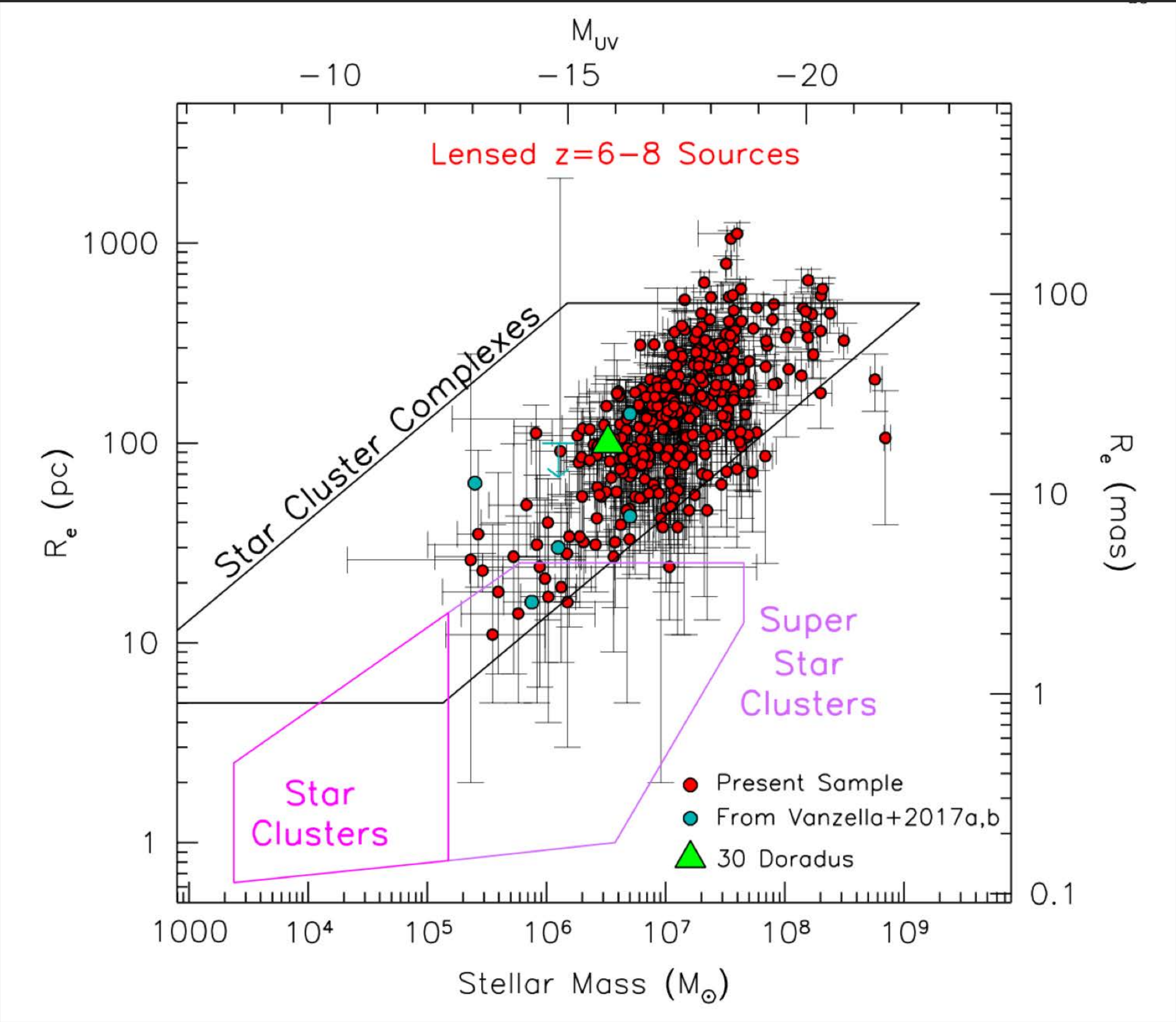
z~6-8 galaxy size comparison to star-forming regions

sizes of 307 z~6-8 galaxies in HFFs

compared to star-forming clusters or complexes and super star clusters

☞ observed sizes of z~6-8 galaxies are similar to lower redshift star-forming complexes – note 30 Doradus ☜

Bouwens+2017c



see also Vanzella+2017a,b; Laporte+2016; Kawamata+2015,2017

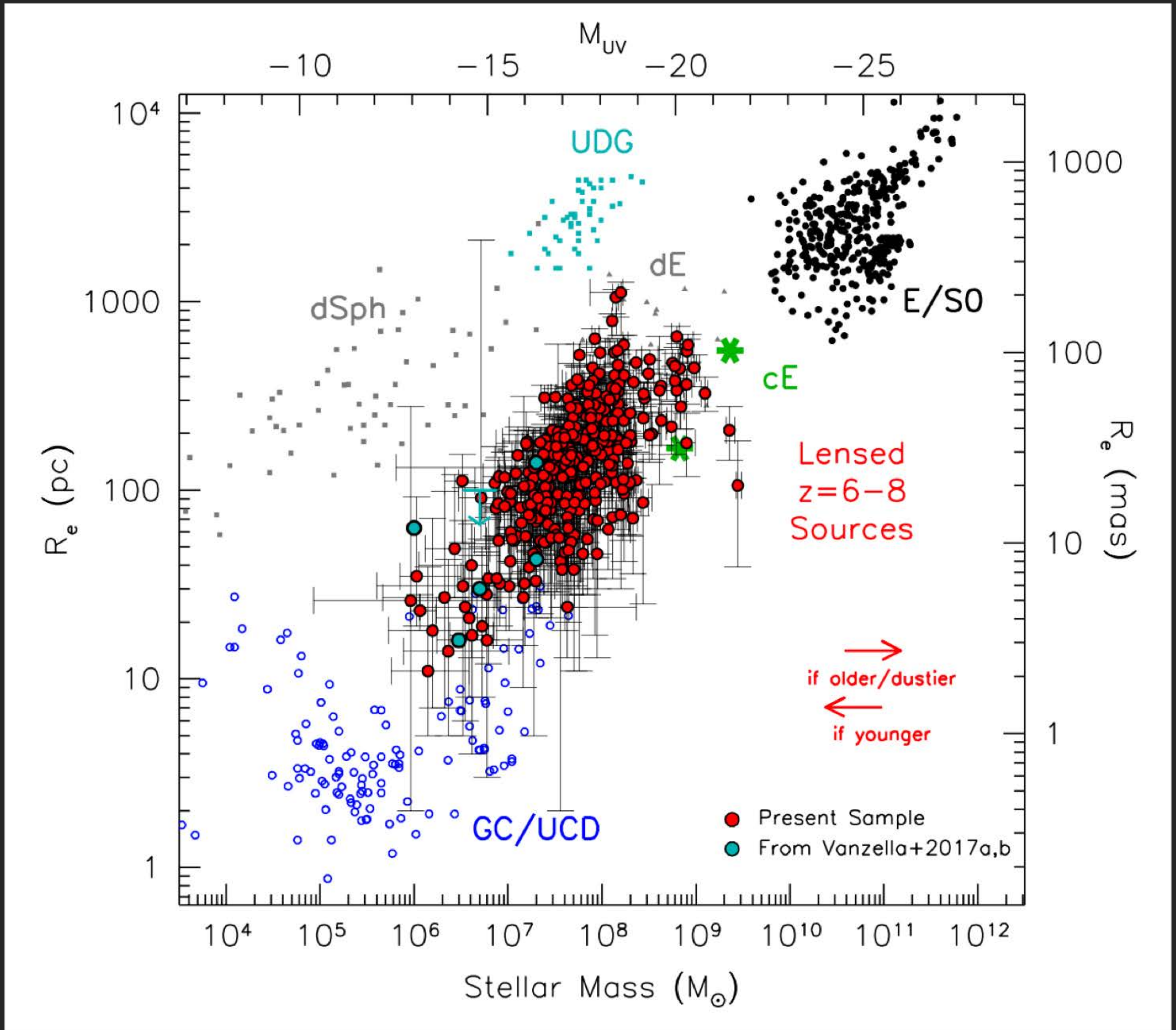
gdi

z~6-8 galaxy size comparison to nearby evolved objects

sizes of z~6-8 galaxies in HFFs compared to nearby evolved objects

local objects from Norris+2014 (see also Brodie+2011)

☞ could we be seeing some globular clusters forming at very high redshift? ☜



Bouwens+2017c

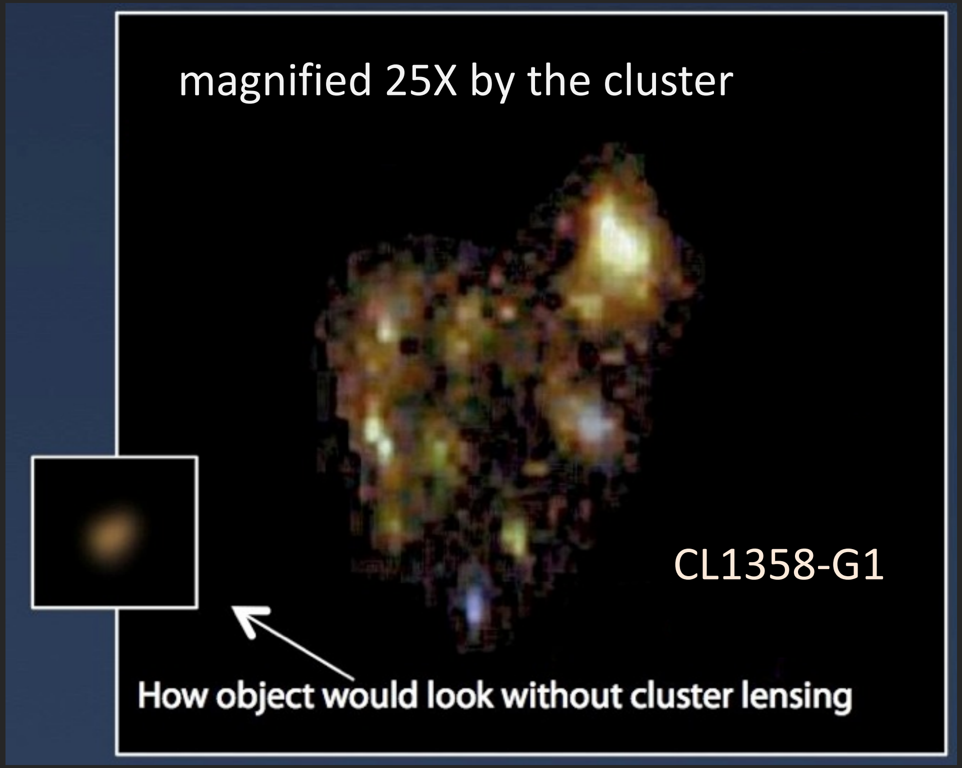
see also Vanzella+2017a,b; Laporte+2016; Kawamata+2015,2017

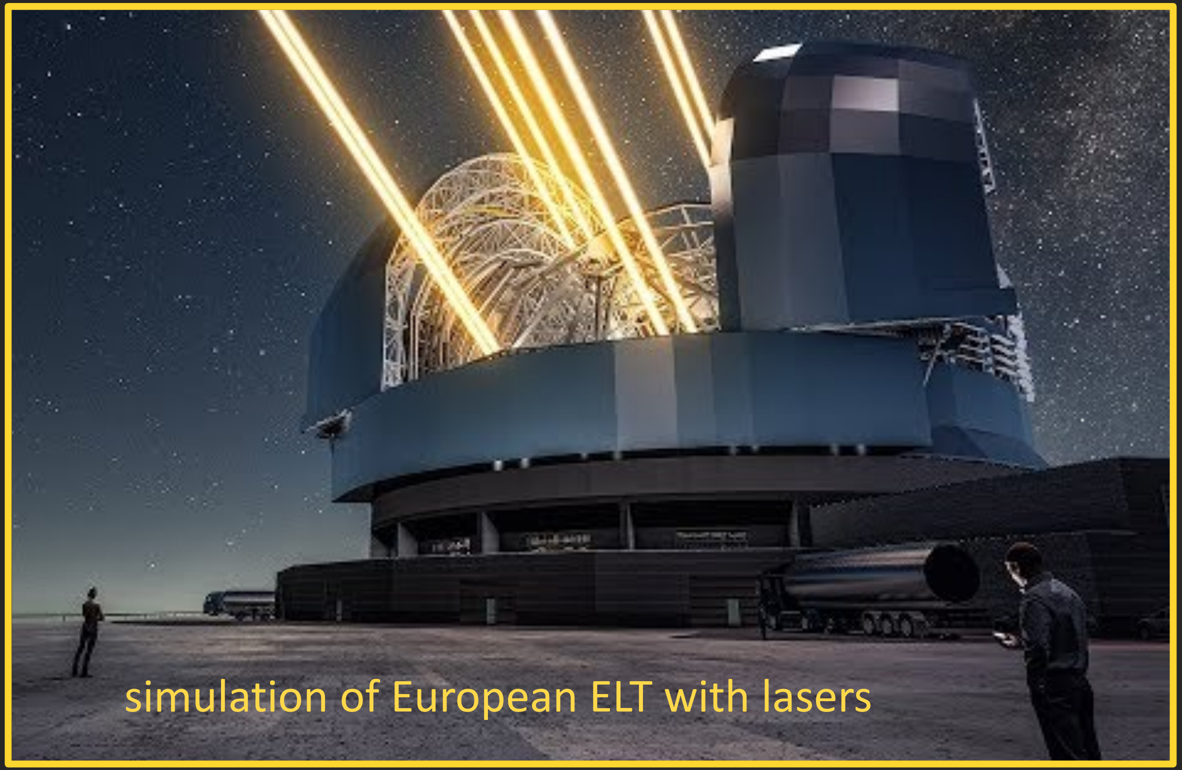
gdi

how will we find more?

>100 clusters have been searched – CL1358G1 is still the best and only one at high redshift

30-40 m telescopes will give <100 pc resolution from lasers and adaptive optics





30-40 m ELT with adaptive optics needed to measure the sizes of star-forming regions in a large sample of early galaxies

ELT – Extremely Large Telescope

gdi

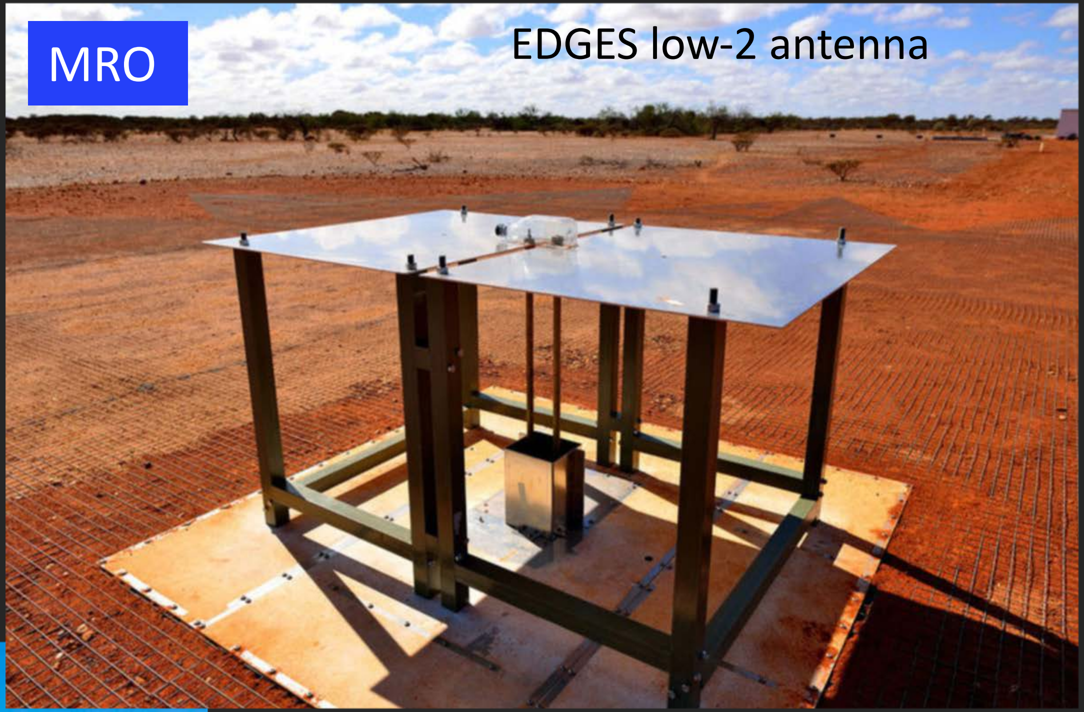
our first indication of when the “first stars” appeared

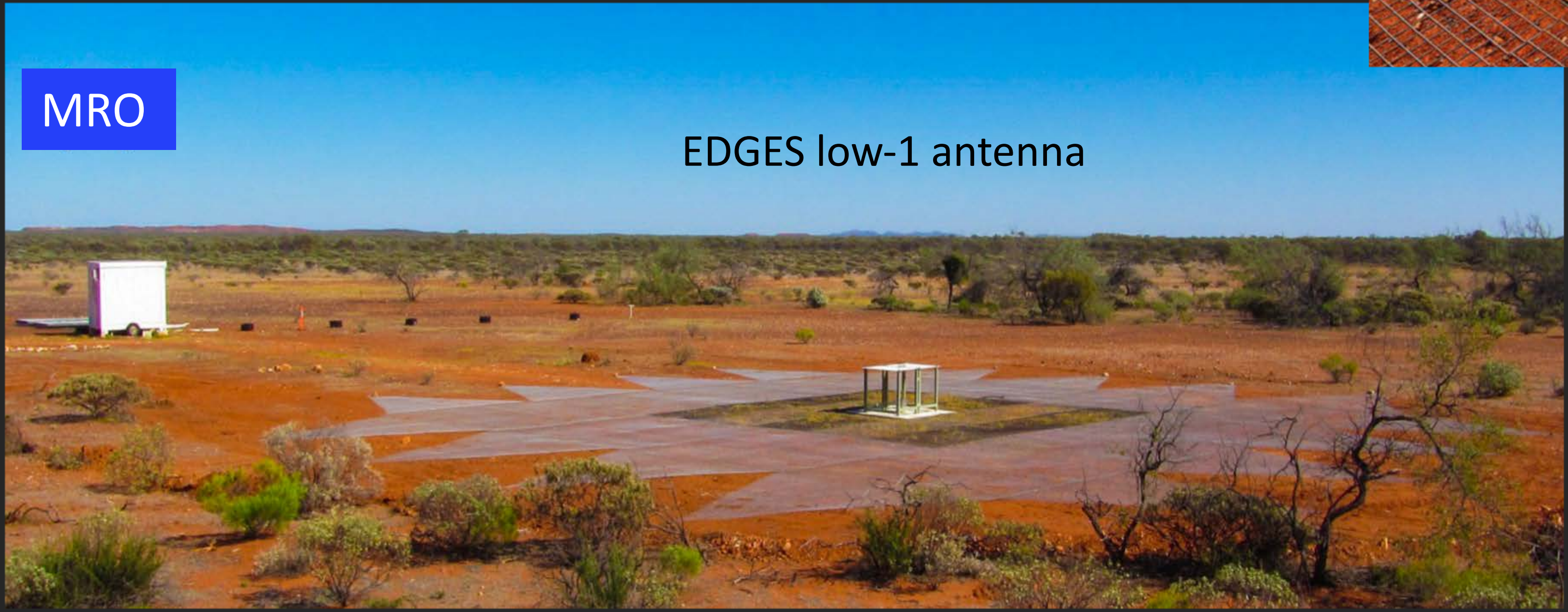
gdi

first evidence for when the "first stars" started to shine brightly

found by these funny-looking (and small) radio antennae in the desert of Western Australia….

Experiment to Detect the Global Epoch of Reionization Signature





Murchison Radio-astronomy Observatory (MRO) in Western Australia

NEW RESULT

published March 01
Nature

NSF

Bowman, Rogers, Monsalve, Mozdzen & Mahesh

National Science Foundation

gdi

first evidence for when the "first stars" started to shine brightly

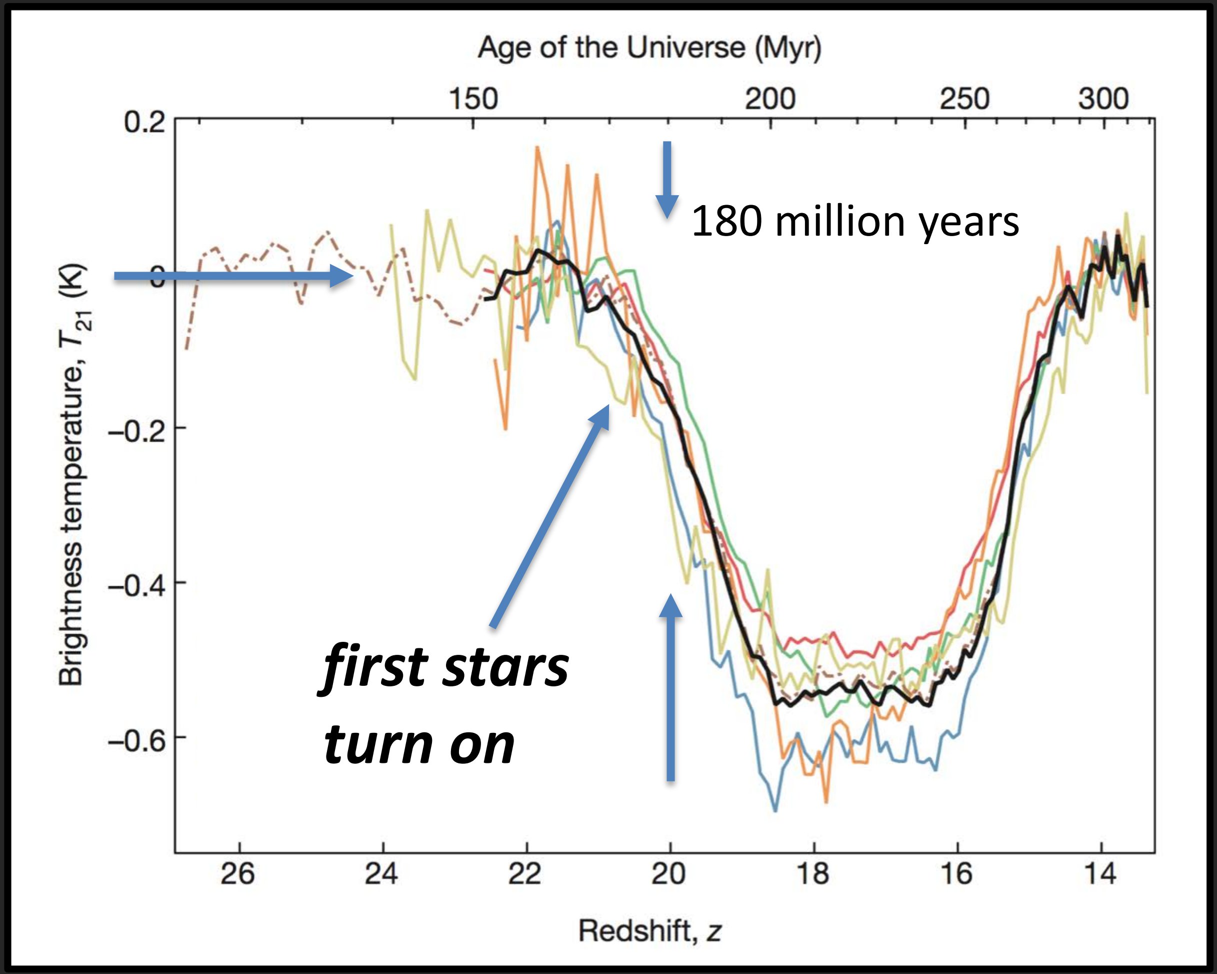
cosmic microwave background

NEW RESULT

published March 01 Nature

is this correct?

confirmation?



first stars become prominent at redshift z~20 (~180 million years)

Bowman+2018

gdi

first evidence for when the "first stars" started to shine brightly

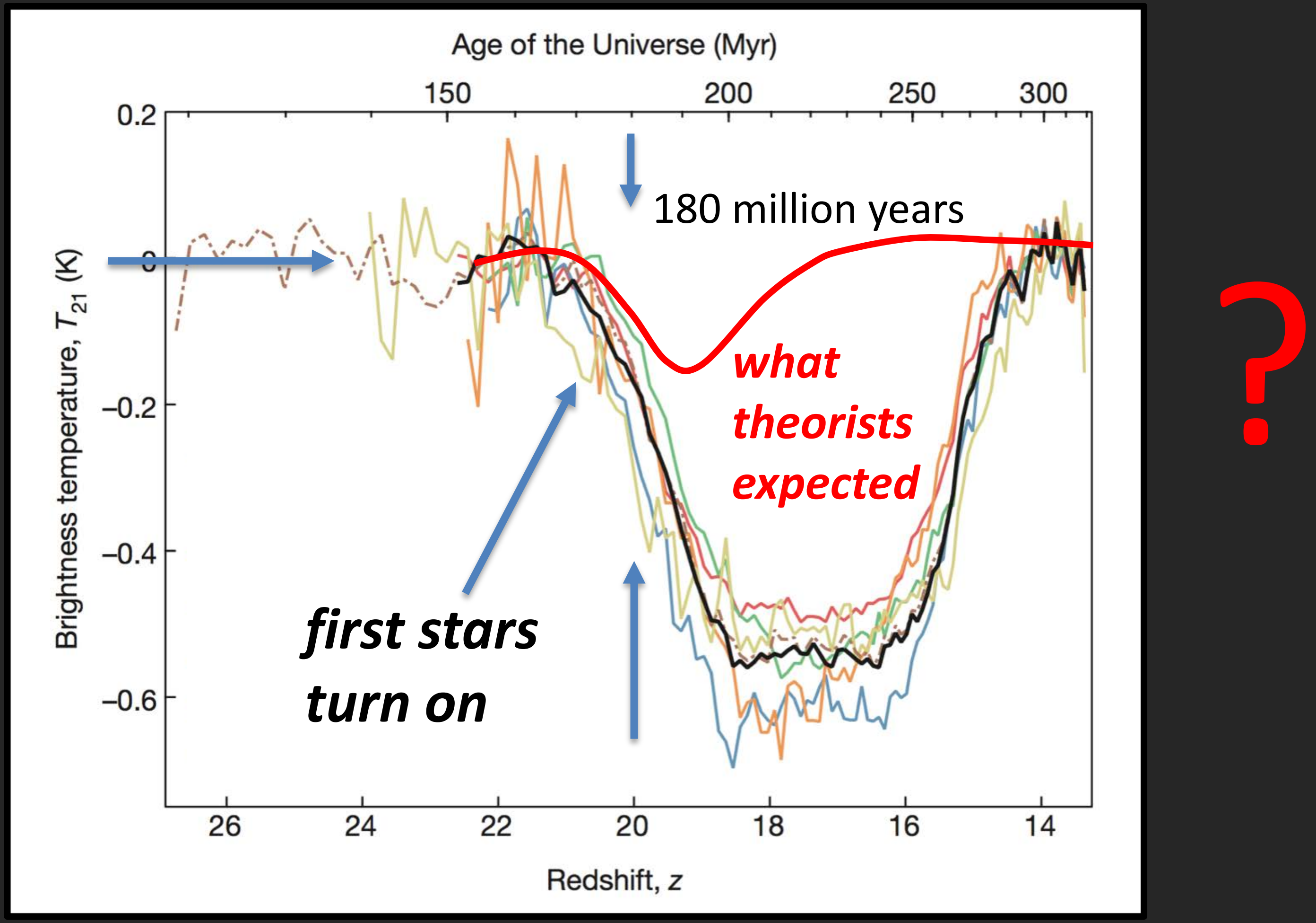
cosmic microwave background

NEW RESULT

published March 01
Nature

is this correct?

confirmation?



first stars become prominent at redshift z~20 (~180 million years)

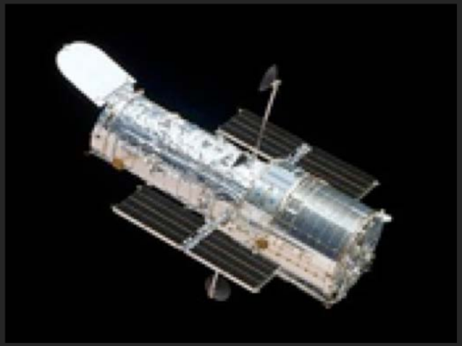
Bowman+2018

gdi

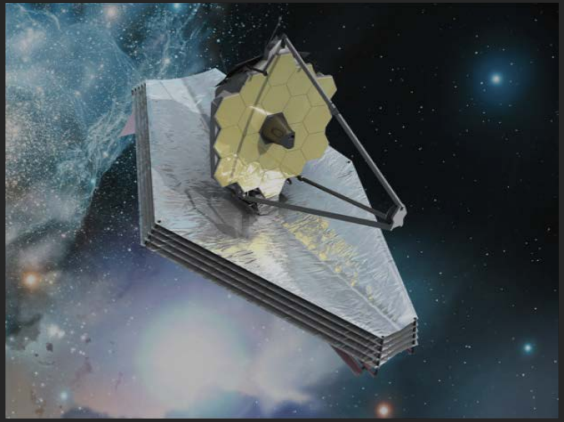
what do we know about the first galaxies?

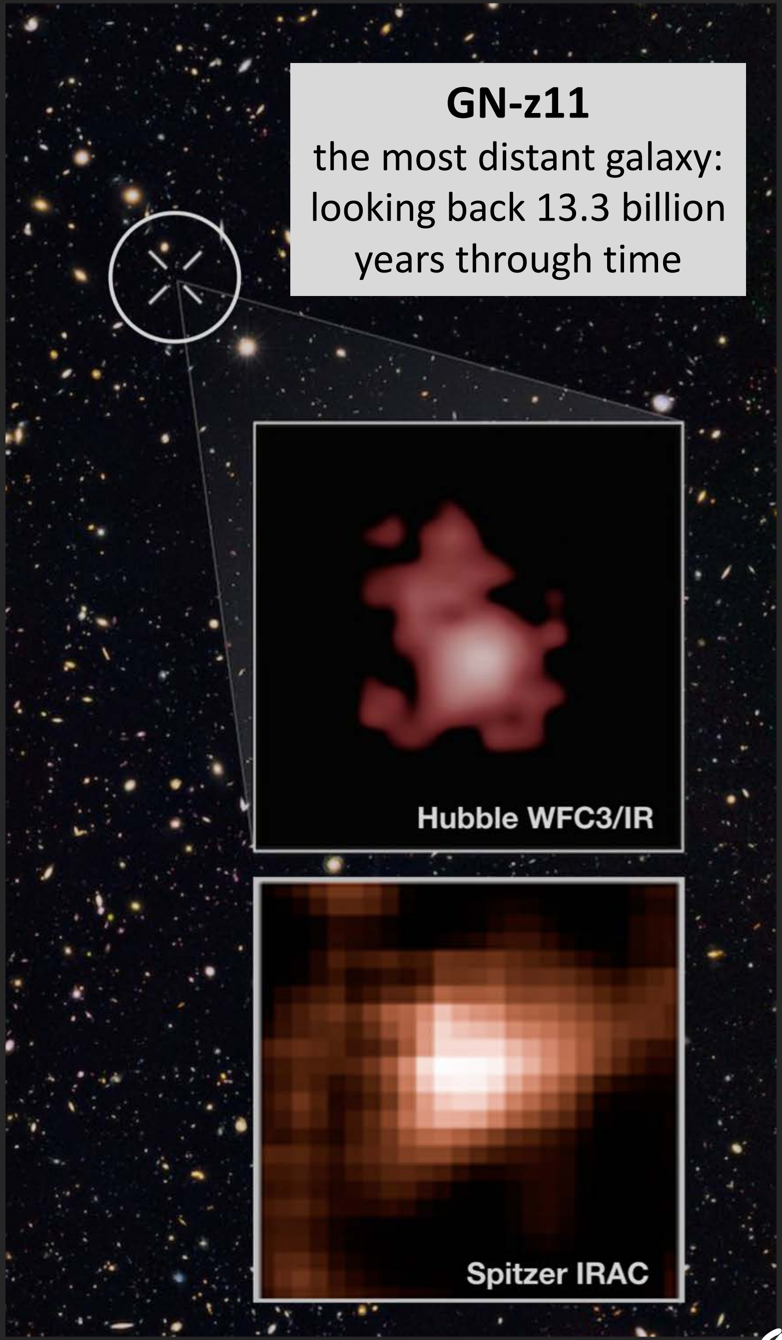
the first galaxies must be earlier than GN-z11

i.e., earlier than 400 million years
but not by much – maybe 100-200 million years?



Hubble and Spitzer have been reaching into JWST territory!

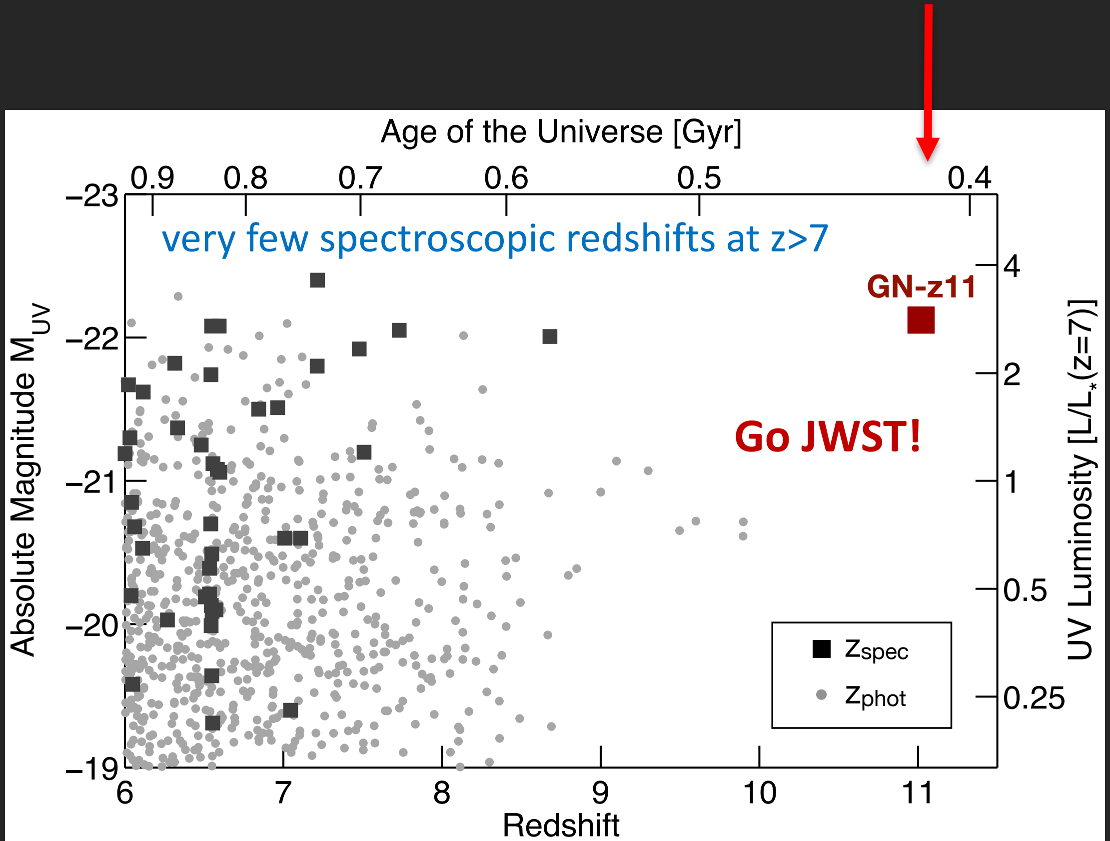


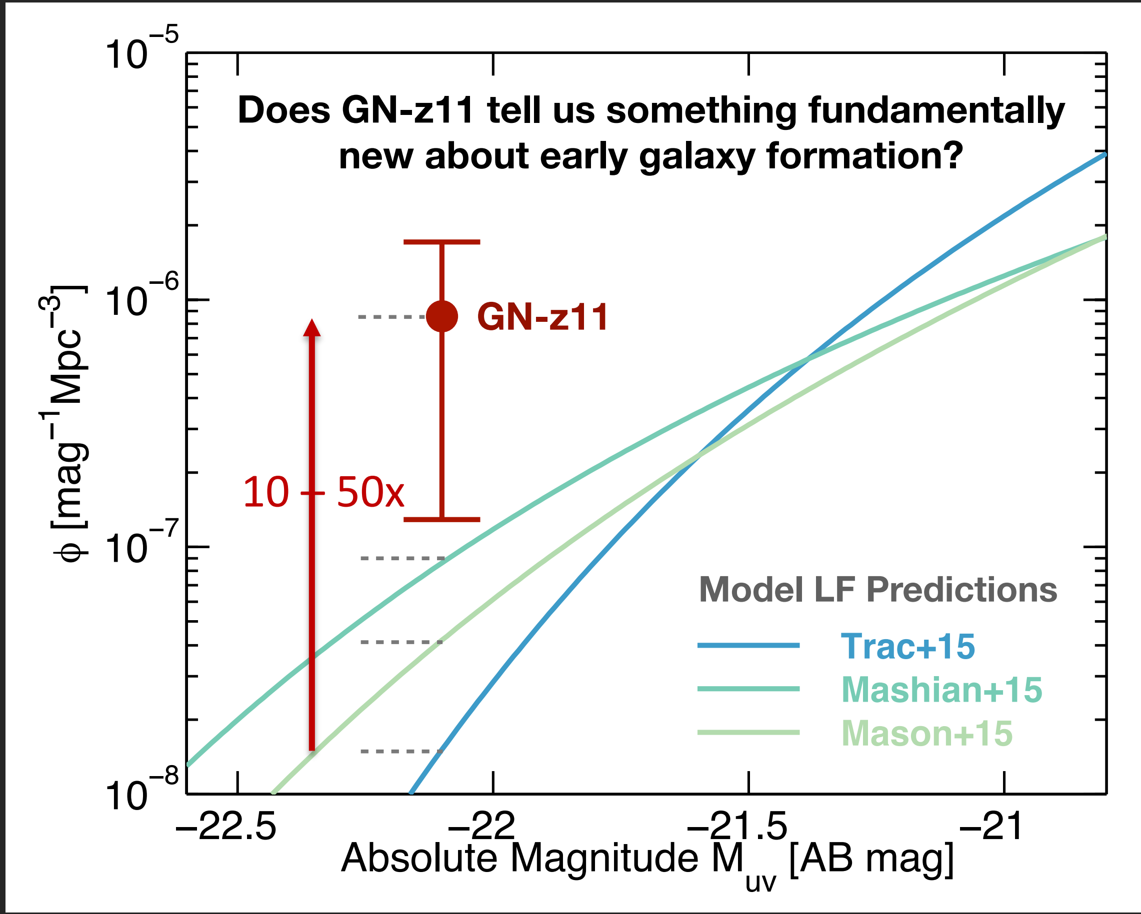


gdi

GN-z11 – the most distant galaxy found to date

surprising discovery of GN-z11





- detection of GN-z11 in *existing data* is unexpected, given current models
- expected to require 10-100x larger areas to find one z~11 galaxy as bright as GN-z11

Oesch+2016

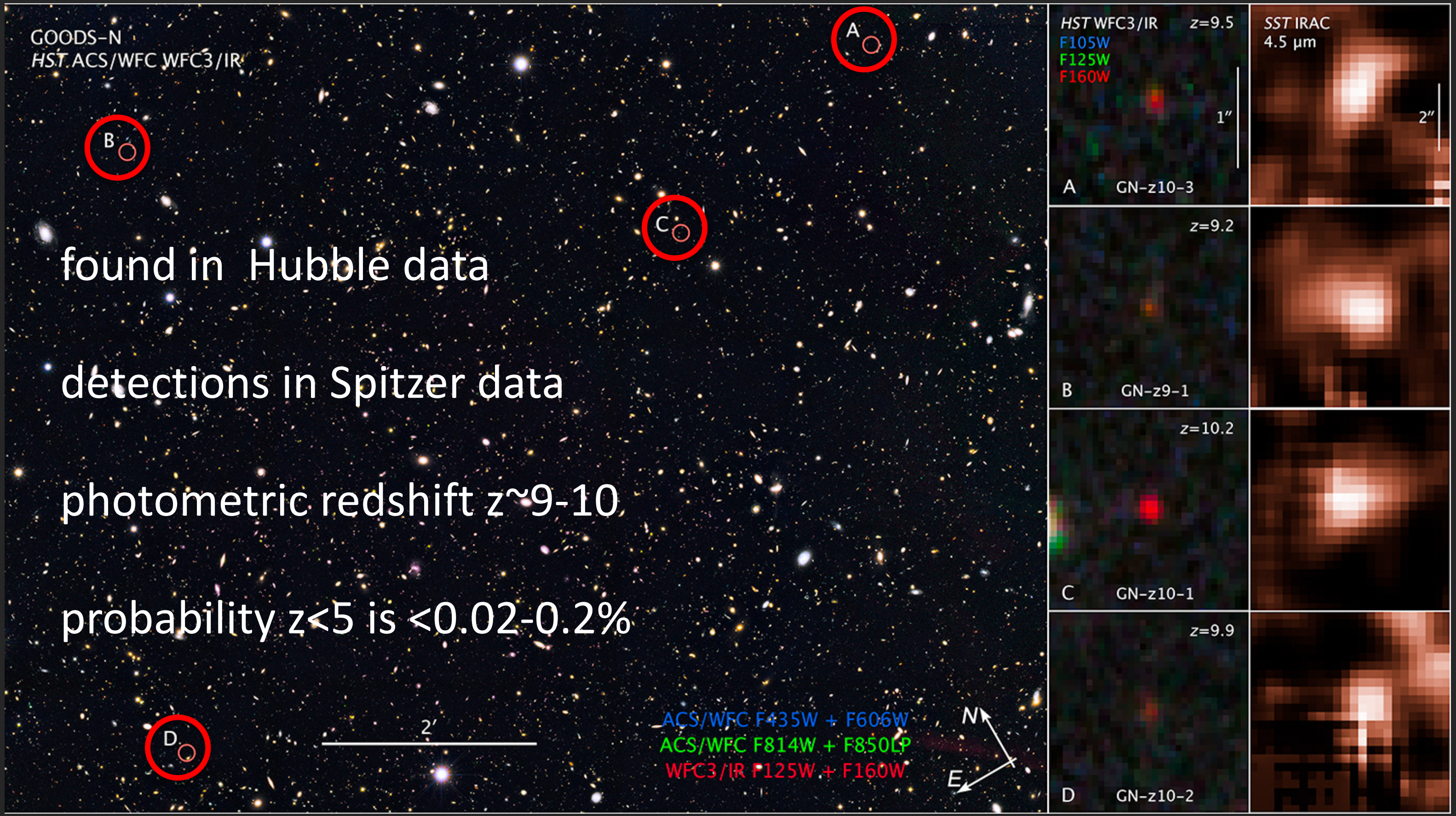
gdi

the highest redshift galaxies

galaxies at z~10 (480 Myr)

gdi

some very luminous galaxy candidates at redshift z~9-10



found in Hubble data

detections in Spitzer data

photometric redshift z~9-10

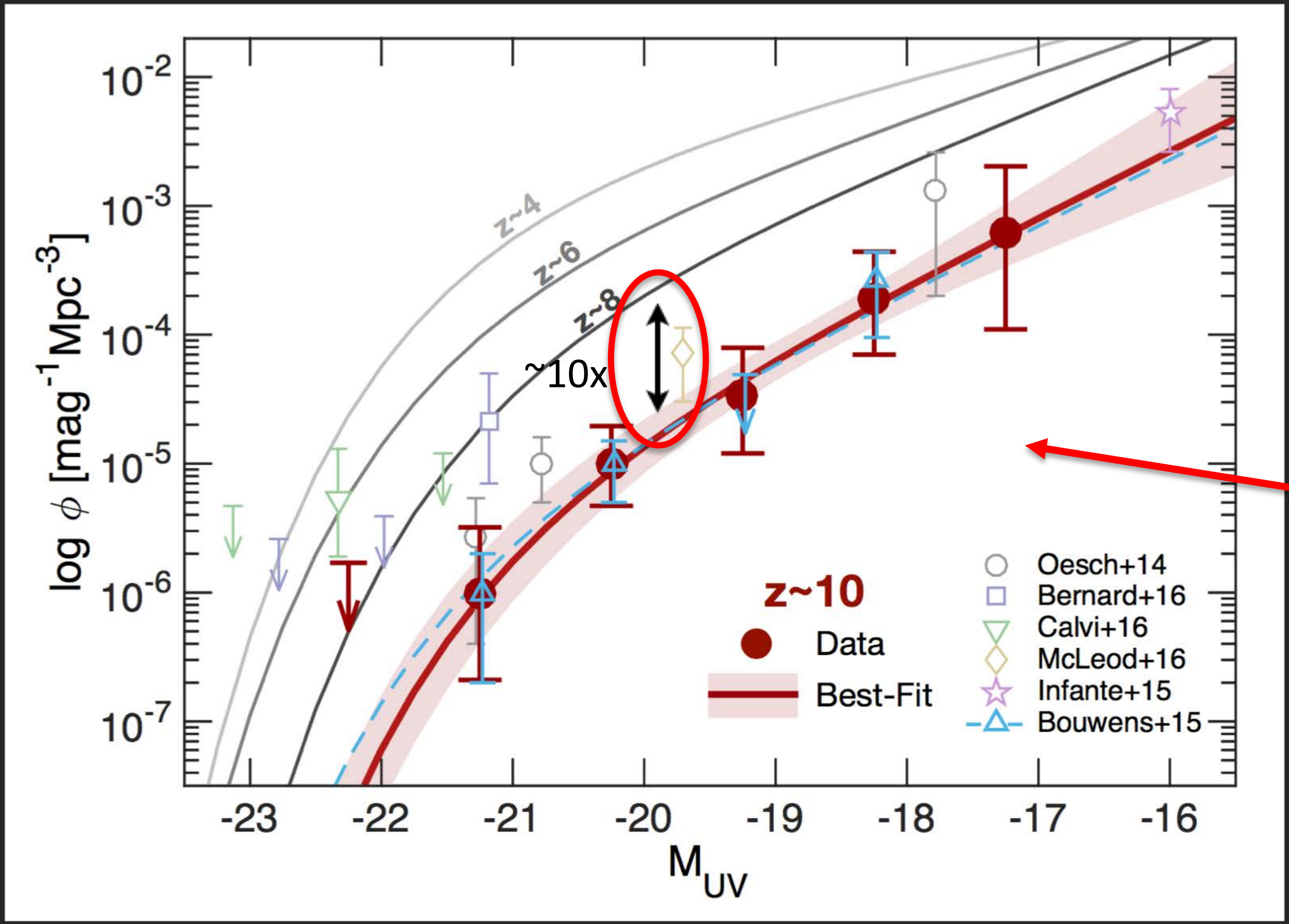
probability z<5 is <0.02-0.2%

GN-z11

Hubble Spitzer

Oesch Bouwens GDI+2014

the luminosity function at z~10



z~10 galaxies are hard to find!

8 years of WFC3/IR imaging

only 9 galaxies in the major Legacy fields: HUDF/XDF + CANDELS/GOODS + HFF

note the change of an order of magnitude between z~8 and z~10

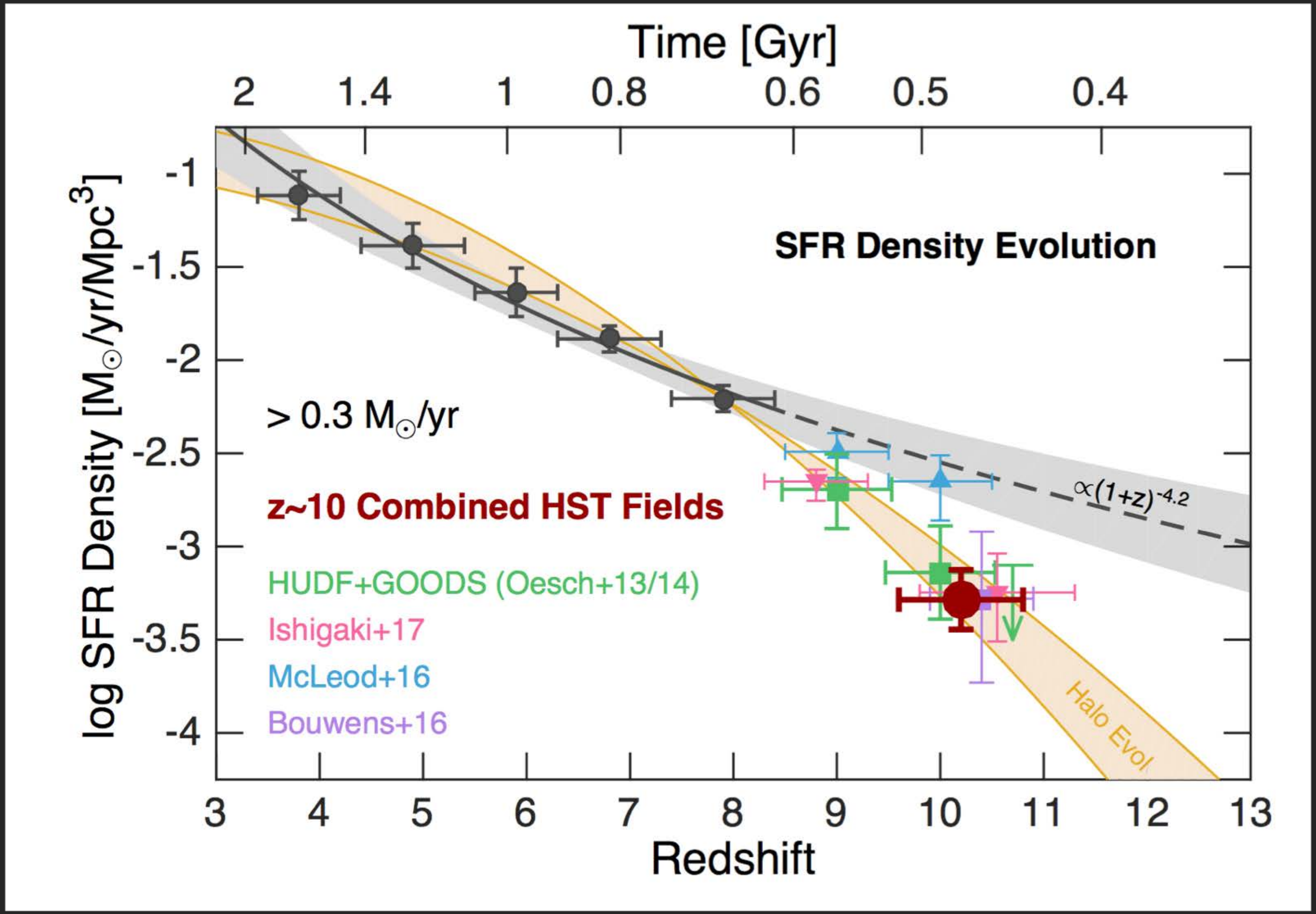
Oesch+2017

gdi

see also: Zheng+2012; Coe+2013; Bouwens+2013,15,16; Ellis+2013; McLure+2013; Ishigaki+2014,17; Infante+2015; Bernard+2016; Calvi+2016; McLeod+2016

“accelerated evolution” – the star formation rate density at z~9-10

clearly a trend to lower SFRD at z>8



Oesch+2013,2014,2017

gdi

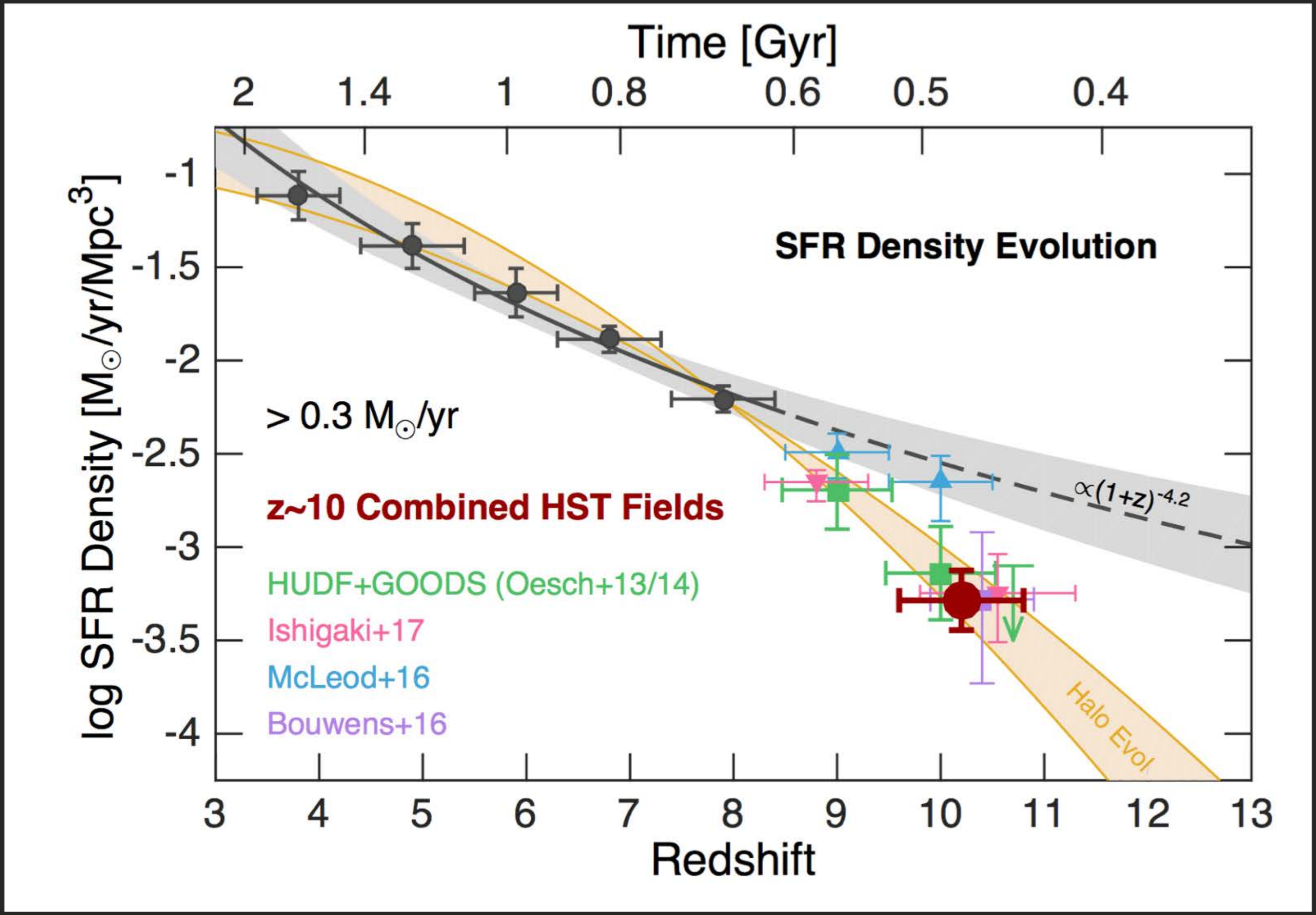
see also: Zheng+2012; Coe+2013; Bouwens+2013,15,16; Ellis+2013; McLure+2013; Ishigaki+2014,17; Infante+2015; Bernard+2016; Calvi+2016; McLeod+2016

“accelerated evolution” – the star formation rate density at z~9-10

clearly a trend to lower SFRD at z>8

“accelerated evolution” is actually consistent with the expected buildup* of dark matter halos over that time

*dark matter halo growth (>~10^{10} $M_\odot$) from HMFcalc – Murray+2013



Oesch+2017

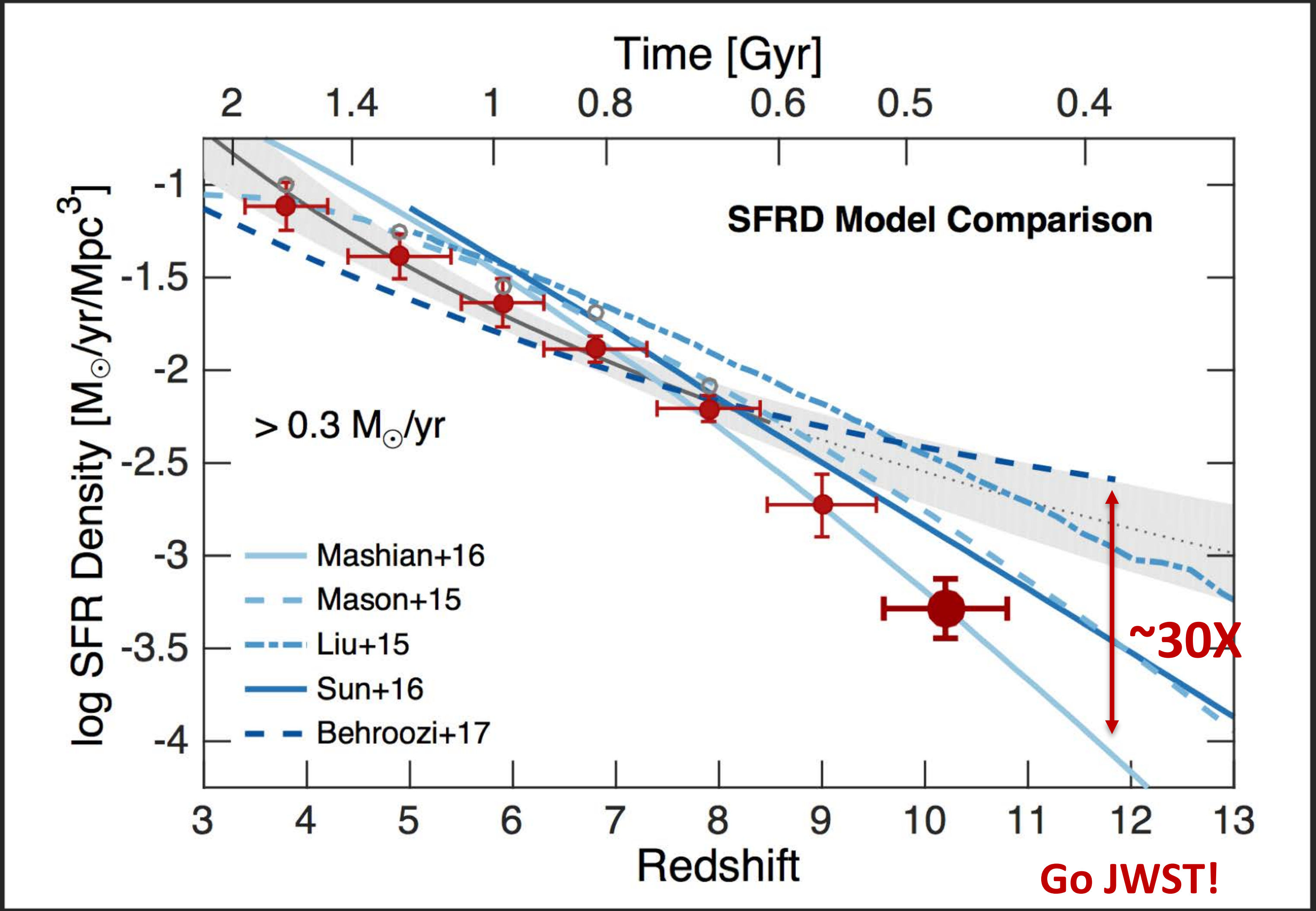
gdi

see also: Zheng+2012; Coe+2013; Bouwens+2013,15,16; Ellis+2013; McLure+2013; Ishigaki+2014,17; Infante+2015; Bernard+2016; Calvi+2016; McLeod+2016

model comparisons – the star formation rate density at z>6

note that "accelerated evolution" is seen in some models, but there is a large range of shapes/slopes

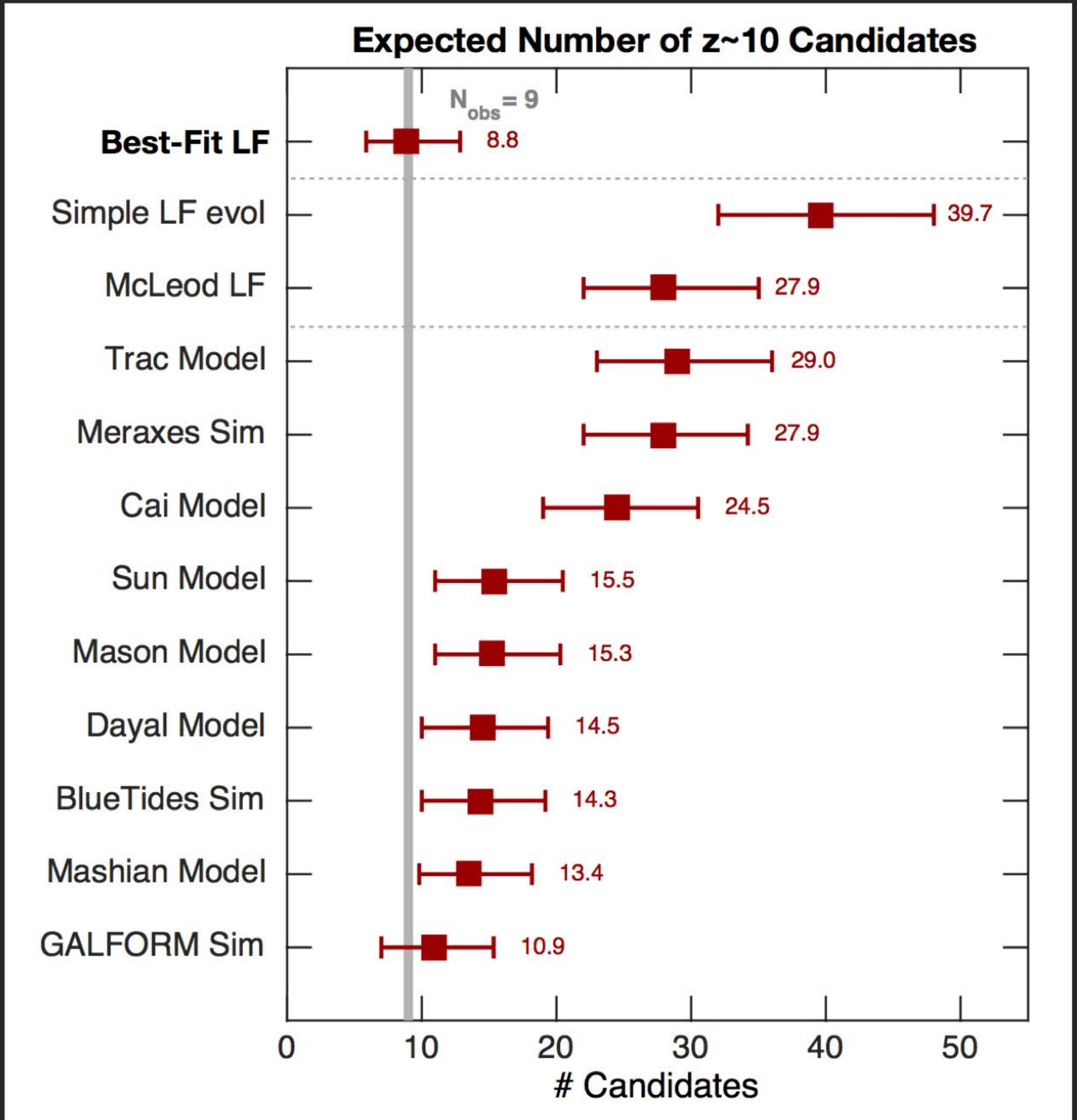
Oesch+2017



gdi

the case of the missing z~10 galaxies

number of z~10 galaxies from
“observed luminosity function”



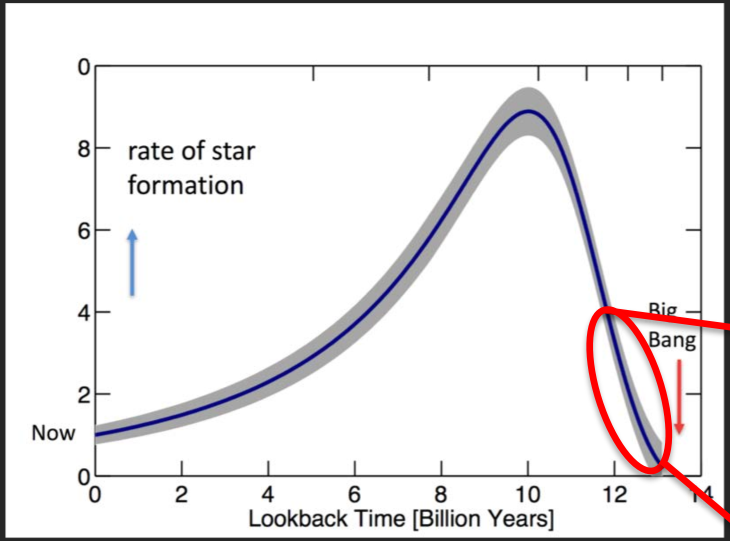
the situation at z~10 is unexpected

the numbers of objects is smaller than any model – the offsets are quite systematic

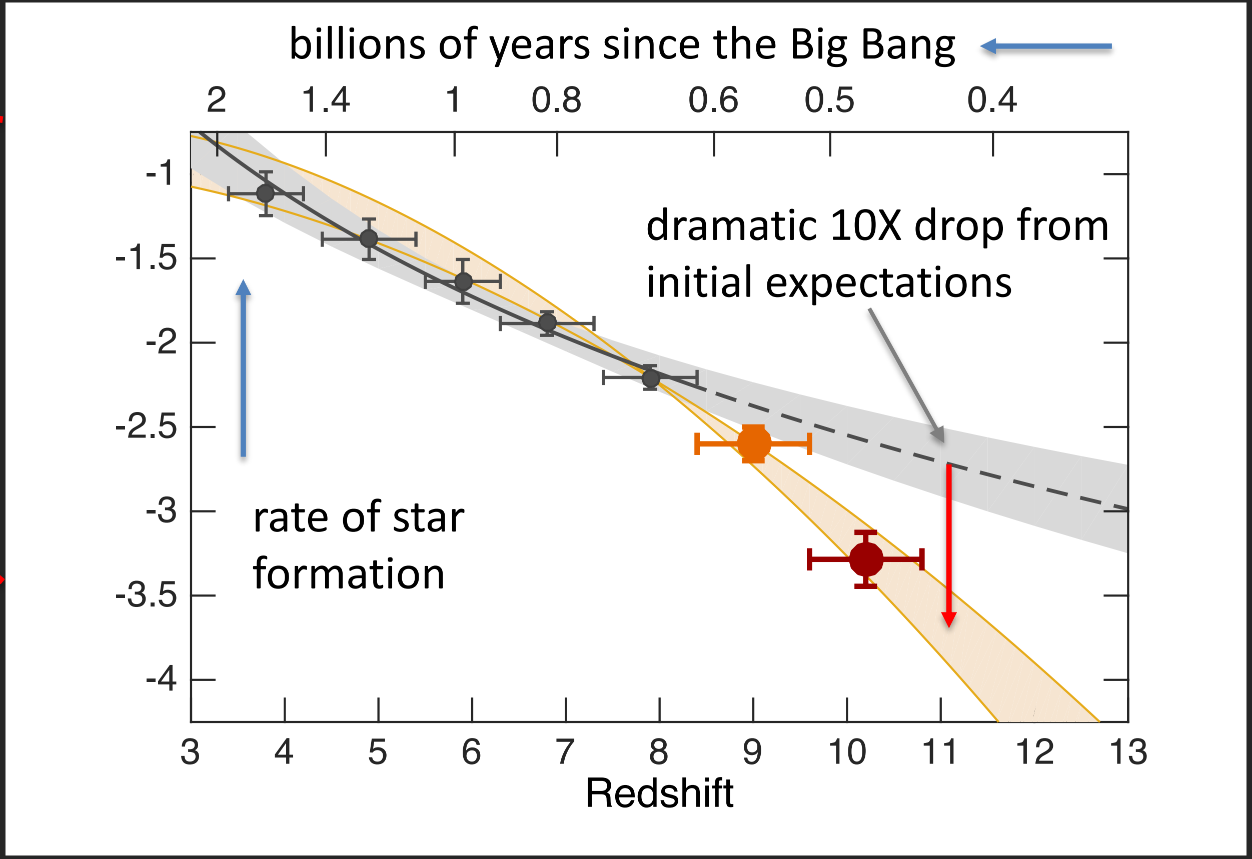
Oesch+2017

gdi

way fewer galaxies than expected at redshift 10



galaxies are evolving rapidly earlier than 650 million years



there are far fewer galaxies than we (naively) expected at early times

"accelerated evolution" is a very important result for JWST

Oesch+2017

z~9: Bouwens+2018

gdi

☛ *what does this mean for JWST and our search for the “first galaxies”?* ☚

gdi

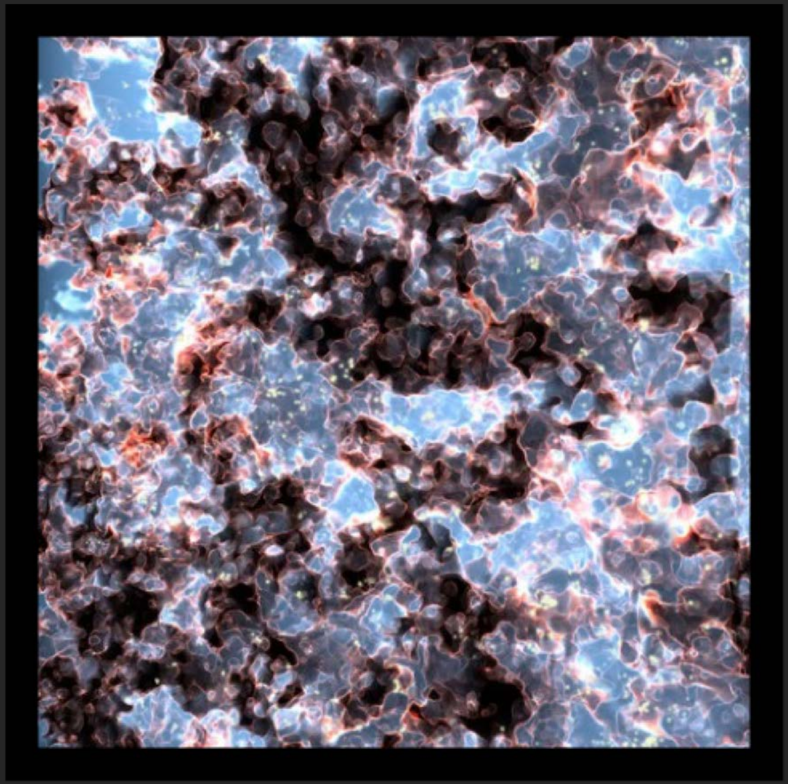
can JWST find the first galaxies?

will they be so rare that they will be hard to find?

will they occur at such high redshifts that they will be hard for JWST to see?

gdi

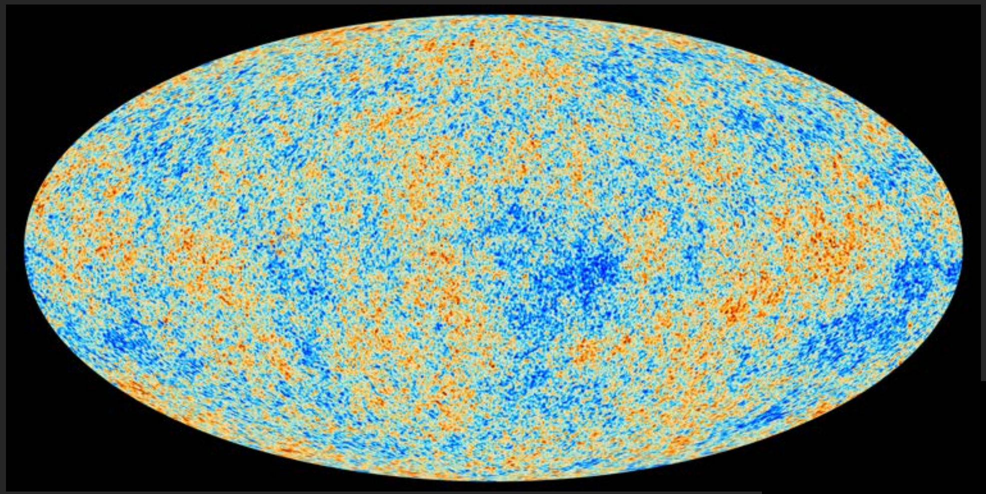
reionization epoch – latest Planck results



simulation: Alvarez et al. 2009

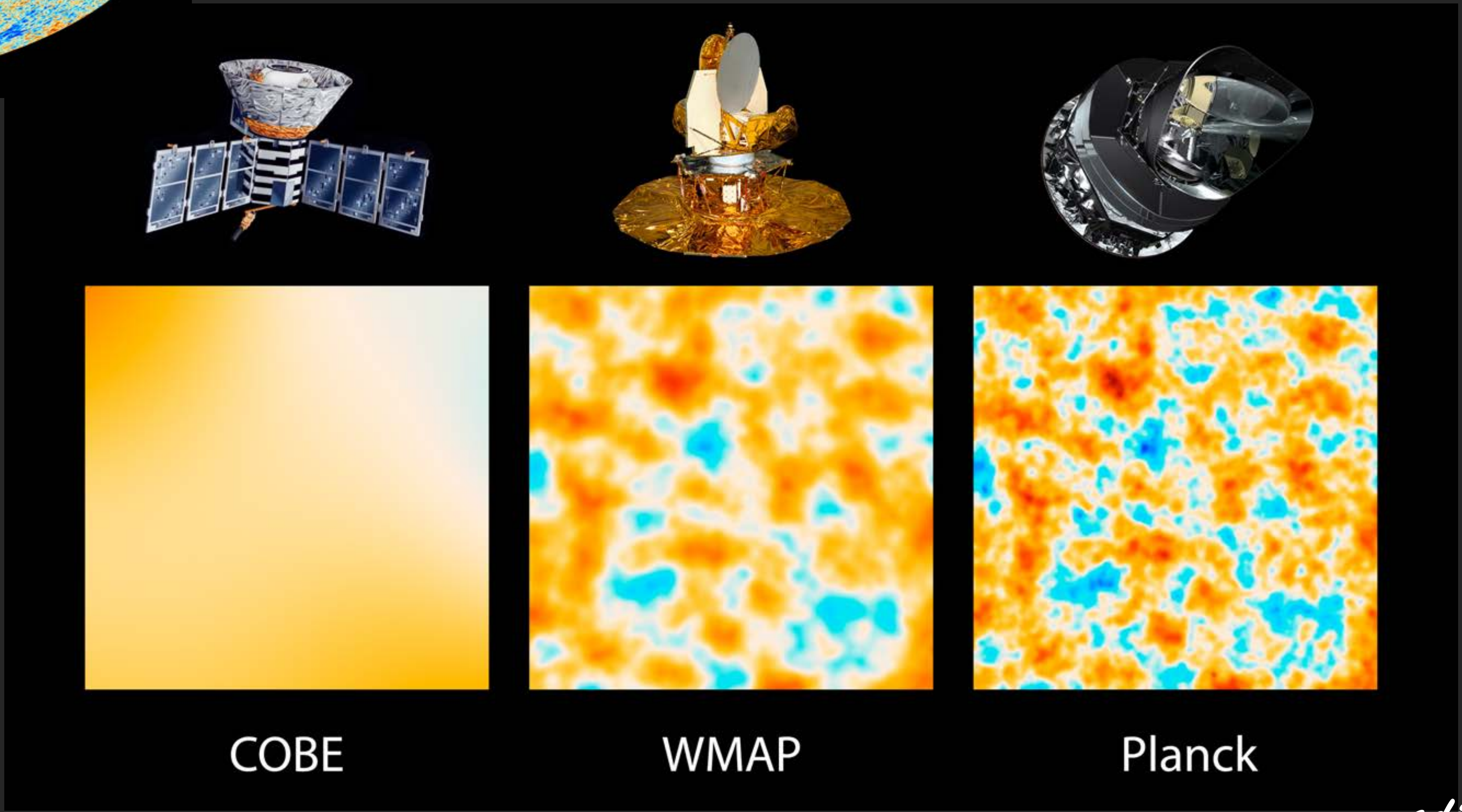
striking concordance between 2016
Planck results and galaxy constraints

measuring the fluctuations in the 3ºK microwave background across the whole sky



Planck all-sky map of the microwave 3ºK background

three amazing missions



1989 2001 2009

Planck 2016

constraints on the reionization history

remarkable mission

- …Thomson optical depth τ = 0.058±0.012….
- …average redshift at which reionization occurs is found to lie between *z* = 7.8 and 8.8…
- …upper limit to the width of the reionization period of Δ*z* < 2.8.
- …the Universe is ionized at less than the 10% level at redshifts above *z* ≃ 10…
- …an early onset of reionization is strongly disfavored by the *Planck* data.

Planck intermediate results

XLVII. *Planck* constraints on reionization history

Planck Collaboration: R. Adam[67], N. Aghanim[53], M. Ashdown[63,7], J. Aumont[53], C. Baccigalupi[75], M. Ballardini[29,45,48], A. J. Banday[85,10], R. B. Barreiro[58], N. Bartolo[28,59], S. Basak[75], R. Battye[61], K. Benabed[54,84], J.-P. Bernard[85,10], M. Bersanelli[32,46], P. Bielewicz[72,10,75], J. J. Bock[60,11], A. Bonaldi[61], L. Bonavera[16], J. R. Bond[9], J. Borrill[12,81], F. R. Bouchet[54,79], F. Boulanger[53], M. Bucher[1], C. Burigana[45,30,48], E. Calabrese[82], J.-F. Cardoso[66,1,54], J. Carron[21], H. C. Chiang[23,8], L. P. L. Colombo[19,60], C. Combet[67], B. Comis[67], F. Couchot[64], A. Coulais[65], B. P. Crill[60,11], A. Curto[58,7,63], F. Cuttaia[45], R. J. Davis[61], P. de Bernardis[31], A. de Rosa[45], G. de Zotti[42,75], J. Delabrouille[1], E. Di Valentino[54,79], C. Dickinson[61], J. M. Diego[58], O. Doré[60,11], M. Douspis[53], A. Ducout[54,52], X. Dupac[36], F. Elsner[20,54,84], T. A. Enßlin[70], H. K. Eriksen[56], E. Falgarone[65], Y. Fantaye[34,3], F. Finelli[45,48], F. Forastieri[30,49], M. Frailis[44], A. A. Fraisse[23], E. Franceschi[45], A. Frolov[78], S. Galeotta[44], S. Galli[62], K. Ganga[1], R. T. Génova-Santos[57,15], M. Gerbino[83,74,31], T. Ghosh[53], J. González-Nuevo[16,58], K. M. Górski[60,87], A. Gruppuso[45,48], J. E. Gudmundsson[83,74,23], F. K. Hansen[56], G. Helou[11], S. Henrot-Versillé[64], D. Herranz[58], E. Hivon[54,84], Z. Huang[9], S. Ilić[85,10,6], A. H. Jaffe[52], W. C. Jones[23], E. Keihänen[22], R. Keskitalo[12], T. S. Kisner[69], L. Knox[25], N. Krachmalnicoff[32], M. Kunz[14,53,3], H. Kurki-Suonio[22,41], G. Lagache[5,53], A. Lähteenmäki[2,41], J.-M. Lamarre[65], M. Langer[53], A. Lasenby[7,63], M. Lattanzi[30,49], C. R. Lawrence[60], M. Le Jeune[1], F. Levrier[65], A. Lewis[21], M. Liguori[28,59], P. B. Lilje[56], M. López-Caniego[36], Y.-Z. Ma[61,76], J. F. Macías-Pérez[67], G. Maggio[44], A. Mangilli[53,64], M. Maris[44], P. G. Martin[9], E. Martínez-González[58], S. Matarrese[28,59,38], N. Mauri[48], J. D. McEwen[71], P. R. Meinhold[26], A. Melchiorri[31,50], A. Mennella[32,46], M. Migliaccio[55,63], M.-A. Miville-Deschênes[53,9], D. Molinari[30,45,49], A. Moneti[54], L. Montier[85,10], G. Morgante[45], A. Moss[77], P. Naselsky[73,35], P. Natoli[30,4,49], C. A. Oxborrow[13], L. Pagano[31,50], D. Paoletti[45,48], B. Partridge[40], G. Patanchon[1], L. Patrizii[48], O. Perdereau[64], L. Perotto[67], V. Pettorino[39], F. Piacentini[31], S. Plaszczynski[64], L. Polastri[30,49], G. Polenta[4,43], J.-L. Puget[53], J. P. Rachen[17,70], B. Racine[56], M. Reinecke[70], M. Remazeilles[61,53,1], A. Renzi[34,51], G. Rocha[60,11], M. Rossetti[32,46], G. Roudier[1,65,60], J. A. Rubiño-Martín[57,15], B. Ruiz-Granados[86], L. Salvati[31], M. Sandri[45], M. Savelainen[22,41], D. Scott[18], G. Sirri[48], R. Sunyaev[70,80], A.-S. Suur-Uski[22,41], J. A. Tauber[37], M. Tenti[47], L. Toffolatti[16,58,45], M. Tomasi[32,46], M. Tristram[64,*], T. Trombetti[45,30], J. Valiviita[22,41], F. Van Tent[68], P. Vielva[58], F. Villa[45], N. Vittorio[33], B. D. Wandelt[54,84,27], I. K. Wehus[60,56], M. White[24], A. Zacchei[44], and A. Zonca[26]

Received 11 May 2016 / Accepted 23 July 2016

ABSTRACT

We investigate constraints on cosmic reionization extracted from the *Planck* cosmic microwave background (CMB) data. We combine the *Planck* CMB anisotropy data in temperature with the low-multipole polarization data to fit ΛCDM models with various parameterizations of the reionization history. We obtain a Thomson optical depth $\tau = 0.058 \pm 0.012$ for the commonly adopted instantaneous reionization model. This confirms, with data solely from CMB anisotropies, the low value suggested by combining *Planck* 2015 results with other data sets, and also reduces the uncertainties. We reconstruct the history of the ionization fraction using either a symmetric or an asymmetric model for the transition between the neutral and ionized phases. To determine better constraints on the duration of the reionization process, we also make use of measurements of the amplitude of the kinetic Sunyaev-Zeldovich (kSZ) effect using additional information from the high-resolution Atacama Cosmology Telescope and South Pole Telescope experiments. The average redshift at which reionization occurs is found to lie between $z = 7.8$ and 8.8, depending on the model of reionization adopted. Using kSZ constraints and a redshift-symmetric reionization model, we find an upper limit to the width of the reionization period of $\Delta z < 2.8$. In all cases, we find that the Universe is ionized at less than the 10% level at redshifts above $z \simeq 10$. This suggests that an early onset of reionization is strongly disfavoured by the *Planck* data. We show that this result also reduces the tension between CMB-based analyses and constraints from other astrophysical sources.

Key words. cosmic background radiation – dark ages, reionization, first stars – polarization

Plank Collaboration XLVII + 2016

gdi

constraints on the reionization history

Planck 2016

remarkable mission

- …Thomson optical depth τ = 0.058±0.012….
- …average redshift at which reionization occurs is found to lie between *z* = 7.8 and 8.8…
- …upper limit to the width of the reionization period of Δ*z* < 2.8.
- …the Universe is ionized at less than the 10% level at redshifts above *z* ≃ 10…
- …an early onset of reionization is strongly disfavored by the *Planck* data.

Planck intermediate results

XLVII. _Planck_ constraints on reionization history

Planck Collaboration: R. Adam[67], N. Aghanim[53], M. Ashdown[63,7], J. Aumont[53], C. Baccigalupi[75], M. Ballardini[29,45,48], A. J. Banday[85,10], R. B. Barreiro[58], N. Bartolo[28,59], S. Basak[75], R. Battye[61], K. Benabed[54,84], J.-P. Bernard[85,10], M. Bersanelli[32,46], P. Bielewicz[72,10,75], J. J. Bock[60,11], A. Bonaldi[61], L. Bonavera[16], J. R. Bond[9], J. Borrill[12,81], F. R. Bouchet[54,79], F. Boulanger[53], M. Bucher[1], C. Burigana[45,30,48], E. Calabrese[82], J.-F. Cardoso[66,1,54], J. Carron[21], H. C. Chiang[23,8], L. P. L. Colombo[19,60], C. Combet[67], B. Comis[67], F. Couchot[64], A. Coulais[65], B. P. Crill[60,11], A. Curto[58,7,63], F. Cuttaia[45], R. J. Davis[61], P. de Bernardis[31], A. de Rosa[45], G. de Zotti[42,75], J. Delabrouille[1], E. Di Valentino[54,79], C. Dickinson[61], J. M. Diego[58], O. Doré[60,11], M. Douspis[53], A. Ducout[54,52], X. Dupac[36], F. Elsner[20,54,84], T. A. Enßlin[70], H. K. Eriksen[56], E. Falgarone[65], Y. Fantaye[34,3], F. Finelli[45,48], F. Forastieri[30,49], M. Frailis[44], A. A. Fraisse[23], E. Franceschi[45], A. Frolov[78], S. Galeotta[44], S. Galli[62], K. Ganga[1], R. T. Génova-Santos[57,15], M. Gerbino[83,74,31], T. Ghosh[53], J. González-Nuevo[16,58], K. M. Górski[60,87], A. Gruppuso[45,48], J. E. Gudmundsson[83,74,23], F. K. Hansen[56], G. Helou[11], S. Henrot-Versillé[64], D. Herranz[58], E. Hivon[54,84], Z. Huang[9], S. Ilić[85,10,6], A. H. Jaffe[52], W. C. Jones[23], E. Keihänen[22], R. Keskitalo[12], T. S. Kisner[69], L. Knox[25], N. Krachmalnicoff[32], M. Kunz[14,53,3], H. Kurki-Suonio[22,41], G. Lagache[5,53], A. Lähteenmäki[2,41], J.-M. Lamarre[65], M. Langer[53], A. Lasenby[7,63], M. Lattanzi[30,49], C. R. Lawrence[60], M. Le Jeune[1], F. Levrier[65], A. Lewis[21], M. Liguori[28,59], P. B. Lilje[56], M. López-Caniego[36], Y.-Z. Ma[61,76], J. F. Macías-Pérez[67], G. Maggio[44], A. Mangilli[53,64], M. Maris[44], P. G. Martin[9], E. Martínez-González[58], S. Matarrese[28,59,38], N. Mauri[48], J. D. McEwen[71], P. R. Meinhold[26], A. Melchiorri[31,50], A. Mennella[32,46], M. Migliaccio[55,63], M.-A. Miville-Deschênes[53,9], D. Molinari[30,45,49], A. Moneti[54], L. Montier[85,10], G. Morgante[45], A. Moss[77], P. Naselsky[73,35], P. Natoli[30,4,49], C. A. Oxborrow[13], L. Pagano[31,50], D. Paoletti[45,48], B. Partridge[40], G. Patanchon[1], L. Patrizii[48], O. Perdereau[64], L. Perotto[67], V. Pettorino[39], F. Piacentini[31], S. Plaszczynski[64], L. Polastri[30,49], G. Polenta[4,43], J.-L. Puget[53], J. P. Rachen[17,70], B. Racine[56], M. Reinecke[70], M. Remazeilles[61,53,1], A. Renzi[34,51], G. Rocha[60,11], M. Rossetti[32,46], G. Roudier[1,65,60], J. A. Rubiño-Martín[57,15], B. Ruiz-Granados[86], L. Salvati[31], M. Sandri[45], M. Savelainen[22,41], D. Scott[18], G. Sirri[48], R. Sunyaev[70,80], A.-S. Suur-Uski[22,41], J. A. Tauber[37], M. Tenti[47], L. Toffolatti[16,58,45], M. Tomasi[32,46], M. Tristram[64,⋆], T. Trombetti[45,30], J. Valiviita[22,41], F. Van Tent[68], P. Vielva[58], F. Villa[45], N. Vittorio[33], B. D. Wandelt[54,84,27], I. K. Wehus[60,56], M. White[24], A. Zacchei[44], and A. Zonca[26]

Received 11 May 2016 / Accepted 23 July 2016

ABSTRACT

We investigate constraints on cosmic reionization extracted from the *Planck* cosmic microwave background (CMB) data. We combine the *Planck* CMB anisotropy data in temperature with the low-multipole polarization data to fit ΛCDM models with various parameterizations of the reionization history. We obtain a Thomson optical depth $\tau = 0.058 \pm 0.012$ for the commonly adopted instantaneous reionization model. This confirms, with data solely from CMB anisotropies, the low value suggested by combining *Planck* 2015 results with other data sets, and also reduces the uncertainties. We reconstruct the history of the ionization fraction using either a symmetric or an asymmetric model for the transition between the neutral and ionized phases. To determine better constraints on the duration of the reionization process, we also make use of measurements of the amplitude of the kinetic Sunyaev-Zeldovich (kSZ) effect using additional information from the high-resolution Atacama Cosmology Telescope and South Pole Telescope experiments. The average redshift at which reionization occurs is found to lie between $z = 7.8$ and 8.8, depending on the model of reionization adopted. Using kSZ constraints and a redshift-symmetric reionization model, we find an upper limit to the width of the reionization period of $\Delta z < 2.8$. In all cases, we find that the Universe is ionized at less than the 10% level at redshifts above $z \simeq 10$. This suggests that an early onset of reionization is strongly disfavoured by the *Planck* data. We show that this result also reduces the tension between CMB-based analyses and constraints from other astrophysical sources.

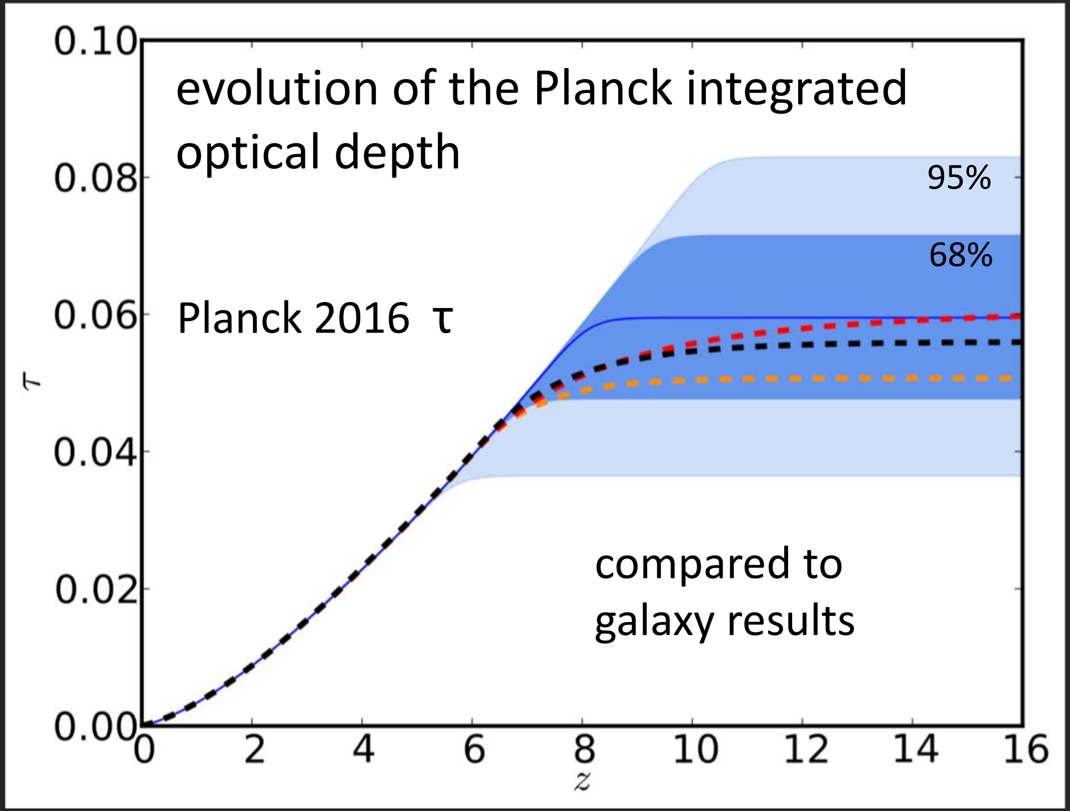
Key words. cosmic background radiation – dark ages, reionization, first stars – polarization

Plank Collaboration XLVII + 2016

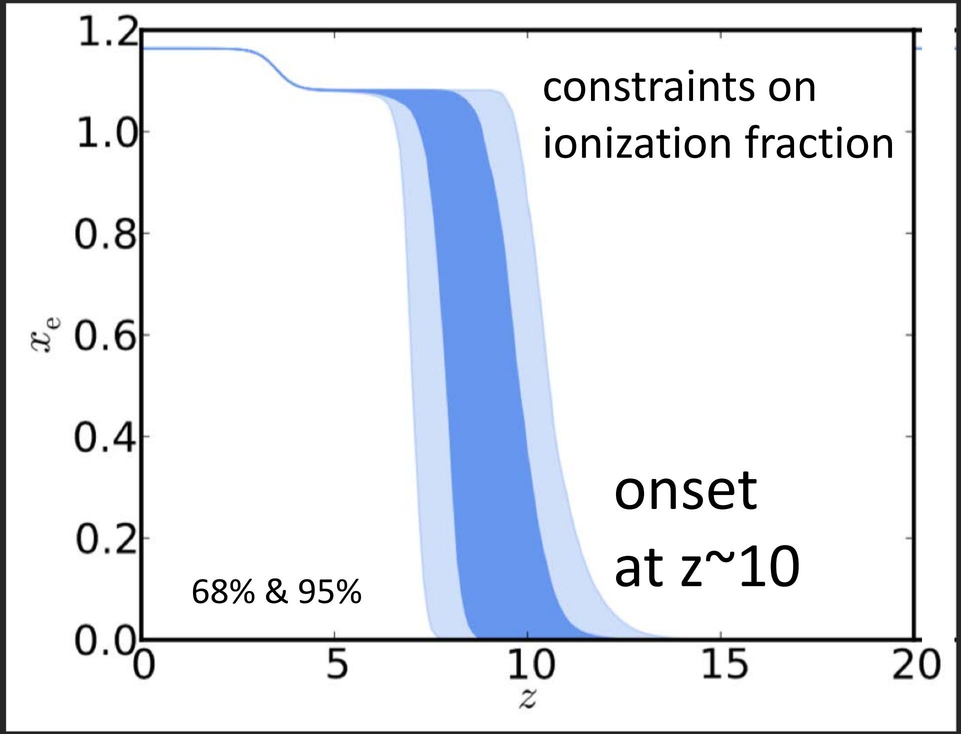
gdi

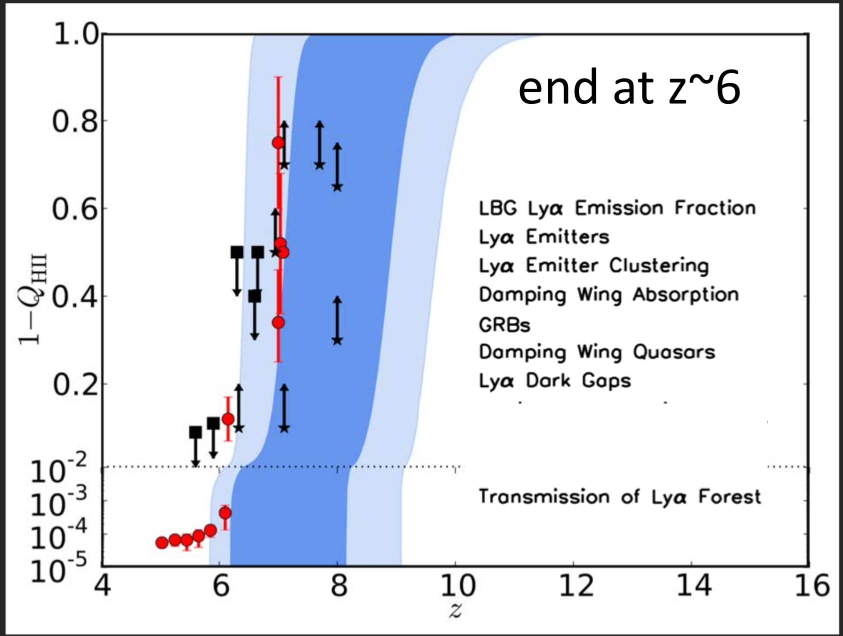
reionization constraints from Planck 2016

striking consistency with galaxy results



Bouwens+2015
Robertson+2015
Ishigaki+2015





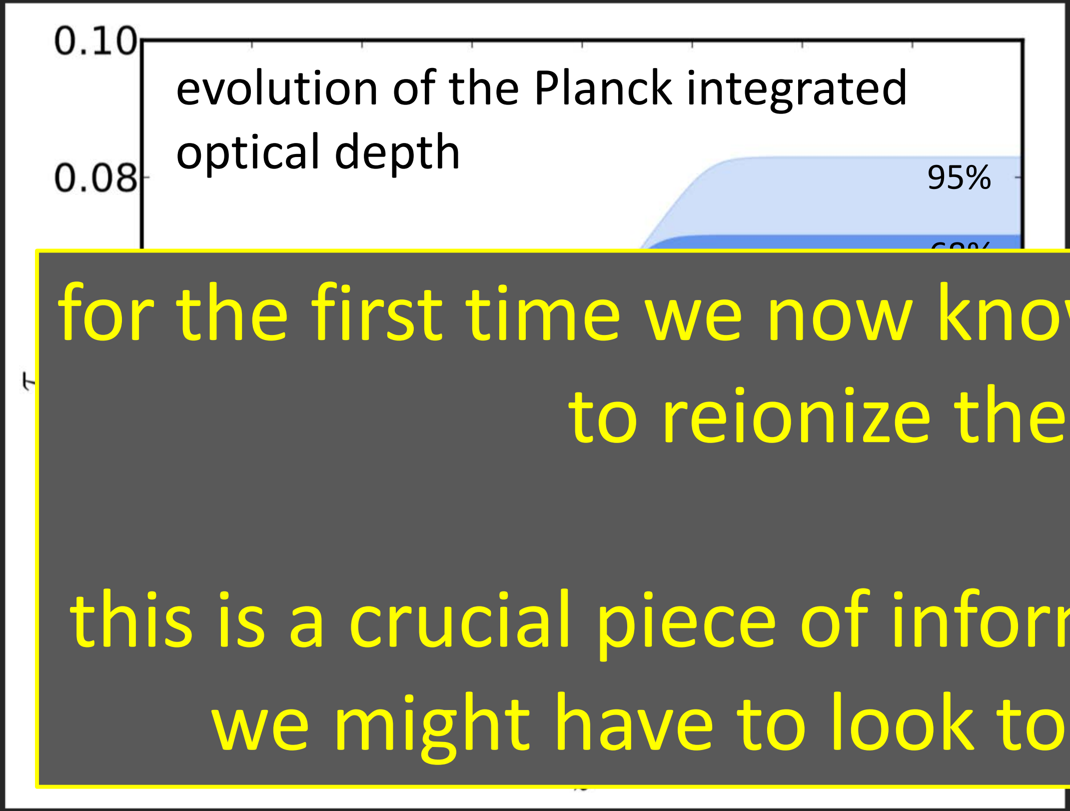
reionization history compared with observational constraints

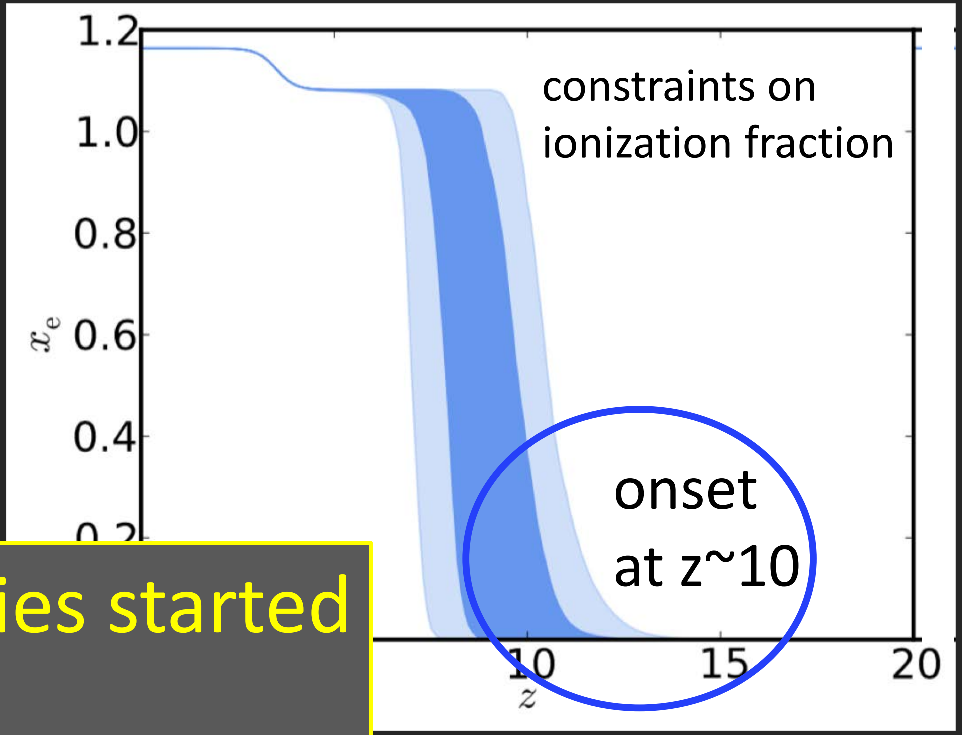
Plank Collaboration XLVII + 2016

gdi

reionization constraints from Planck 2016

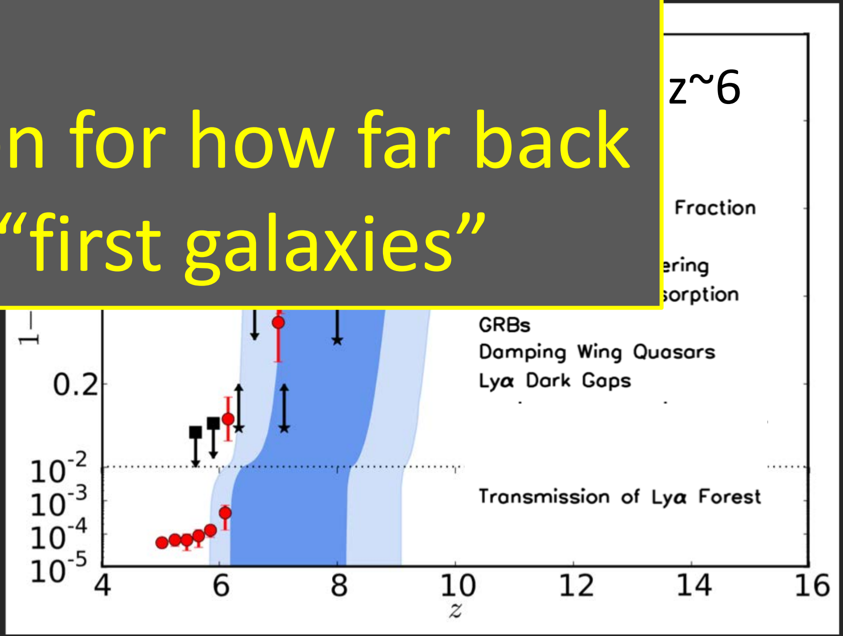
striking consistency with galaxy results





for the first time we now know when galaxies started to reionize the universe

this is a crucial piece of information for how far back we might have to look to find "first galaxies"



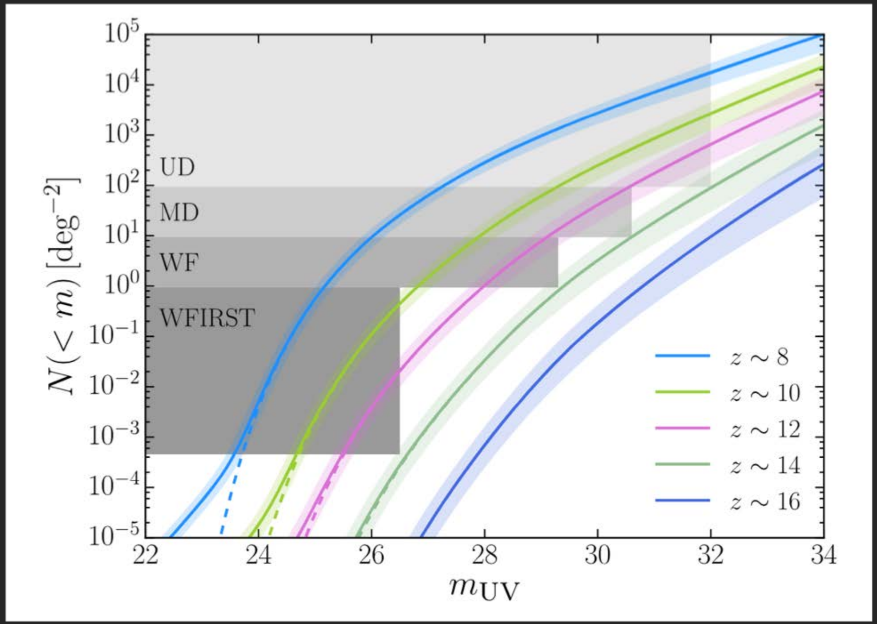
reionization history compared with observational constraints

Plank Collaboration XLVII + 2016

gdi

"First Light and Reionization"
one of JWST's four science themes

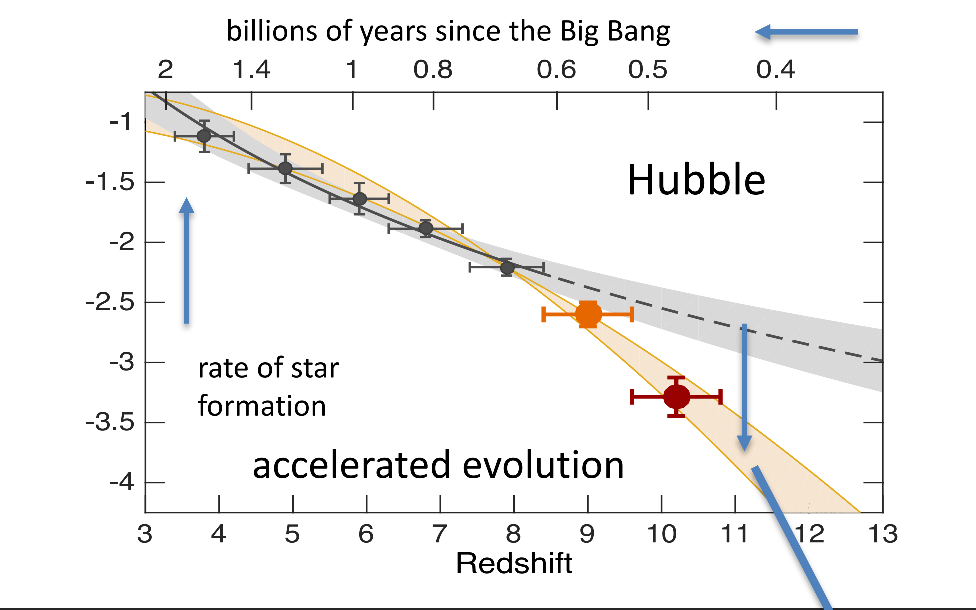
can JWST see the "first galaxies"?



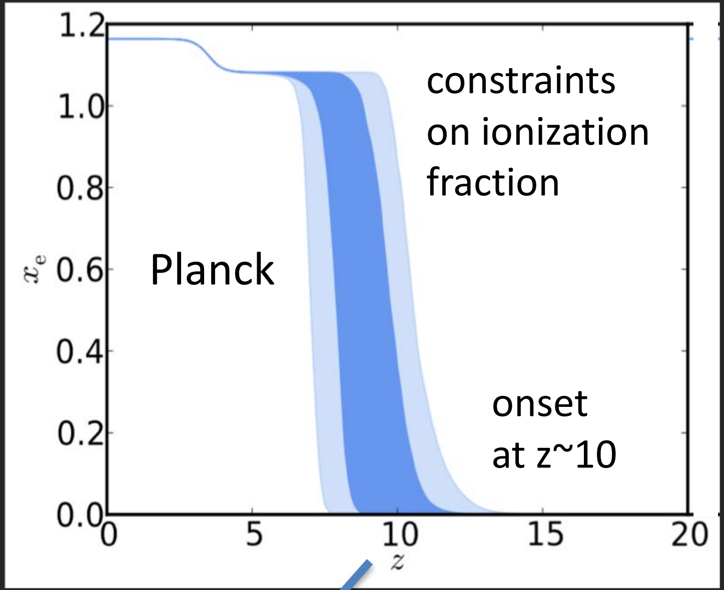
recent studies suggest that JWST will reach to redshift z~ 14-15 in the deepest studies.

Mason+2015

gdi



*can JWST see the "first galaxies"?***



**depends on what one means by "first galaxies" but yes…

large 10X drop from expected at z~11 + galaxy turn-on at z~10-11

☛ suggest major changes in galaxy population at z~10-12 ☚

great for JWST's "first light" goal since galaxies are turning on at z~10-12

likely major changes from z~10-15 – where JWST can see them

☞ exciting times ahead at "Cosmic Sunrise"! ☜

gdi

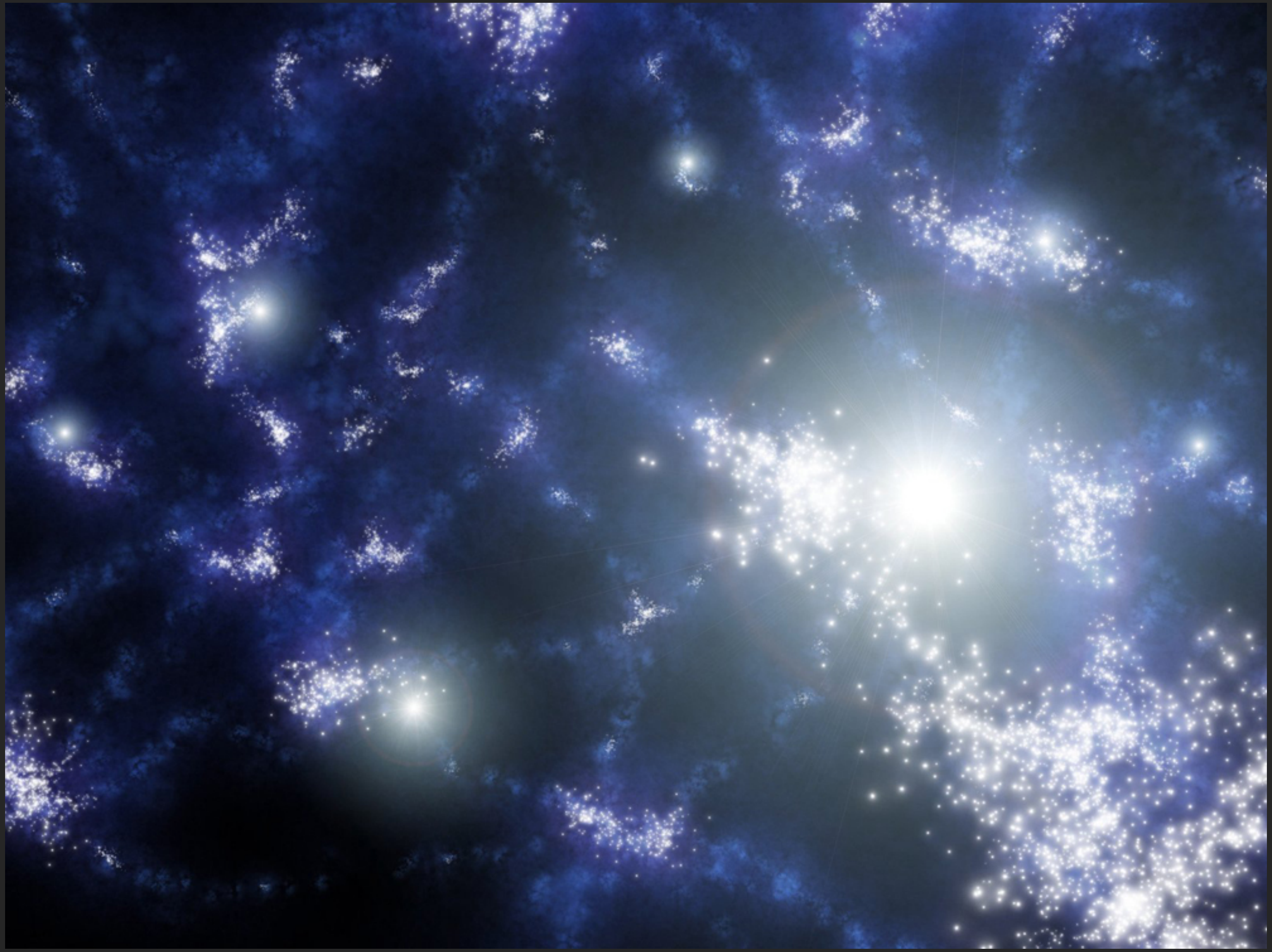
the dramatic brightening after dawn

desert sunrise

gdi

the dramatic brightening after cosmic dawn

"Cosmic Sunrise" as the first galaxies burst forth at z~12-15

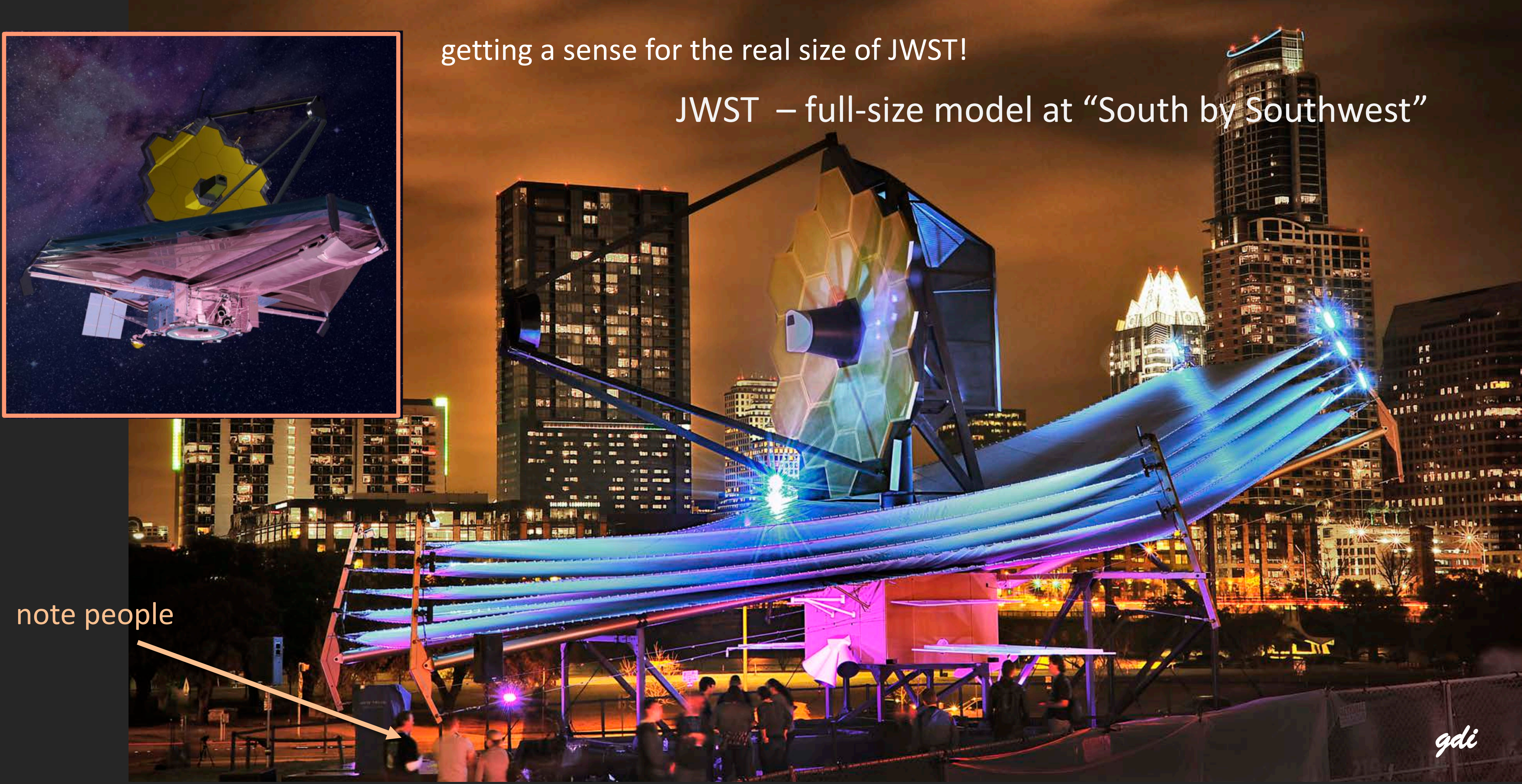


gdi

JWST is the "what's next" for the earliest galaxies

getting a sense for the real size of JWST!

JWST – full-size model at "South by Southwest"

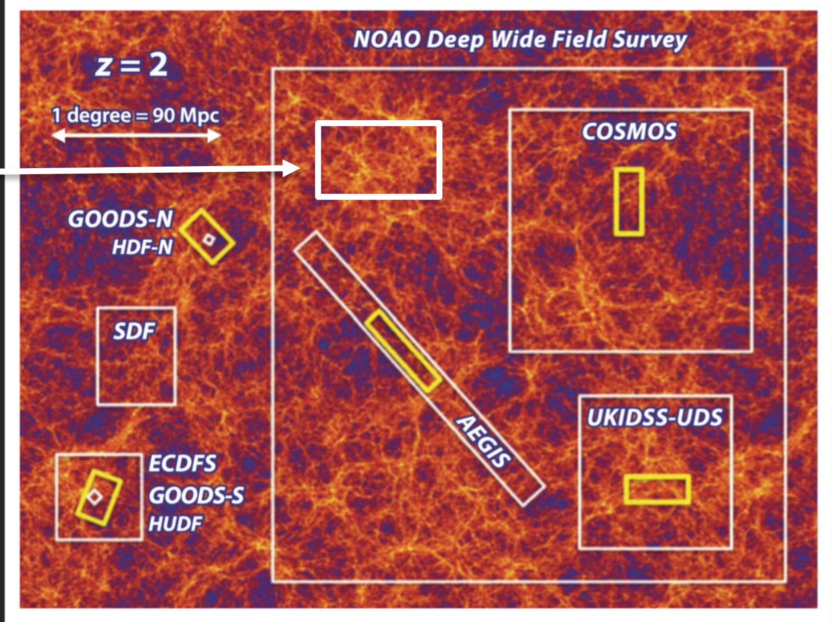


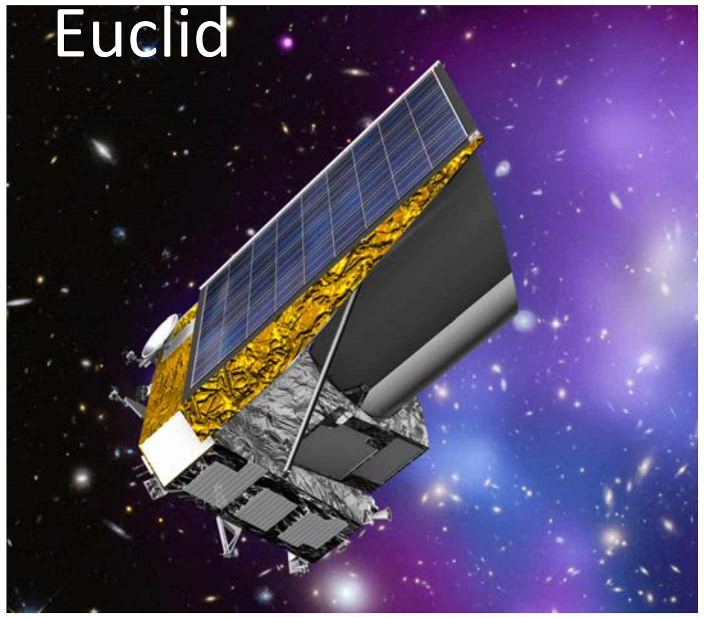
note people

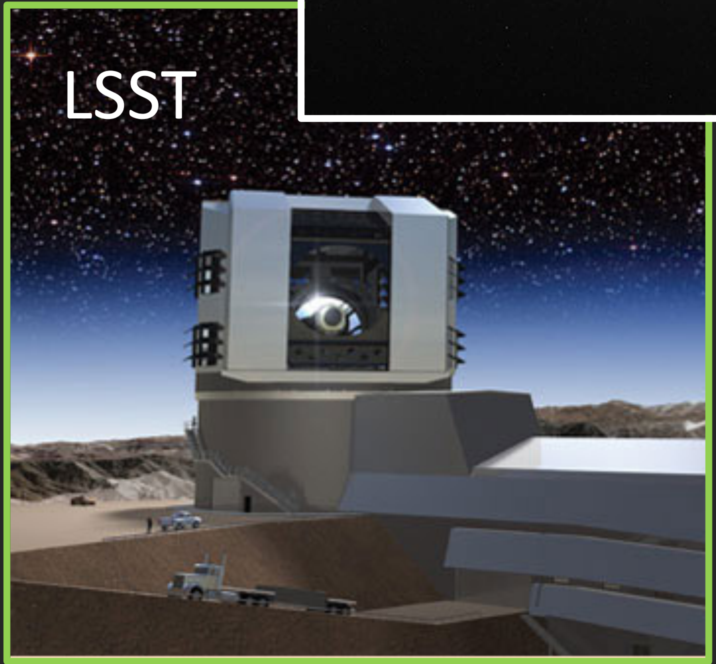
surveys to minimize cosmic variance

wide-area imaging >>Hubble or JWST

WFIRST →







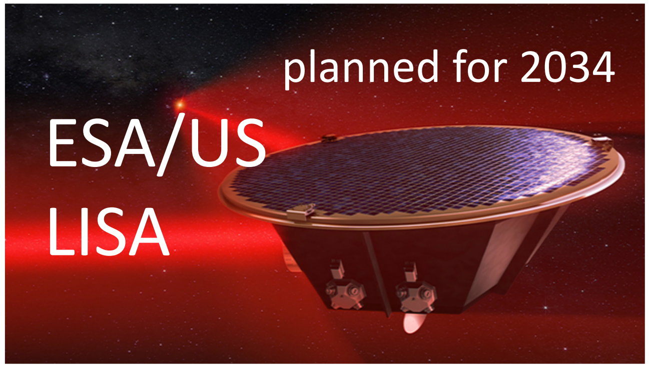
☛ *the long-term future – after JWST* ☚

great opportunities, but great challenges…..

gdi

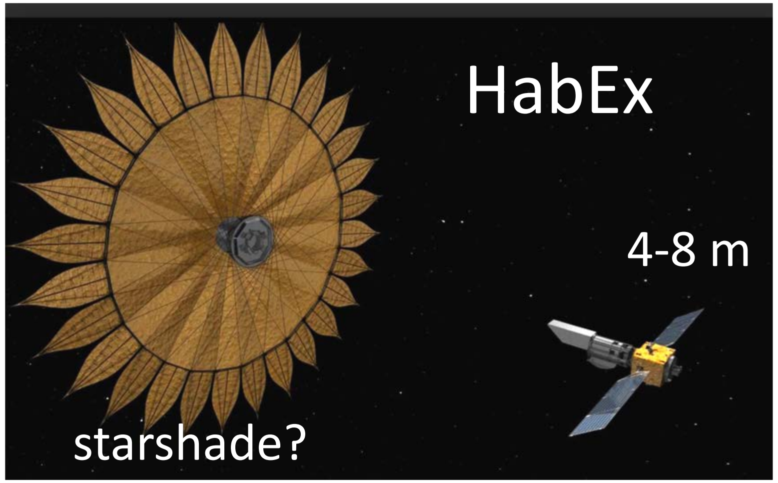
☛ *the flagships of the 2030s (?)* ☚

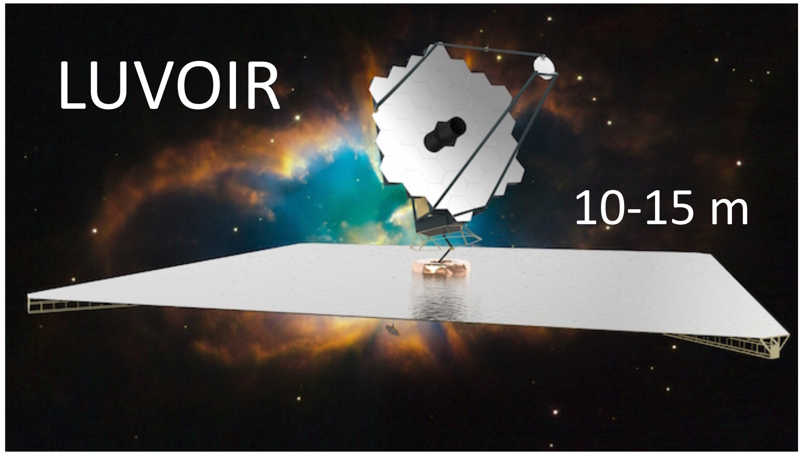
✓



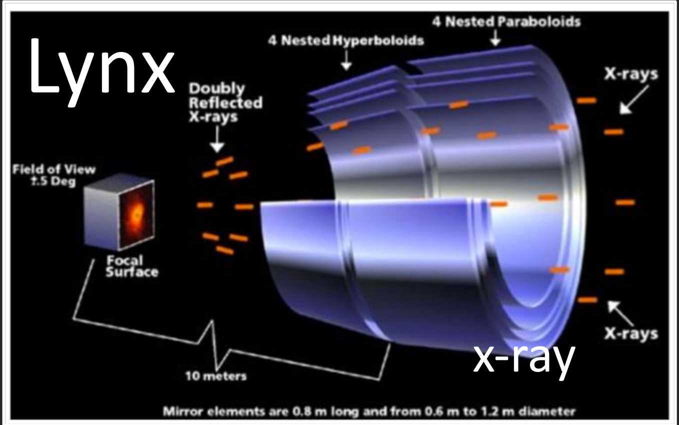
✓

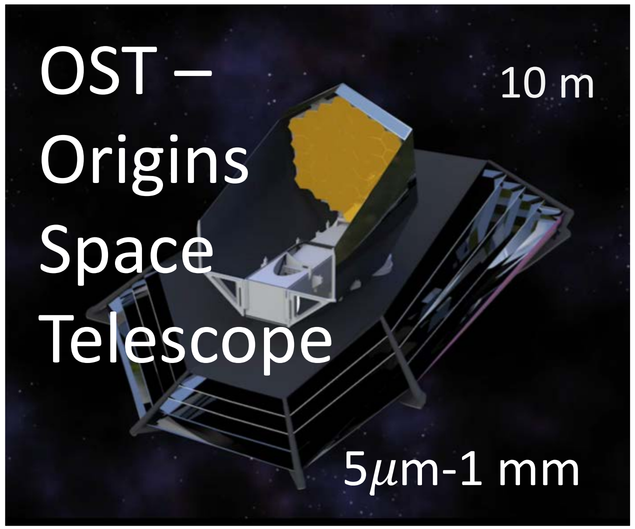
NASA strategic missions under study for the 2020 Decadal





?

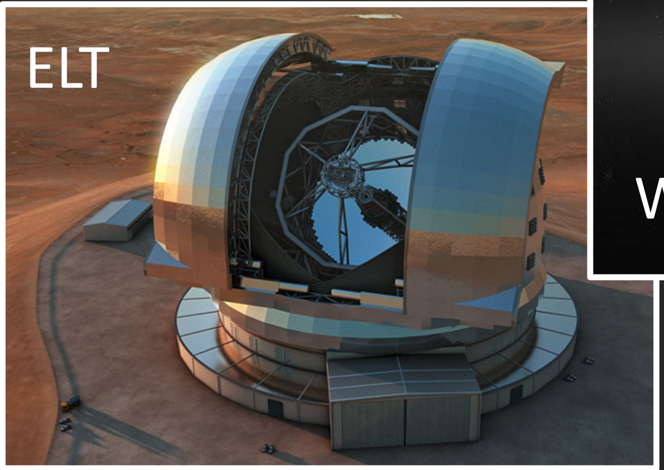


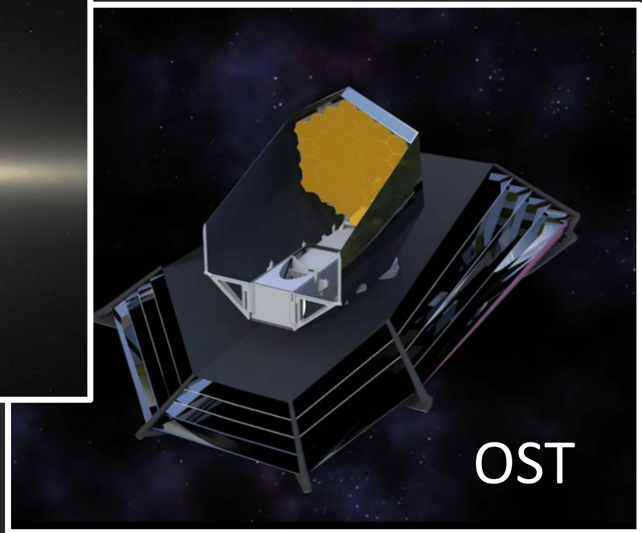


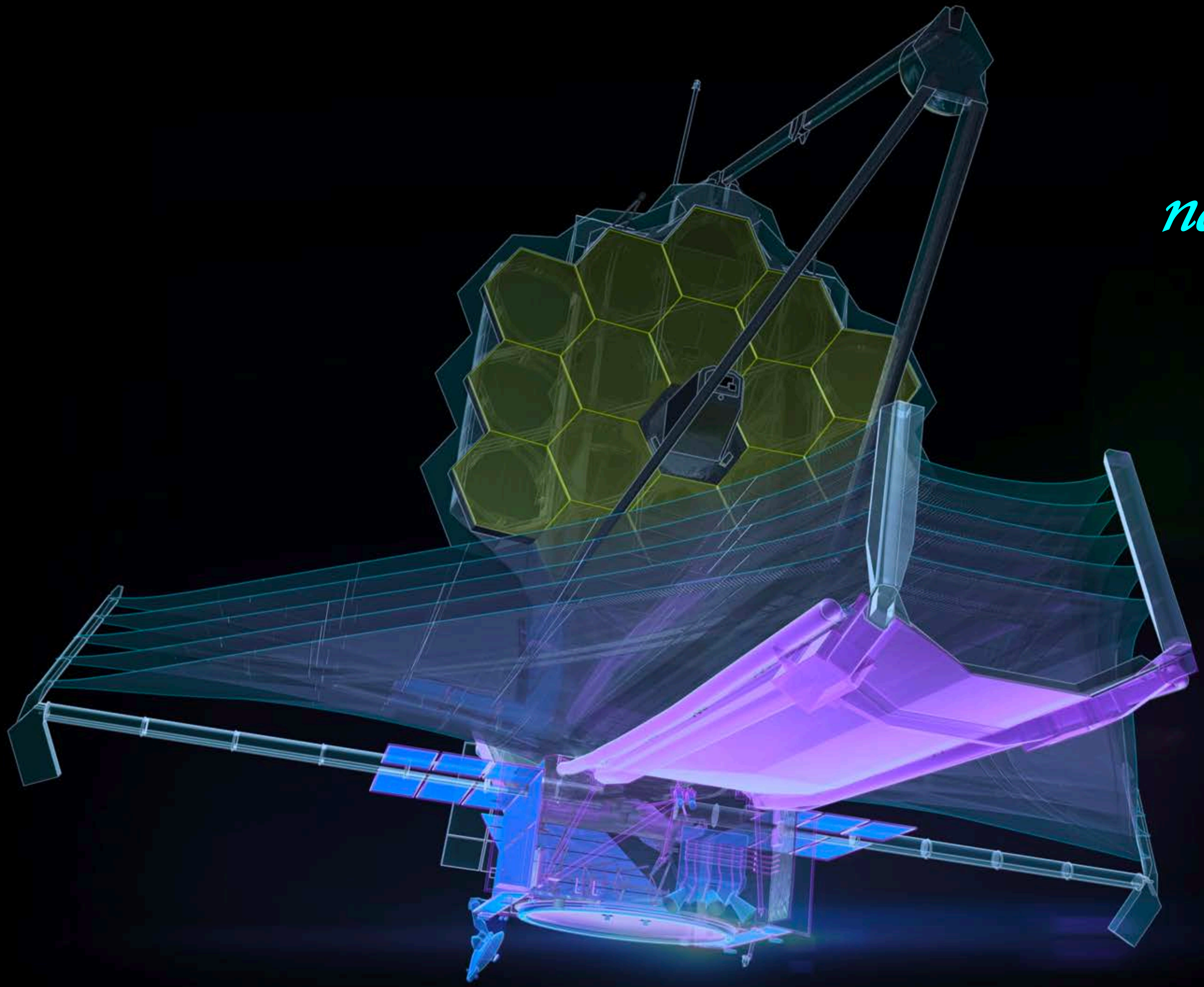
only 0-1 of these!

gdi

☛ *JWST, along with WFIRST (and similar telescopes) and the ELT, will transform our understanding of distant galaxies in the next decade, but, for distant galaxies, another "next generation telescope" will be needed in the decade beyond* ☚



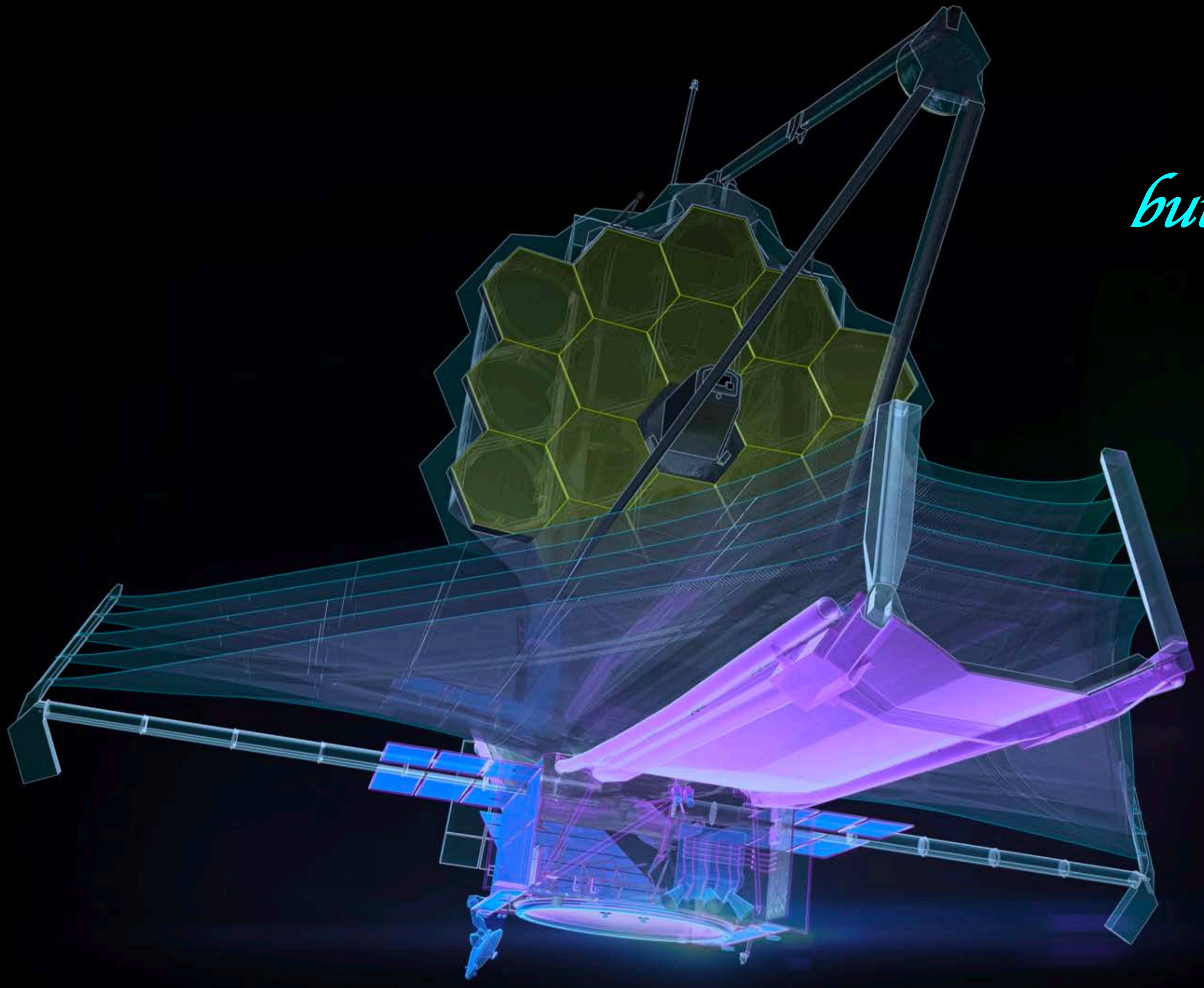




not the end...

gdi

but the beginning...



gdi